Detection of a magnetic field in three old and inactive solar-like planet-hosting stars

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\textbf{ABSTRACT}

\textbf{Aims.} Our understanding of magnetic fields in late-type stars is strongly driven by what we know of the solar magnetic field. For this reason, it is crucial to understand how typical the solar dynamo is. To do this, we need to compare the solar magnetic field with that of other stars as similar to the Sun as possible, both in stellar parameters and age, hence activity. We present here the detection of a magnetic field in three planet-hosting solar-like stars having a mass, age, and activity level comparable to that of the Sun.\textbf{\quad Methods.} We used the HARPSpol spectropolarimeter to obtain high-resolution high-quality circularly polarised spectra of HD 70642, HD 117207, and HD 154088, using the least-squares deconvolution technique to detect the magnetic field. From the Stokes $I$ spectra, we calculated the log $R'\!HK$ activity index for each star. We compared the position of the stars in the Hertzsprung-Russell diagram to evolutionary tracks, to estimate their mass and age. We used the lithium abundance, derived from the Stokes $I$ spectra, to further constrain the ages.

\textbf{Results.} We obtained a definite magnetic field detection for both HD 70642 and HD 154088, while for HD 117207 we obtained a marginal detection. Due to the lower signal-to-noise ratio of the observations, we were unable to detect the magnetic field in the second set of observations available for HD 117207 and HD 154088. On the basis of effective temperature, mass, age, and activity level the three stars can be considered solar analogs.

\textbf{Conclusions.} HD 70642, HD 117207, and HD 154088 are ideal targets for a comparative study between the solar magnetic field and that of solar analogs.

\textbf{Key words.} magnetic fields – stars: activity – stars: magnetic field

\section{1. Introduction}

Our knowledge about magnetic fields in late-type stars is strongly driven by what we know of the Sun; however, how typical is the solar magnetic field? For example, the Sun’s magnetic field and activity cycle might be rather peculiar due to the presence of planets, Jupiter in particular (Zaqarashvili 1997; Wilson et al. 2008), which has an orbital period consistent with the periodicity of the solar cycle. One way to look for an answer to “how typical is the solar dynamo?” is to look for and characterise the magnetic field of stars as massive, old, and inactive as the Sun.

Petit et al. (2008) determined the geometry of the magnetic field for four solar-like stars, concluding that slowly rotating (inactive) stars tend to host a stronger poloidal component of the magnetic field compared to the toroidal one, with the latter increasing its strength at the expense of the former with increasing rotation (activity). Of the four stars analysed by Petit et al. (2008), only 18 Sco (HD 146233) is as old and inactive as the Sun, while the other three stars are progressively younger and therefore more active. The magnetic field geometry of 18 Sco is similar to that of the Sun, but one comparison star is not statistically significant. In general, any such comparison should also aim to sample over the stellar activity cycle.

In April 2012 we observed 29 planet-hosting stars aiming to look for and study their magnetic fields and possible connection to star-planet interaction. Most of the stars for which we detected a magnetic field were either young and active or evolved away from the main sequence (Hertzsprung-gap stars). The results of this programme will be published in a separate paper. Here we present the detection of a magnetic field in three old and inactive solar-like planet-hosting stars: HD 70642, HD 117207, and HD 154088. Together with 18 Sco, these objects will ultimately allow us to understand how typical is the solar dynamo, as well as detecting a connection between the stellar activity cycle and giant planets, if any.

\section{2. Observations and magnetic field detection}

We observed the three stars with the HARPSpol polarimeter (Snik et al. 2011; Piskunov et al. 2011) feeding the HARPS spectrograph (Mayor et al. 2003) attached to the ESO 3.6-m telescope in La Silla, Chile. The observations, covering the 3780–6910 Å wavelength range with a spectral resolution $R \sim 115 000$, were obtained using the circular polarisation analyser. We observed each star with a sequence of four sub-exposures obtained rotating the quarter-wave retarder plate by 90° after each exposure, i.e., 45°, 135°, 225° and 315°. The exposure times have been set according to the stellar brightness and sky conditions;
the line mask used by the LSD code separately for each star adopting the stellar parameters listed in Table 2. We extracted the line parameters from the Vienna Atomic Line Database (VALD; Piskunov et al. 1995; Kupka et al. 1999; Ryabchikova et al. 1999). For each star we used about 3000 metal lines, all stronger than 30% of the continuum, avoiding lines with extended wings (e.g. Mg I b and hydrogen lines) and in spectral regions affected by the presence of telluric lines.

Figure 1 shows the obtained LSD profiles, while Table 3 gives the results gathered from their analysis. We defined the magnetic field detection making use of the false alarm probability (FAP; Donati et al. 1992), considering a profile with FAP < 10^-5 as a definite detection (DD), 10^-5 < FAP < 10^-3 as a marginal detection (MD), and FAP > 10^-2 as a non-detection (ND). To further check the magnetic field detections are not spurious, we calculated the FAP for the null profile in the same velocity range as used for the magnetic field, obtaining ND in all cases. We also calculated the FAP for equivalent velocity ranges displaced both redwards and bluewards to sample the continuum in the Stokes I spectrum. The results for both Stokes V and the null profile are given in Col. 6 of Table 3. The much higher FAP obtained in these tests are consistent with our detections being genuine. In addition, we checked both Stokes V and the null profile are consistent with the expected noise properties (e.g., Stokes V uncertainties consistent with the standard deviation of the null profile).

The non-detections resulting from the analysis of the spectra of HD 117207 and HD 154088 on the night of April 21st are most likely due to the lower S/N compared to that obtained on the night of April 17th. This was caused by the highly variable seeing on the night of April 21st, which did not allow us to have enough control on the exposure times in order to achieve the desired S/N.

### 3. Lithium abundance, age, and activity

To find whether the three stars presented here are good comparisons to the Sun, we need to establish their age and activity level as accurately as possible. We estimated the mass and age of the stars by fitting evolutionary tracks to their position in the Hertzsprung-Russell (HR) diagram. We derived luminosities on the basis of the V magnitudes and HIPPARCOS distances (van Leeuwen 2007), adopting the bolometric correction by Flower (1996).

We used the STARS stellar evolution code (Eggleton 1971; Stancliffe & Eldridge 2009) to construct models for stars of between 0.9 to 1.5 M⊙ (with a mass spacing of 0.05 M⊙), for metallicities of Z = 0.01, 0.02 and 0.04. These models have been evolved from the pre-main sequence to the giant branch using 999 meshpoints, a mixing length parameter of 2.0 and without the inclusion of convective overshooting. Figure 2 shows the position of the three target stars in the HR diagram, in comparison with STARS evolutionary tracks calculated with a metallicity of [Fe/H] = 0.3 dex (Z = 0.04). By taking into account the uncertainties on T eff, log L/L⊙, and [Fe/H], for HD 70642, HD 117207 and HD 154088 we estimated an age of 3 ± 3, 5 ± 2, and 5 ± 3 Gyr, respectively. The derived stellar masses are listed in Table 4.

To further constrain the age, we measured the Li abundance from the feature at ~6707 Å, adopting hyperfine structure from Smith et al. (1998) and the meteoritic/terrestrial isotopic ratio Li6/Li7 = 0.08 (Rosman & Taylor 1998). To measure the Li abundance, we fitted synthetic spectra, calculated adopting the LLMODELS (Shulyak et al. 2004) model atmosphere code.
Notes. The S/N (per-pixel) of Stokes $I$ is that of the observed spectrum and it has been calculated over an 0.5 Å region at ~5060 Å. The S/N of Stokes $V$ is that of the LSD profile. Column six lists the FAP calculated in the continuum region bluewards and redwards of the spectral line, over a range as broad as that adopted to derive ($B_i$). The last column lists the number of lines used in the line mask. For all stars we adopted a range of 22 km s$^{-1}$ (i.e., ±11 km s$^{-1}$ from the line center) for the calculation of the magnetic field.

Table 3. Results from the LSD analysis.

<table>
<thead>
<tr>
<th>Star Name</th>
<th>Date</th>
<th>($B_i$) ($V$)</th>
<th>($B_i$) ($N$)</th>
<th>Detection $V$</th>
<th>FAP ($V$) blue cont./red cont.</th>
<th>FAP ($N$) blue cont./red cont.</th>
<th>$S/N$ $I$</th>
<th>$S/N$ $V_{LS}$</th>
<th># lines</th>
</tr>
</thead>
<tbody>
<tr>
<td>HD 70642</td>
<td>17</td>
<td>0.54 ± 0.59</td>
<td>9.7 × 10$^6$</td>
<td>DD 1.4 × 10$^{-1}$/5.4 × 10$^{-1}$</td>
<td>339 17 864</td>
<td>2873</td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>HD 117207</td>
<td>17</td>
<td>1.36 ± 0.62</td>
<td>1.1 × 10$^4$</td>
<td>ND 1.4 × 10$^{-1}$/1.6 × 10$^{-1}$</td>
<td>389 17 471</td>
<td>2859</td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>21</td>
<td>−0.55 ± 0.77</td>
<td>9.9 × 10$^{-1}$</td>
<td>ND 4.8 × 10$^{-1}$/2.3 × 10$^{-1}$</td>
<td>251 13 808</td>
<td>2859</td>
<td></td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>HD 154088</td>
<td>17</td>
<td>2.72 ± 0.49</td>
<td>3.0 × 10$^{-6}$</td>
<td>DD 3.3 × 10$^{-1}$/2.7 × 10$^{-1}$</td>
<td>288 21 280</td>
<td>3299</td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>21</td>
<td>0.96 ± 0.71</td>
<td>1.2 × 10$^{-3}$</td>
<td>ND 5.1 × 10$^{-1}$/9.9 × 10$^{-1}$</td>
<td>206 14 353</td>
<td>3299</td>
<td></td>
<td></td>
<td></td>
<td></td>
</tr>
</tbody>
</table>

The chromospheric activity $S$-index was measured by comparing the flux in the core of the Ca II H&K lines against the flux in the surrounding continuum region, as explained by Jenkins et al. (2006, 2008). We convert this value into the classic $R_{HIK}$ index using the recipes described in Noyes et al. (1984) after conversion to the Mt. Wilson system of measurements (see Duncan et al. 1991). A more thorough description of the calibration method will be presented in a forthcoming paper (Fossati et al., in prep.). The derived $R_{HIK}$ for all three stars lies between ~4.93 and ~5.05, typical of quiescent solar-like stars on the main sequence (Jenkins et al. 2011). In addition, the derived $R_{HIK}$ activity index agrees with that previously obtained by other authors (Jenkins et al. 2006; Wright et al. 2004), showing that our observations have not been taken during a period of excess activity. The higher log $R_{HIK}$ value for HD 70642 agrees with the slightly younger age found for this star from the isochrone fitting, assuming a rotationally driven magnetic dynamo model that weakens with age.

4. Discussion

Except for a slight metal enrichment, HD 70642 and HD 117207 have stellar parameters very similar to that of the Sun. The two
stars have an age in the range of 3–6 Gyr for HD 70642 and 3–7 Gyr for HD 117207, proving they are about as old as the Sun, if not older. For both stars, the derived mass and age values are in agreement in comparison with evolutionary tracks calculated with [Fe/H] $\approx$ 0.3 dex. The position of the Sun is also indicated and its displacement very close to our values. For HD 117207, Wright et al. (2004) measured a rotation period of 42 days, almost double than that of the Sun.

On the basis of the estimated temperature, mass, age and activity, and following the definition given by Soderblom & King (1998), HD 70642, HD 117207, and HD 154088 can be considered solar analogs (note that within the uncertainties, the metallicity of HD 154088 is consistent with the definition of solar analog). Although HD 117207 would require further confirmation of the presence of a detectable magnetic field, the three stars are ideal targets to characterise their magnetic field geometry and compare it to that of the Sun, to study how typical the solar dynamo is. The cooler temperature of HD 154088 might not make it such a good comparison for the Sun, but it will allow one to look for differences in the stellar dynamo as a function of mass, for stars of similar ages. In addition, HD 70642 and HD 117207 host Jupiter-like planets with Jupiter-like orbital periods ($\sim$6 and 7 years, respectively). It will be very valuable to determine the periodicity of their activity cycle, to check whether it is comparable with the orbital period of the hosted planets.

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#### Table 4. Derived age range, mass, Li abundance and log $R'_{HIK}$ in comparison to that of the quiet Sun.

<table>
<thead>
<tr>
<th>Star name</th>
<th>Age [Gyr]</th>
<th>Mass $M_\odot$</th>
<th>log n(Li)</th>
<th>log $R'_{HIK}$</th>
</tr>
</thead>
<tbody>
<tr>
<td>HD 70642</td>
<td>3–6</td>
<td>1.05(10)</td>
<td>$&lt;$0.85</td>
<td>$-$4.935(14)</td>
</tr>
<tr>
<td>HD 117207</td>
<td>3–7</td>
<td>1.08(03)</td>
<td>$&lt;$0.85</td>
<td>$-$5.043(24)</td>
</tr>
<tr>
<td>HD 154088</td>
<td>3–8</td>
<td>0.97(05)</td>
<td>$&lt;$0.75</td>
<td>$-$4.992(40)</td>
</tr>
<tr>
<td>Sun</td>
<td>4.6</td>
<td>1.00</td>
<td>1.05</td>
<td>$-$4.990</td>
</tr>
</tbody>
</table>

Notes. The uncertainties given for mass and log $R'_{HIK}$ corresponds to that of the last digits.

#### Fig. 2. Position of HD 70642, HD 117207 and HD 154088 in the HR diagram in comparison with evolutionary tracks calculated with [Fe/H] = 0.3 dex. The position of the Sun is also indicated and its displacement in comparison with evolutionary tracks calculated with [Fe/H] = 0.3 dex. The position of the Sun is also indicated and its displacement in comparison with evolutionary tracks calculated with [Fe/H] = 0.3 dex.