

A HOT MOLECULAR OUTFLOW DRIVEN BY THE IONIZED JET ASSOCIATED WITH IRAS 16562–3959

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ABSTRACT

We report molecular line observations in the CO $J = 3 \rightarrow 2$, $6 \rightarrow 5$, and $7 \rightarrow 6$ transitions, made using the Atacama Pathfinder Experiment Telescope, toward the massive and dense core IRAS 16562–3959. This core harbors a string of radio sources thought to be powered by a central collimated jet of ionized gas. The molecular observations show the presence of high-velocity gas exhibiting a quadrupolar morphology, most likely produced by the presence of two collimated outflows. The southeast–northwest (SE–NW) molecular outflow is aligned with the string of radio continuum sources, suggesting it is driven by the jet. We find that the excitation temperature of the gas in the SE–NW outflow is high, with values of 145 and 120 K for the blueshifted and redshifted lobes, respectively. This outflow has a total mass of $1.92 M_{\odot}$, a total momentum of $\sim 89 M_{\odot} \text{ km s}^{-1}$, and an averaged momentum rate of $\sim 3.0 \times 10^{-2} M_{\odot} \text{ km s}^{-1} \text{ yr}^{-1}$, values characteristic of flows driven by young massive stellar objects with high luminosities ($L_{\text{bol}} \sim 2 \times 10^4 L_{\odot}$). Complementary data taken with the Atacama Submillimeter Telescope Experiment in high density and shock tracers support the picture that IRAS 16562–3959 is an accreting young massive star associated with an ionized jet, which is the energy source of a molecular outflow.

Key words: ISM: individual objects (IRAS 16562–3959) – ISM: jets and outflows – radio continuum: stars – stars: formation

Online-only material: color figure

1. INTRODUCTION

IRAS 16562–3959 (G345.4938+01.4677) is a mid-infrared (MIR) source, with *Midcourse Space Experiment* (MSX) color characteristics of massive young stellar objects (MYSOs; Lumsden et al. 2002). Observations of dust continuum emission at 1.2 mm, made with SIMBA at the SEST telescope, show that IRAS 16562–3959 is associated with a massive ($1.3 \times 10^3 M_{\odot}$) and dense ($9 \times 10^5 \text{ cm}^{-3}$) molecular core (López et al. 2011). Maser emission has been detected in transitions of OH (Caswell 1998, 2004), giving further support to the hypothesis that IRAS 16562–3959 is a young high-mass star-forming region.

Recently, Guzmán et al. (2010) reported the detection toward IRAS 16562–3959 of a string of radio continuum emission consisting of a bright compact central object and four outer lobes. They argue that the continuum emission from the central object corresponds to free–free emission from a thermal jet, whereas the radio emission from the lobes correspond to thermal emission arising in shocks generated from the interaction of the collimated wind with the surrounding medium. Assuming that IRAS 16562–3959 is located at a distance of $1.6 \pm 0.1 \text{ kpc}$ ($V_{\text{LSR}} = -12.6 \text{ km s}^{-1}$; Urquhart et al. 2008),⁴ the total far-infrared luminosity, computed using the observed IRAS fluxes (see Casoli et al. 1986), is $\sim 7.0 \times 10^4 L_{\odot}$. This implies that IRAS 16562–3959 harbors the most luminous MYSO known to date associated with an ionized jet.

MIR images taken from the *Spitzer Space Telescope*-GLIMPSE survey data (Benjamin et al. 2003) show bright emission associated with the jet in all IRAC bands. In particular, at $8.0 \mu\text{m}$ there is extended emission roughly aligned with the jet axis. Further, there appears to be an excess of emission

at $4.5 \mu\text{m}$, which is thought to be indicative of shock activity (Cyganowski et al. 2008; Chambers et al. 2009). The GLIMPSE data also show that there is at least one more young stellar object (YSO) within the core, located about $15''$ northeast of the jet, that is not detected at $2 \mu\text{m}$ and hence likely to be deeply embedded.

In this paper, we present molecular line observations toward IRAS 16562–3959 in the CO ($3 \rightarrow 2$), CO ($6 \rightarrow 5$), and CO ($7 \rightarrow 6$) transitions, made using the Atacama Pathfinder Experiment Telescope (APEX), which have revealed the presence of two high-velocity bipolar molecular outflows. Their observed and derived characteristics suggest that the driving sources are young, luminous protostellar objects. One of the outflows is aligned in the same direction as the string of radio sources strongly suggesting that it is being driven by the thermal radio jet.

2. OBSERVATIONS

The observations were made using the 12 m APEX located at Llano de Chajnantor, Chile. A detailed description of APEX and its performance are given by Güsten et al. (2006). The observed transitions and basic observational parameters are summarized in Table 1. Columns 1 and 2 give, respectively, the observed transition and line frequency. Columns 3, 4, and 5 give the telescope used, the half-power beam width, and the main beam efficiency at the observed frequency. Columns 6–9 give, respectively, the number of positions observed, the angular spacing, the channel width, and the resulting rms noise in antenna temperature, for each of the observed transitions.

The CO ($3 \rightarrow 2$) observations were made during 2008 June and 2009 July. The front-end consisted of a single pixel heterodyne SiS receiver operating in the 275–370 GHz band (Risacher et al. 2006). For the back-end we used the APEX fast Fourier transform (FFT) spectrometer (Klein et al. 2006)

⁴ The near/far kinematic distance ambiguity was resolved in Faúndez et al. (2004).

Table 1
Observational Parameters

Line	Frequency (GHz)	Tel.	Beam (")	η_{mb}	Map	Spacing (")	Δv (km s ⁻¹)	Noise (K)
CO(3 → 2)	345.796	APEX	17.3	0.73	5 × 5	20"	0.423	0.1
CO(6 → 5)	691.627	APEX	8.8	0.47	29 × 21	4"	0.318	0.5
CO(7 → 6)	806.898	APEX	7.7	0.43	27 × 19	4"	0.272	1.0
SiO(8 → 7)	347.331	APEX	17.	0.73	3 × 3	20"	0.423	0.05
HCO ⁺ (4 → 3)	356.734	ASTE	22	0.7	Central Pos.	...	0.11	0.1
C ¹⁸ O(3 → 2)	329.331	ASTE	23	0.7	Central Pos.	...	0.43	0.05

with a bandwidth of 1 GHz and 2048 frequency channels. The velocity coverage was ~ 870 km s⁻¹ centered at -13.6 km s⁻¹ V_{LSR} with a channel width of 0.42 km s⁻¹. In this transition the emission was mapped within a region of $100'' \times 100''$ in size, with $20''$ angular spacing (roughly full beam spacing), centered on the position of the radio source. System temperatures were typically 270 K. The pointing accuracy is $2''$ – $3''$ calibrated toward the source RT-SCO ($\alpha = 17^{\text{h}}03^{\text{m}}32^{\text{s}}.6$, $\delta = -36^{\circ}55'14''$) and the uncertainty in the absolute flux scale is $\sim 10\%$. The observations were performed in position-switching mode under good atmospheric conditions ($\tau \sim 0.07$) using $\alpha = 16^{\text{h}}58^{\text{m}}15^{\text{s}}$ and $\delta = -39^{\circ}00'00''$ as off-position. The integration time on-source in each position was ~ 30 s, resulting in an rms noise of typically 0.1 K in antenna temperature per channel.

The CO(6 → 5) and CO(7 → 6) observations were made during 2009 August, under good weather conditions (0.35–0.45 mm of water vapor, corresponding to $0.5 < \tau < 0.8$) using an off-position located at $\Delta\alpha = +6^{\text{m}}$. The pointing calibration source was NGC 6334, attaining an accuracy of $\sim 2''$. The front-end consisted of a dual heterodyne SIS receiver array of 7 pixels operating in the 600–720 GHz and 750–950 GHz atmospheric windows, known as the CHAMP⁺ receiver (Güsten et al. 2008). In these two transitions we mapped the emission within a region of $110'' \times 90''$ in size, with $4''$ angular spacing. The back-ends were operated with 2.4 GHz total bandwidth with 2048 channels each.

We complement the above data with additional spectra taken toward the radio source position in the SiO(8 → 7), HCO⁺(4 → 3), and C¹⁸O(3 → 2) lines. The SiO(8 → 7) observations were performed using APEX with the same APEX FFT spectrometer. The integration time was ~ 450 s giving an rms noise of 0.05 K.

The HCO⁺(4 → 3) and C¹⁸O(3 → 2) spectra were taken as part of a larger mapping of the core using the Atacama Submillimeter Telescope Experiment (ASTE) 10 m submillimeter telescope (Ezawa et al. 2004) during 2010 June–July, under good atmospheric conditions ($\tau \sim 0.08$). The front-end consisted in a heterodyne single pixel SiS receiver operating between 324 and 372 GHz. As back-end we used the MAC XF-type digital spectrorcorrelator with 1024 channels. For the C¹⁸O(3 → 2) line we used a total bandwidth of 512 MHz, giving a velocity resolution of about 0.43 km s⁻¹, whereas for the HCO⁺(4 → 3) transition we used a 125 MHz bandwidth, giving a velocity resolution of ~ 0.11 km s⁻¹. A detailed description of these observations will be given elsewhere.

3. RESULTS

Figure 1 shows a grid of the CO spectra in the three observed transitions within a region of $100'' \times 100''$ in size, with $20''$ spacing. The spectra of the CO(6 → 5) and CO(7 → 6)

transitions were convolved to match the angular resolution of the CO(3 → 2) spectra. Clearly seen in this figure is the presence of broad and strong wing emission across the molecular core.

Figure 2 shows the HCO⁺(4 → 3), C¹⁸O(3 → 2), and SiO(8 → 7) spectra observed toward the peak position of IRAS 16562–3959. The HCO⁺(4 → 3) spectrum shows a double-peaked line profile, with a bright blueshifted peak at -14.0 ± 0.2 km s⁻¹ and a weaker redshifted peak at -10.1 ± 0.2 km s⁻¹. On the other hand, the C¹⁸O(3 → 2) line profile shows a symmetric single component with a peak line center velocity of -12.5 km s⁻¹. We also note that the HCO⁺(4 → 3) and the SiO(8 → 7) spectra show prominent high-velocity wings.

The ¹²CO profiles exhibit redshifted wing emission mainly toward the west and south, up to a local standard of rest (LSR) velocity of 14 km s⁻¹, whereas blueshifted wing emission is seen mainly toward the east and north, up to an LSR velocity of -46 km s⁻¹. The full velocity range of the wing emission observed is then ~ 60 km s⁻¹. The radial flow-velocity is defined as $|v_{\text{lsr}} - v_{\text{amb}}|$, where v_{amb} is the systemic velocity of the ambient gas. For the later we adopt a value of -12.5 km s⁻¹, corresponding to the peak velocity of the C¹⁸O(3 → 2) line profile observed toward the central position of the core. The maximum outflow velocity is then ~ 33.5 km s⁻¹ toward the blue and ~ 26.5 km s⁻¹ toward the red. Figure 3 shows contour maps of the wing emission in the three CO lines integrated over the velocity range -46 km s⁻¹ $< v_{\text{lsr}} < -22.0$ km s⁻¹ (blue contours) and -3 km s⁻¹ $< v_{\text{lsr}} < 14$ km s⁻¹ (red contours). The star marks the position of the radio jet (Guzmán et al. 2010). These velocity limits define the ranges in which emission from the high-velocity outflow gas is seen. We note that the blueshifted flow emission is slightly stronger and reaches higher radial flow velocities compared to the redshifted emission, hence the wider blueshifted velocity range. The inner limits of the ranges were selected as to leave out the contribution of the ambient cloud and chosen symmetrically with respect to the ambient cloud velocity.

The spatial distribution of the blueshifted and redshifted emissions shows a quadrupolar morphology, with blueshifted emission seen mainly toward the east and north and redshifted emission seen toward the west and south. We propose that the quadrupolar morphology is produced by the superposition of two bipolar outflows, one along a southeast–northwest direction (hereafter the SE–NW flow) with a position angle (P.A.) of $\sim 107^{\circ}$, and a second outflow aligned roughly in a north–south (N–S) direction. The collimation factors—length divided by width—of the molecular outflows are not high: for the blueshifted and redshifted lobes of the SE–NW outflow are $\gtrsim 2$ and ~ 1 , respectively, whereas for the lobes of the N–S outflow we found values close to 1. These values are similar to those derived for massive outflows (Beuther et al. 2002). We

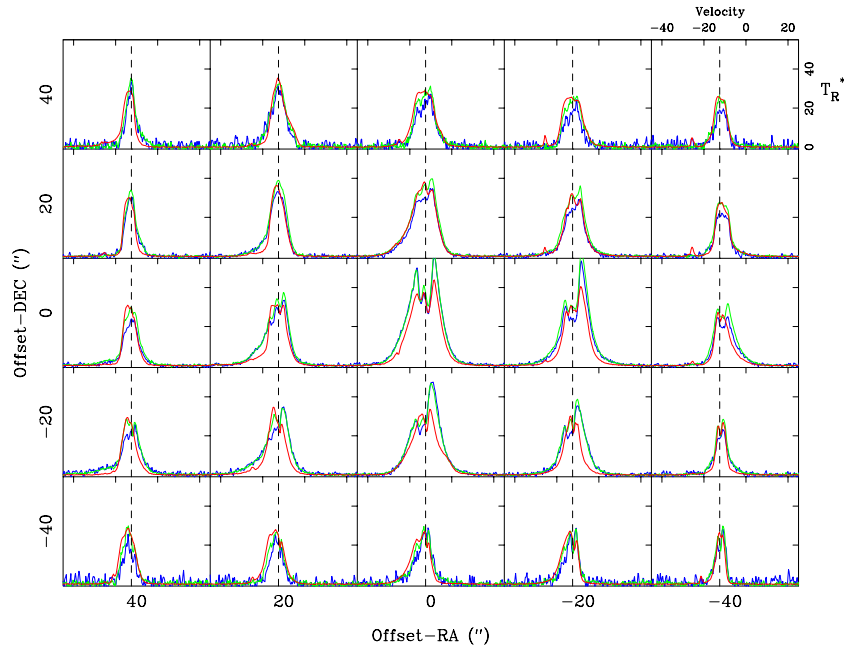


Figure 1. Spectral grid of the CO(3 \rightarrow 2) (red), CO(6 \rightarrow 5) (green), and CO(7 \rightarrow 6) (blue) emission observed toward IRAS 16562–3959. The grid spacing is 20". Offsets are from the radio source reference position at $\alpha_{2000} = 16^{\text{h}}59^{\text{m}}41.61^{\text{s}}0$, $\delta_{2000} = -40^{\circ}03'43''$. Velocity scale ranges from -45 to 25 km s $^{-1}$. The dashed line in each spectrum marks the ambient velocity adopted of -12.5 km s $^{-1}$.

note, however, that the highest velocity gas emission of the SE–NW outflow appears to be associated with more collimated structures.

Figure 4 presents a contour map of the outflow emission in the CO(6 \rightarrow 5) line overlaid with the 8.6 GHz radio continuum emission observed toward IRAS 16562–3959 (Guzmán et al. 2010). It appears that the SE–NW outflow is associated with the string of radio sources, the peak position of the blueshifted and redshifted lobes being symmetrically displaced from the bright central radio source. The SE blueshifted and NW redshifted lobes extend up to $\sim 27''$ and $\sim 32''$ from the central radio source, respectively. The symmetry axis of the SE–NW outflow is along a direction with a P.A. of $\sim 107^{\circ}$, roughly the same as the P.A. of the symmetry axis of the jet of 110° . As noted in Guzmán et al. (2010), the radio lobes and the jet are not completely aligned, showing a small bending, which is also seen in the SE–NW outflow. Possible bending mechanisms of protostellar jets are discussed in Fendt & Zinnecker (1998), but with the available data we cannot discern between the various alternatives.

The spectroscopic signatures of the HCO $^{+}$ (4 \rightarrow 3) and C 18 O(3 \rightarrow 2) transitions suggest that the bulk of the molecular gas toward IRAS 16562–3959 is undergoing large-scale inward motions (e.g., Sanhueza et al. 2010). Infalling motions traced by optically thick molecular lines are expected to produce line profiles showing blue asymmetry, whereas optically thin lines are expected to exhibit symmetrical profiles (Mardones et al. 1997).

4. ANALYSIS AND DISCUSSION

4.1. K_s -band Evidence of Shocked H_2

Figure 5 presents an image of the Two Micron All Sky Survey (2MASS; Skrutskie et al. 2006) K_s -band emission across an $8' \times 8'$ region of the sky, centered near IRAS 16562–3959. Clearly seen toward the center is diffuse emission along the SE–NW direction extending by more than 1.5 on each side of the bright radio source. The P.A. of this diffuse K_s -band emission is

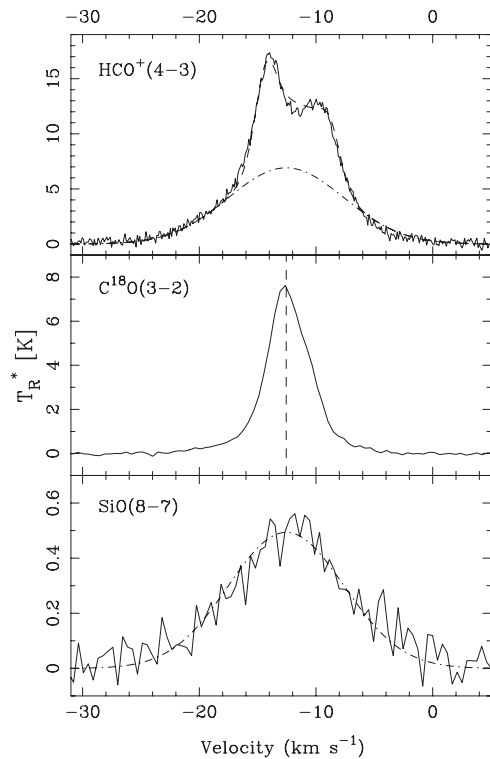


Figure 2. Spectra observed toward the peak position of IRAS 16562–3959. Top: HCO $^{+}$ (4 \rightarrow 3) spectrum. The dashed line shows the best fit using a collapsing envelope model from Myers et al. (1996) plus an outflow component (dot-dashed line). Middle: C 18 O(3 \rightarrow 2) spectrum. The dashed line indicates the adopted ambient cloud velocity of -12.5 km s $^{-1}$. Bottom: SiO(8 \rightarrow 7) spectrum. The dot-dashed line is not a fit but shows the outflow profile fitted to the HCO $^{+}$ (4 \rightarrow 3) spectrum scaled by 1/14.

$\sim 104^{\circ}$, a value similar to that of the SE–NW molecular outflow. We suggest that the outer parts of this diffuse emission, which are at greater distances from the central jet than the molecular

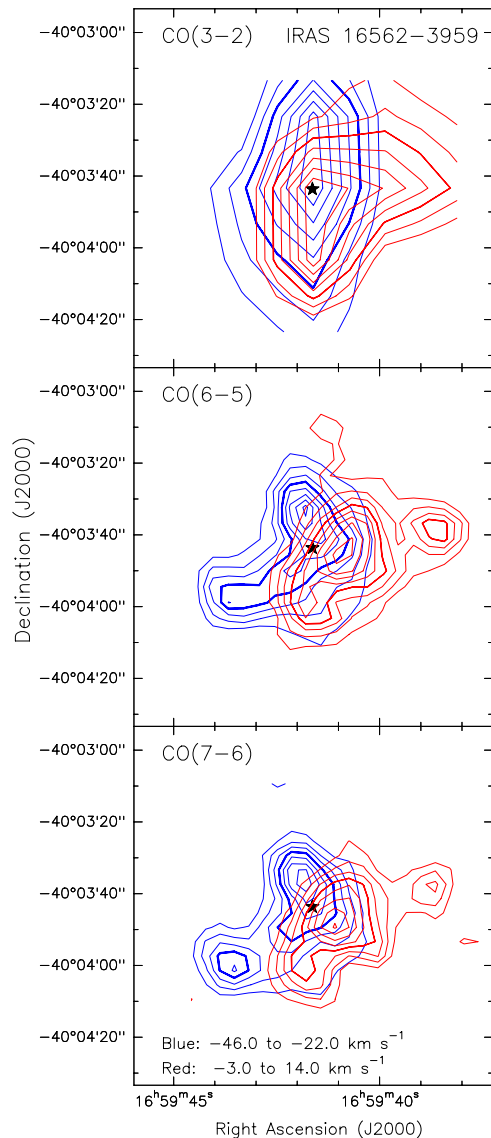


Figure 3. Contour maps of the velocity-integrated CO line wing emission toward IRAS 16562–3959. Blue lines represent emission integrated over the velocity range $-46 \text{ km s}^{-1} < v_{\text{lsr}} < -22 \text{ km s}^{-1}$, which is blueshifted with respect to the ambient velocity of -12.5 km s^{-1} , and red lines represent emission integrated over the velocity range $-3 \text{ km s}^{-1} < v_{\text{lsr}} < 14 \text{ km s}^{-1}$, which is redshifted with respect to the ambient velocity. The star marks the position of the jet source. Contour levels are 20%, 30%, 40%, 50%, 60%, 70%, 80%, and 90% of the peak emission. Top: CO(3 → 2) emission. Peak blueshifted emission: 49.0 K km s^{-1} . Peak redshifted emission: 26.4 K km s^{-1} . Middle: CO(6 → 5) emission. Peak blueshifted emission: 85.2 K km s^{-1} . Peak redshifted emission: 49.6 K km s^{-1} . Bottom: CO(7 → 6) emission. Peak blueshifted emission: 85.1 K km s^{-1} . Peak redshifted emission: 54.3 K km s^{-1} .

lobes, are tracing shocked H_2 – $2.12 \mu\text{m}$ gas produced by an older episode of mass ejection from the jet.

We note that the near-infrared (NIR) emission detected close to the central jet comes mainly from a region associated with the SE blueshifted lobe. Much less emission is seen associated with the NW redshifted lobe (see also Figure 4 from Guzmán et al. 2010). A similar situation has been observed in other bipolar outflows, for example in BHR 71 (Bourke 2002). Most likely, the strong NIR emission from the eastern, closer to the jet, feature corresponds to scattered light from the inner walls of the outflow cavity.

The second outflow detected toward IRAS 16562–3959 is roughly aligned in the N–S direction, having a symmetry axis

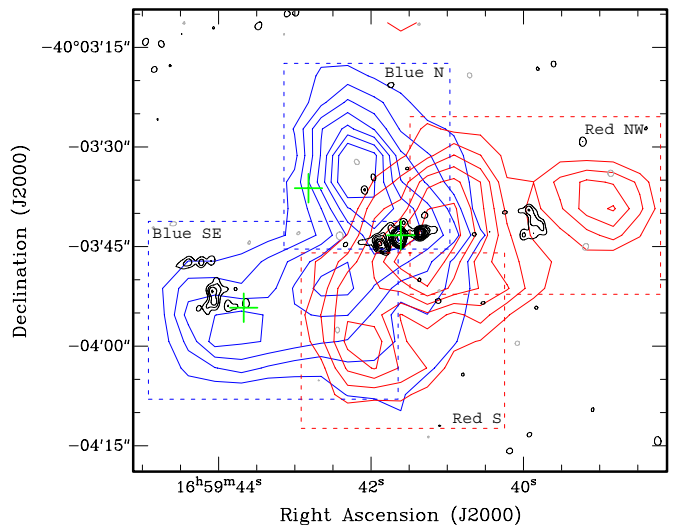


Figure 4. Map of the CO(6 → 5) velocity-integrated blueshifted (blue contours) and redshifted (red contours) wing emission, as in Figure 3, overlaid with a map of the 8.6 GHz emission (black contours) showing the string of radio emission described in Guzmán et al. (2010). The red and blue dashed-line boxes mark the regions of the sky where the CO emission was integrated to determine the parameters of the red and blue lobes, respectively. The green crosses mark the position of the OH masers.

with a P.A. of 7° . The blueshifted lobe extends up to $\sim 21''$ to the north of the jet and the redshifted lobe up to $\sim 17''$ to the south. Interestingly, Figure 5 also shows the presence of diffuse K_s -band emission roughly along the N–S direction but located about $3.4'$ north and $2'$ south from the N–S molecular outflow. This emission exhibits a bow shock morphology, characteristic of Herbig–Haro objects, most likely produced by the interaction of a collimated flow with the ambient medium. However, as noted in Guzmán et al. (2010), this diffuse emission is bright in all MIR *MSX* bands. In consequence, it is likely that this K_s -band emission has an important NIR continuum component, in addition to the H_2 – $2.12 \mu\text{m}$ line.

We suggest that there is an independent source of energy driving the N–S outflow, although from the available data we are not able to pinpoint its location. Guzmán et al. (2010) have already argued, on the basis that a re-orientation of the central jet in $\sim 90^\circ$ seems physically unfeasible, for the presence of two high-mass YSOs driving the outflows.

4.2. Parameters of the Outflows

To compute physical parameters of the molecular outflows we followed the standard formalism described in Bourke et al. (1997), Garden et al. (1991), and Goldsmith & Langer (1999), assuming that the high-velocity gas is optically thin and its excitation state can be described by a single excitation temperature. If the transitions are sub-thermally excited, then the derived excitation temperature would correspond to a lower limit of the kinetic temperature of the outflowing gas.

A general discussion of the sources of errors have been given by Margulis & Lada (1985) and Cabrit & Bertout (1990). The main sources of error arise from the difficulty in determining the contribution to the outflow in the velocity range of the ambient cloud, and not knowing the flow inclination. To be conservative we adopt as velocity boundary between the blue and red wing emission and the ambient emission the values of -22.0 and -3.0 km s^{-1} , respectively.

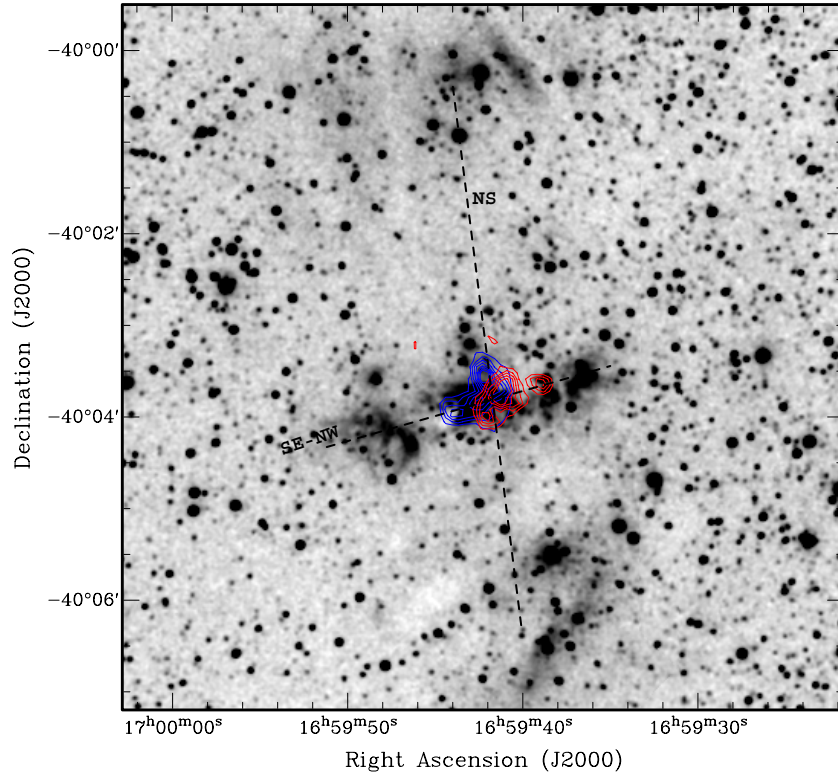


Figure 5. Gray scale: K_s -band 2MASS emission. Overlaid are contours of the blueshifted (blue contours) and redshifted (red contours) CO(6 \rightarrow 5) emission. The approximate directions of the SE-NW and N-S flows are indicated by dashed lines. (A color version of this figure is available in the online journal.)

4.2.1. Column Densities and Excitation Temperatures

The column density of CO molecules in the velocity range $[v_1, v_2]$ is given by

$$N_{\text{CO}} = 2.31 \times 10^{14} \frac{(T_{\text{ex}} + hB/3k) \exp(E_J/kT_{\text{ex}})}{1 - \exp(-hv/kT_{\text{ex}}) (J+1)} \times \int_{v_1}^{v_2} \tau_v dv \text{ cm}^{-2}, \quad (1)$$

where T_{ex} is the rotational excitation temperature in K, J is the rotational quantum number of the lower state, and $E_J = hBJ(J+1)$ is the energy of level J , with B the rotational constant of the CO molecule. The frequency of the CO($J+1 \rightarrow J$) transition is $\nu \approx 2B(J+1)$ and τ_v is the opacity of the material moving at velocity v (the latter measured in km s^{-1}). Assuming a beam filling factor of 1, the observed main-beam brightness temperature is related to the opacity by

$$T_{\text{mb}}^*(\nu) = [J_r(\nu, T_{\text{ex}}) - J_r(\nu, T_{\text{bg}})](1 - e^{-\tau_\nu}), \quad (2)$$

where $J_r(\nu, T)$ is defined as

$$J_r(\nu, T) = \frac{hv/k}{\exp(hv/kT) - 1}.$$

In the optically thin limit,

$$T_{\text{mb}}^*(\nu) = [J_r(\nu, T_{\text{ex}}) - J_r(\nu, T_{\text{bg}})]\tau_\nu. \quad (3)$$

From the observations of the emission in the three transitions of CO it is possible to determine both the excitation temperature

and column density of the outflowing gas using Equations (1) and (3), and taking $T_{\text{bg}} = 2.7$ K. Figure 6 presents plots of the velocity-integrated emission versus the J quantum number from the blueshifted and redshifted lobes of both outflows and from the ambient gas. The velocity intervals of integration for the blueshifted and redshifted emissions are $[-46, -22]$ and $[-3, 14]$ km s^{-1} , respectively. The areas over which the emission has been spatially integrated for the different outflow lobes are shown in Figure 4. In the positions in which lobes overlap we assumed equal contributions from each lobe, except at offset position ($\Delta\alpha = +20''$, $\Delta\delta = +20''$) where we assumed that the contribution to the blueshifted emission from the N-S and SE-NW outflows are in a 2:1 ratio. Dashed lines show the results of the best fit to the observed integrated emission assuming optically thin conditions and filling factors of 1. The derived CO column densities and excitation temperatures are given in Table 2. The parameters for the ambient cloud were derived using the spectra observed at offset position ($\Delta\alpha = +40''$, $\Delta\delta = -40''$), which does not show evidence of self-absorption or high-velocity gas.

One of the less well-known physical parameter of the gas in molecular outflows is its kinetic temperature. Our observations of three lines of CO allowed us to determine the excitation temperature of both outflows, concluding that their kinetic temperatures are high. The derived excitation temperatures of the high-velocity gas are 145 and 120 K, respectively, for the blueshifted and redshifted lobes of the SE-NW outflow, and 76 and 88 K, respectively, for the blueshifted and redshifted lobes of the N-S outflow. All of these temperatures are considerably higher than the ambient cloud temperature, indicating that part of the energy used in accelerating the gas has also heated it.

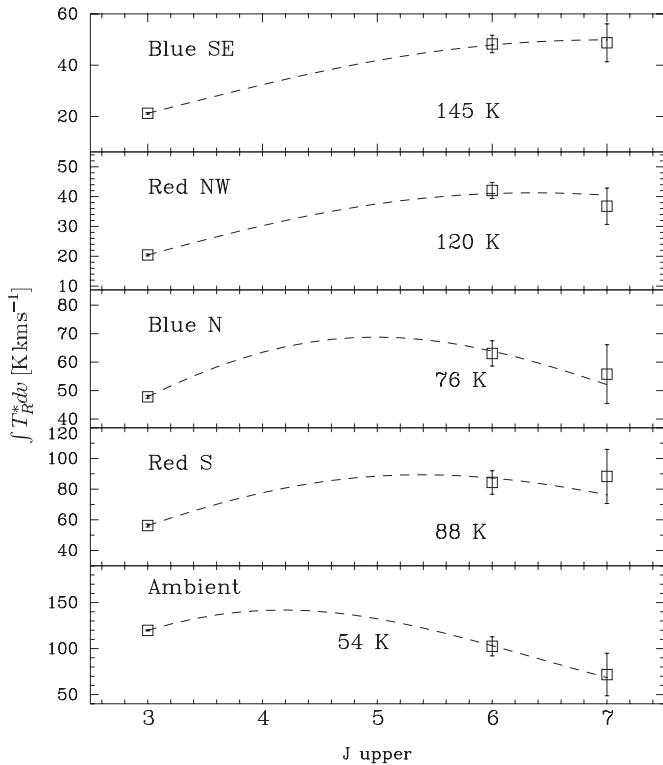


Figure 6. Velocity-integrated CO emission vs. upper rotational quantum level J . The range of velocity integration for the different features is indicated in Column 2 of Table 2. The outflow emission from the lobes has been averaged over the zones indicated in Figure 4. Dashed lines represent fits to the data using Equation (1) and optically thin conditions. The derived column densities and excitation temperatures are shown in Table 2, and the latter are also indicated in each panel. Top: wing emission from the blueshifted lobe of the SE–NW outflow. Middle top: wing emission from the redshifted lobe of the SE–NW outflow. Middle: wing emission from the blueshifted lobe of the N–S outflow. Middle bottom: wing emission from the redshifted lobe of the N–S outflow. Bottom: ambient cloud emission, measured from spectrum observed at offsets $\Delta\text{R.A.} = +40''$, $\Delta\text{decl.} = -40''$.

4.2.2. Mass, Momentum, and Momentum Rates

The mass in the outflows can be computed from the derived column densities as

$$M = [\text{H}_2/^{12}\text{CO}]\mu_m \sum N_{\text{CO}} dA, \quad (4)$$

where N_{CO} is given by Equation (1), μ_m is the mean molecular mass per H_2 molecule, $[\text{H}_2/^{12}\text{CO}]$ is the molecular hydrogen to carbon monoxide abundance ratio, dA is the size of the emitting area in an individual position, and the sum is over all the observed positions.

The derived masses of the SE–NW and N–S flows are given in Columns 2–4 of Table 3. Column 2 gives the mass of the outflows in the high-velocity (HV) ranges, computed using the column densities given in Table 2, $\mu_m = 2.3m_{\text{H}}$, and $[\text{H}_2/^{12}\text{CO}] = 10^4$. To estimate the contribution to the mass from the outflowing gas emitting in the same velocity range as the ambient cloud gas, referred as the low-velocity (LV) outflow gas, we followed the prescription of Margulis & Lada (1985). Using expression (A16) of Bourke et al. (1997), adopting as velocity boundary between the wing and ambient emissions values of -22 km s^{-1} in the blue side and -3 km s^{-1} in the red side, we estimate that the mass of the low-velocity flow is 1.26 and $0.92 M_{\odot}$ for the SE–NW and N–S flow, respectively. The total masses for the SE–NW and N–S outflows are then 1.92 and $1.32 M_{\odot}$, respectively.

Table 2
Parameters of High-velocity and Ambient Gas

Lobe	V range (km s^{-1})	N_{CO} (10^{16} cm^{-2})	T_{ex} (K)
Blue SE	$[-46.0, -22]$	8.72	145
Red NW	$[-3, 14.0]$	6.06	120
Blue N	$[-46.0, -22]$	5.70	76
Red S	$[-3, 14.0]$	3.37	88
Ambient	$[-22, -3]$	11.30	54

The momentum of the gas can be estimated from the first moment of the line emission. This procedure has the advantage of using the detailed information of the spectra. The derivation of flow kinematic and dynamical parameters by using moments of the spectrum has been described in detail by Calvet et al. (1983). For this purpose we used the $\text{CO}(3 \rightarrow 2)$ spectra since it has the best signal-to-noise ratio, but note that using any of the three transitions render consistent values within 10%. In the low-velocity range we selected -12.5 km s^{-1} , the ambient cloud velocity, as the limiting value between what we consider blueshifted or redshifted gas. The derived momentums are given in Columns 5–7 of Table 3. They were computed using the observed radial velocities, and thus correspond to strict lower limits. To correct for inclination they should be multiplied by a factor $(\cos i)^{-1}$ where i is the angle between the flow and the line of sight.

The average momentum rates can be estimated from the second moment of the $\text{CO}(3 \rightarrow 2)$ spectra, from the expression

$$\dot{P} = [\text{H}_2/^{12}\text{CO}]\mu_m \sum_{\text{lobe}} F(J, T_{\text{ex}}) \frac{\int \tau_v v^2 dv}{r_{\text{char}}} dA, \quad (5)$$

where $F(J, T_{\text{ex}})$ is the term that multiplies the integral in the right-hand side of Equation (1), and we select $J = 2$ as indicated previously. The opacity τ_v can be estimated from Equation (3) in the optically thin limit. The characteristic radius of the flow, r_{char} , is a sensitive parameter that regulates the size of the flow and the timescale (after dividing it by the flow velocity) over which the momentum has been deposited in the pre-stellar core. In Equation (5), we consider r_{char} as the projected radius, taken as the length subtended by $30''$ (0.23 pc at 1.6 kpc) for the high-velocity flow; and as the length subtended by $20''$ for the low-velocity gas. The momentum rates derived for the outflows, with no correction for inclination, are given in Columns 8–10 of Table 3. A correction for inclination will require multiplication by the factor $(\cos i)^{-2} \sin i$.

Toward each of the lobes of the SE–NW flow we observe wing emission at both blueshifted and redshifted velocities. This allows us to make an estimate of its inclination using the expression (Cabrit & Bertout 1990, 1986)

$$\tan i = \frac{R + 1}{R - 1} \frac{1}{\tan \theta},$$

where R is the ratio between the maximum observed blueshifted and redshifted velocities toward a single lobe and θ is half of the total aperture angle of the outflow. We estimate the latter from the 2MASS K_s -band image, obtaining a value of $\sim 23^\circ$. Using this and the observed values of R , we derive an inclination angle for the SE–NW outflow of 80° . Correction factors for the momentum and momentum rate are then 5.8 and 33, respectively. The inclination-corrected momentum of the SE–NW flow is $89 M_{\odot} \text{ km s}^{-1}$ and the inclination-corrected momentum rate (calculated with Equation (5)) is $3.0 \times 10^{-2} M_{\odot} \text{ km s}^{-1} \text{ yr}^{-1}$.

Table 3
Outflow Parameters^a

Outflow	Mass (M_{\odot})			Momentum ^b ($M_{\odot} \text{ km s}^{-1}$)			Momentum Rate ^b ($10^{-4} M_{\odot} \text{ km s}^{-1} \text{ yr}^{-1}$)		
	HV	LV	Total	HV	LV	Total	HV	LV	Total
SE–NW outflow									
Blue SE	0.39	0.61	1.00	5.8	2.9	8.7	4.3	1.2	5.5
Red NW	0.27	0.65	0.92	3.5	3.1	6.6	2.2	1.3	3.5
Total	0.66	1.26	1.92	9.3	6.0	15.3	6.5	2.5	9.0
N–S outflow									
Blue N	0.25	0.49	0.74	3.4	2.4	5.8	3.3	1.0	4.3
Red S	0.15	0.43	0.58	1.9	2.0	3.9	1.7	0.8	2.5
Total	0.40	0.92	1.32	5.3	4.4	9.7	5.0	1.8	6.8

Notes.

^a HV and LV denote parameters of the high-velocity and low-velocity flows, derived as described in Section 4.2.2.

^b Not corrected for inclination.

One of the principal problems of the previous estimation is that one assumes that the momentum has been injected into the molecular gas continuously, despite the molecular high-velocity gas and the radio image evidence that the emission comes from discrete episodes of mass ejection. Therefore, the momentum rates calculated are time-averaged over the history of mass ejections from the protostar, that probably includes long periods of inactivity.

The momentum flux of the ionized jet reported by Guzmán et al. (2010)—corrected for 80° inclination—is $\sim 4.2 \times 10^{-4} M_{\odot} \text{ km s}^{-1} \text{ yr}^{-1}$, while that of the associated SE–NW molecular flow ($3.0 \times 10^{-2} M_{\odot} \text{ km s}^{-1} \text{ yr}^{-1}$) is approximately 70 times greater. Also, the observed momentum of the flow could not have been deposited in the ambient gas by such a jet in less than $\sim 10^5$ yr, a time larger than the estimated life of the YSO. A possible solution to this inadequacy of the jet as the driving source of the molecular outflow could be that the protostellar jet is only partially ionized, being its mass and momentum loads greater than the estimated ones.

4.3. Further Tracers of Shocks

In addition to the high-velocity CO gas and its high excitation temperature, further evidence for the presence of shocks is provided by the detection of SiO emission. SiO is thought to be generated when strong shocks pass through dense molecular gas disrupting dust grains (Hartquist et al. 1980; Caselli et al. 1997). Therefore, we expect that SiO emission should be associated with the outflowing gas. This suggestion is strongly supported by the almost identical shape of the profile of the high-velocity outflowing gas emission observed in the $\text{HCO}^+(4 \rightarrow 3)$ line and the $\text{SiO}(8 \rightarrow 7)$ profile (see next section and the bottom panel of Figure 2).

The observed velocity-integrated SiO emission is $\sim 4.6 \text{ K km s}^{-1}$. Assuming the SiO line is emitted in optically thin conditions with an excitation temperature of 75 K, we derive a column density of SiO of $7.8 \times 10^{12} \text{ cm}^{-2}$. Assuming $T_{\text{ex}} = 100 \text{ K}$ only changes this value by 3%. The critical density for the $\text{SiO}(8 \rightarrow 7)$ transition is $\sim 10^8 \text{ cm}^{-3}$, considerably larger than the ambient density, indicating that thermal collision with ambient H_2 is not the main mechanism for excitation to this rotational state. The good correlation between the SiO emission and the high-velocity component suggests that this mechanism is shock excitation of the gas.

There are three OH masers reported in the literature toward the region (Caswell 1998, 2004). Figure 4 shows their positions

with green crosses. The closest to the peak of the SE blueshifted lobe corresponds to a 1720 MHz OH maser, which has an LSR velocity of -24.5 km s^{-1} . The 1720 MHz OH masers are usually associated with shocked gas (Brogan 2007), and the location and velocity of this particular one is consistent with it being pumped by the shock produced by the high-velocity blueshifted flow. Caswell (2004) reported Zeeman splitting on this maser implying a $\sim 10 \text{ mG}$ magnetic field, the latter probably enhanced through shock compression. We also note that the excitation temperature of the outflowing gas—derived in Section 4.2.1—is above the minimum temperature required to pump this maser collisionally ($T \approx 90 \text{ K}$; Elitzur 1976; Hoffman et al. 2003). We conclude that this maser is associated with the IRAS 16562–3959 core and to the shocked gas in the SE–NW outflow, and not located farther away at $\sim 2.5 \text{ kpc}$ as previously suggested by Caswell (2004).

The other two masers correspond to 1665/1667 MHz OH masers: one is coincident with the central jet source ($V_{\text{OH}} = -12.7 \text{ km s}^{-1}$) and the other, located about $15''.6$ NE from the jet ($V_{\text{OH}} = -15.5 \text{ km s}^{-1}$), is associated with an MIR source in the field, unseen in 2MASS but conspicuous in the *Spitzer* 8.0 μm band (see Figure 4 from Guzmán et al. 2010).

4.4. Infall of Molecular Core

To derive infall parameters we modeled the $\text{HCO}^+(4 \rightarrow 3)$ profile using the simple analytic model of contracting clouds of Myers et al. (1996). We first subtracted a broad Gaussian profile, shown as a dash-dotted line in the upper panel of Figure 2, to take into account the contribution of the high-velocity outflow. We note that the shape of the subtracted Gaussian is very similar to the profile of the $\text{SiO}(8 \rightarrow 7)$ emission which is expected to come mainly from the outflowing gas. This is illustrated in the bottom panel of Figure 2, which shows the $\text{SiO}(8 \rightarrow 7)$ spectrum and superimposed the broad Gaussian profile fitted to the $\text{HCO}^+(4 \rightarrow 3)$ emission (scaled by a factor of $1/14$).

The simple model can very well fit the observed profile (see the top panel of Figure 2). The parameters of the best fitting collapse model are: an infall velocity (V_{in}) of 0.35 km s^{-1} , a velocity dispersion (σ) of 1.54 km s^{-1} , an optical depth (τ_0) of 3.8, a kinetic temperature of 32 K, and a rest-frame velocity of the collapsing envelope of -11.4 km s^{-1} . This velocity is about 1 km s^{-1} greater than the ambient cloud velocity adopted here, but we note that Bronfman et al. (1996) also reported a velocity of -11.6 km s^{-1} in the $\text{CS}(2 \rightarrow 1)$ transition detected toward this source.

In order to estimate the envelope accretion rate we consider the mass distribution model for the parent cloud derived from dust emission in Guzmán et al. (2010). Assuming that the extent of the contracting gas along the line of sight is comparable to the size of the cloud of about 0.7 pc FWHM (1.5 at 1.6 kpc) and using the derived values of the infall speed, molecular density, and core size, we obtain a mass infall rate \dot{M}_{in} of $\sim 3.7 \times 10^{-4} M_{\odot} \text{ yr}^{-1}$, a value one order of magnitude below the estimation from the spectral energy distribution fitting reported by Guzmán et al. (2010). The value derived for the infall velocity is also about one order of magnitude below the free-fall velocity expected for the entire cloud, suggesting a retarded collapse.

5. SUMMARY

We undertook molecular line observations toward the IRAS 16562–3959 massive and dense core in three CO transitions: CO(3 → 2), CO(6 → 5), and CO(7 → 6), using APEX, with angular resolutions ranging from $\sim 8''$ to $\sim 17''$. Additionally, HCO⁺(4 → 3), C¹⁸O(3 → 2), and SiO(8 → 7) observations were performed toward the center of the core. The main results and conclusions presented in this paper are summarized as follows.

1. High-velocity molecular gas was detected in the three CO transitions, spanning a range in radial velocity of $\sim 60 \text{ km s}^{-1}$.
2. The morphology of the high-velocity emission, best shown in the CO(6 → 5) and CO(7 → 6) lines, is quadrupolar. We conclude that this morphology is due to the presence of two collimated bipolar outflows, one lying in the SE–NW direction and another in the N–S direction. The more extended bipolar flow is in the SE–NW direction, consistent with being excited by the ionized jet detected toward this MYSO (Guzmán et al. 2010). Extended K_s -band emission probably tracing excited H₂–2.12 μm is also associated with the SE–NW flow.
3. From an analysis of the emission in the three CO lines we derive that the excitation temperature of the high-velocity gas, assumed optically thin, is high, with values of ~ 140 in the lobes of the SE–NW outflow and ~ 90 in the lobes of the N–S outflow. This suggests that the excitation of the gas is related to the acceleration mechanism, namely, shock-induced acceleration.
4. The total mass in the SE–NW and N–S outflows are, respectively, 1.9 and 1.3 M_{\odot} . The momentum and momentum rates derived for the SE–NW outflow, corrected for an inclination angle of 80° , are $89 M_{\odot} \text{ km s}^{-1} \text{ yr}^{-1}$ and $3.0 \times 10^{-2} M_{\odot} \text{ km s}^{-1} \text{ yr}^{-1}$, respectively. These values are characteristics of flows driven by young massive stellar objects with high luminosities.
5. The molecular core in which the outflow is embedded presents evidence of being in gravitational contraction as shown by the blue asymmetric peak seen in the HCO⁺(4 → 3) transition observations. The derived mass infall rate is of the order of a few times $10^{-4} M_{\odot} \text{ yr}^{-1}$ and the infall velocity is $\sim 0.35 \text{ km s}^{-1}$, which is a factor of ~ 10 below the velocity expected from a free-fall collapse.

This implies that the collapse has been retarded and not gravitationally dominated.

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