

## SPITZER NUMBER COUNTS OF ACTIVE GALACTIC NUCLEI IN THE GOODS FIELDS

EZEQUIEL TREISTER,<sup>1,2,3</sup> C. MEGAN URRY,<sup>2,4</sup> JEFFREY VAN DUYN,<sup>2</sup> MARK DICKINSON,<sup>5</sup> RANGA-RAM CHARY,<sup>6</sup>  
DAVID M. ALEXANDER,<sup>7</sup> FRANZ BAUER,<sup>7</sup> PRIYA NATARAJAN,<sup>2</sup> PAULINA LIRA,<sup>3</sup> AND NORMAN A. GROGIN<sup>8</sup>

Received 2005 August 11; accepted 2005 November 30

### ABSTRACT

We present mid-infrared observations of active galactic nuclei (AGNs) in the GOODS fields, performed with the *Spitzer Space Telescope*. These are the deepest infrared and X-ray fields to date and cover a total area of  $\sim 0.1$  deg<sup>2</sup>. AGNs are selected on the basis of their hard (2–8 keV) X-ray emission. The median AGN infrared luminosity is at least 10 times larger than the median for normal galaxies with the same redshift distribution, suggesting that the infrared emission is dominated by the central nucleus. The X-ray-to-infrared luminosity ratios of GOODS AGNs, most of which are at  $0.5 \lesssim z \lesssim 1.5$ , are similar to the values obtained for AGNs in the local universe. The observed infrared flux distribution has an integral slope of  $\sim 1.5$ , and there are 1000 sources per square degree brighter than  $\sim 50$   $\mu$ Jy at  $\sim 3$ – $6$   $\mu$ m. The counts approximately match the predictions of models based on AGN unification, in which the majority of AGNs are obscured. This agreement confirms that the faintest X-ray sources, which are dominated by the host galaxy light in the optical, are obscured AGNs. Using these *Spitzer* data, the AGN contribution to the extragalactic infrared background light is calculated by correlating the X-ray and infrared catalogs. This is likely to be a lower limit given that the most obscured AGNs are missed in X-rays. We estimate the contribution of AGNs missed in X-rays, using a population synthesis model, to be  $\sim 45\%$  of the observed AGN contribution, making the AGN contribution to the infrared background at most  $\sim 2\%$ – $10\%$  in the 3–24  $\mu$ m range, depending on wavelength, lower than most previous estimates. The AGN contribution to the infrared background remains roughly constant with source flux in the IRAC bands but decreases with decreasing flux in the MIPS 24  $\mu$ m band, where the galaxy population becomes more important.

*Subject headings:* diffuse radiation — galaxies: active — infrared: galaxies

*Online material:* machine-readable table

### 1. INTRODUCTION

Hard X-ray surveys are not strongly affected by dust obscuration and thus provide a relatively unbiased view of the active galactic nucleus (AGN) phenomenon. Deep *Chandra* X-ray surveys have allowed us to test the basic prediction of population synthesis models of the X-ray background, namely, that a combination of obscured and unobscured AGNs is needed to explain the X-ray background hard spectral shape (Setti & Woltjer 1989; Gilli et al. 2001). According to the AGN unification paradigm, obscuration comes from optically thick dust blocking the central engine along some lines of sight. The temperature in this structure, which can range up to 1000 K (dust sublimation temperature), and the roughly isotropic emission toward longer wavelengths should make both obscured and unobscured AGNs very bright in the mid–far-infrared bands.

This prediction was tested by *Infrared Astronomical Satellite* (*IRAS*; e.g., Rush et al. 1993) and *Infrared Space Observatory* (*ISO*; e.g., Fadda et al. 2002; Matute et al. 2002; Franceschini

et al. 2002) observations, which discovered a population of very luminous infrared sources coincident with X-ray emitters and therefore very good candidates to be obscured AGNs. With relatively bright flux limits, however, these surveys sample the nearby universe or extremely luminous, rare AGNs. Very deep surveys were possible with *ISO* (e.g., Fadda et al. 2002), but only over a very small area. With the *Spitzer Space Telescope*, it is now possible to perform very deep infrared surveys,  $\sim 100$ – $1000$  times deeper than typical *IRAS* observations, over appreciable areas,  $\sim 10$ – $1000$  times more area than the deep *ISO* surveys, thus permitting us, for the first time, to study a representative sample of the AGN population at cosmologically significant distances.

Since AGNs comprise most of the X-ray background and are also very bright infrared sources, it is natural to ask what fraction of the extragalactic infrared background they may contribute. Based on previous observations, the predicted contribution of AGNs to the infrared background ranged from  $\sim 20\%$  (Fabian & Iwasawa 1999) to just a few percent (Xu et al. 2001), depending mostly on the assumptions for the amount of energy reradiated in the infrared and the spatial density of obscured AGNs. The deeper infrared observations possible with *Spitzer* are important for understanding the true contribution of AGNs to the extragalactic infrared background light.

Multiwavelength observations of a large, uniformly selected sample of moderate-luminosity AGNs at cosmologically significant distances provide a complete picture of the AGN emission across the frequency spectrum and establish the total contribution of these sources to the energy budget of the universe. This is one of the main goals of the Great Observatories Origins Deep Survey (GOODS), a 0.1 deg<sup>2</sup> study divided in two fields that overlap with the deepest X-ray observations, the *Chandra* Deep Fields North (Brandt et al. 2001) and South (Giacconi et al. 2001). It consists of

<sup>1</sup> Department of Astronomy, Yale University, P.O. Box 208101, New Haven, CT 06520.

<sup>2</sup> Yale Center for Astronomy and Astrophysics, Yale University, P.O. Box 208121, New Haven, CT 06520.

<sup>3</sup> Departamento de Astronomía, Universidad de Chile, Casilla 36-D, Santiago, Chile.

<sup>4</sup> Department of Physics, Yale University, P.O. Box 208120, New Haven, CT 06520.

<sup>5</sup> National Optical Astronomy Observatory, Tucson, AZ 85719.

<sup>6</sup> *Spitzer* Science Center, California Institute of Technology, MS 220-6, Pasadena, CA 91125.

<sup>7</sup> Institute of Astronomy, Madingley Road, Cambridge CB3 0HA, UK.

<sup>8</sup> Department of Physics and Astronomy, Johns Hopkins University, 3400 North Charles Street, Baltimore, MD 21218.

deep imaging in the optical with the *Hubble Space Telescope* (*HST*; Giavalisco et al. 2004) and in the infrared with the *Spitzer Space Telescope* (Dickinson & Giavalisco 2003). These very deep multiwavelength observations allow a highly complete, relatively unbiased view of the AGN phenomenon and a test of the AGN unification paradigm, well established in the local universe, at higher redshifts.

In this paper we present *Spitzer* observations of the AGNs in the GOODS fields, from 3.6 to 24  $\mu\text{m}$ . In § 2 we describe the *Spitzer* Infrared Array Camera (IRAC) and Multiband Imaging Photometer for *Spitzer* (MIPS) observations. Results and discussion are presented in § 3, addressing the infrared luminosities of AGNs and normal galaxies, the X-ray-to-infrared ratios for GOODS and local AGNs, the infrared number counts of AGNs, and the AGN contribution to the extragalactic infrared background light. Our conclusions are presented in § 4. Throughout this paper we assume  $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ ,  $\Omega_m = 0.3$ , and  $\Omega_\Lambda = 0.7$ .

## 2. SPITZER DATA AND AGN SAMPLE

The GOODS-North and South fields, which cover  $\sim 160 \text{ arcmin}^2$  each, were observed with the *Spitzer Space Telescope* as part of the GOODS Legacy program (PI: M. Dickinson). Observations were performed using the IRAC camera (Fazio et al. 2004a) in the four available bands (3.6–8  $\mu\text{m}$ ) and the MIPS instrument (Rieke et al. 2004) in the 24  $\mu\text{m}$  band. Exposure times per sky pixel for each field were  $\sim 23 \text{ hr}$  in the IRAC bands. The MIPS 24  $\mu\text{m}$  observations, available to us only in the GOODS-North field, had an exposure time of  $\sim 10.3 \text{ hr}$ . Individual images were processed through the *Spitzer* Science Center basic calibrated data pipeline and then postprocessed and combined using customized routines developed by the GOODS team<sup>9</sup> (M. Dickinson et al. 2006, in preparation). Sources were detected using SExtractor (Bertin & Arnouts 1996) on a combined 3.6 and 4.5  $\mu\text{m}$  band image filtered using a  $5 \times 5$  pixel Gaussian profile with a FWHM of  $1''.8$ . Photometry was performed using SExtractor. The AUTO magnitude was used, since aperture corrections, which we do not apply, would be very small, according to simulations<sup>10</sup> (M. Dickinson et al. 2006, in preparation). If instead we perform aperture photometry (with an aperture diameter of  $4''$ ) and make the corresponding aperture corrections, which range from 0.3 mag in the 3.6  $\mu\text{m}$  band to 0.7 at 8  $\mu\text{m}$ , then the AUTO and aperture-corrected magnitudes differ by less than 0.2 mag on average, with no obvious dependence on magnitude and no significant outliers. Therefore, our results do not depend significantly on the method used to perform the photometry.

The formal  $5 \sigma$  background-limited flux limits for an isolated point source in the four IRAC channels are 0.13, 0.22, 1.45, and 1.61  $\mu\text{Jy}$ . In practice, the photometric uncertainty for a given source depends on the degree of crowding in the images and on the details of how the photometry and source deblending are carried out. This increases the photometric errors, which we evaluate using Monte Carlo simulations of artificial sources added to the real data. According to simulations, all the IRAC images are 90% complete at  $\sim 14 \mu\text{Jy}$ .

The MIPS 24  $\mu\text{m}$  image of the GOODS-North field was reduced using a combination of the standard *Spitzer* Science Center pipeline and customized routines created by the GOODS team (R. Chary et al. 2006, in preparation). Photometry was ob-

tained at the IRAC-derived positions using a point-spread function fitting technique. The  $5 \sigma$  flux limit for the MIPS observations is  $\sim 25 \mu\text{Jy}$ , and they are 84% complete at that flux. The relatively large aperture used to calculate fluxes in both the IRAC and MIPS observations means that the light from the AGN and the host galaxy are summed; in any case, given the limited spatial resolution, the two contributions cannot be separated from the infrared imaging alone.

In order to define the AGN sample, we start with the X-ray catalogs for the Chandra Deep Fields North and South of Alexander et al. (2003), with the associations to the GOODS *HST* optical catalog from Bauer et al. (2004), which uses the likelihood method described by Bauer et al. (2000). Specifically, our sample consists of sources detected in the hard (2–8 keV) *Chandra* band that do not have multiple optical counterparts in the GOODS ACS *z*-band image. Only three sources in the South field and six in the North have multiple optical counterparts, so excluding these sources from the sample does not bias our results. (We do include sources with no *z*-band counterpart down to the limit  $z \simeq 27 \text{ mag}$ .) A large fraction of the sample,  $\sim 60\%$ , have spectroscopic redshifts, obtained by Szokoly et al. (2004) in the South field and by Barger et al. (2003) in the North. Most of the X-ray sources are obscured AGNs at moderate redshift,  $0.5 < z < 1$ , and unobscured AGNs up to  $z \sim 4$ . More distant obscured AGNs are too faint for optical spectroscopy (Treister et al. 2004). Sources with spectroscopic redshifts and hard X-ray luminosity lower than  $10^{42} \text{ ergs s}^{-1}$  are probably not AGN dominated, since the most X-ray-luminous star-forming galaxy known, NGC 3256, has a total X-ray luminosity  $L_X \simeq 8 \times 10^{41} \text{ ergs s}^{-1}$  in the 0.5–10 keV band (Lira et al. 2002); so these were removed from our sample. This criterion removes 12 sources in the South field and 25 in the North. With these constraints, the final sample includes 126 X-ray sources in the GOODS-South field and 179 in the North field.

In order to look for *Spitzer* counterparts of the *Chandra* sources, several search radii were tried, looking to maximize the number of sources with counterparts while keeping a low number of sources with multiple counterparts. The best results were obtained using a search radius of  $2''$ . Typical astrometric errors are  $\sim 0''.5$  in X-rays and  $\sim 0''.2$  in the IRAC bands; therefore, astrometric errors can be safely neglected when using a search radius of  $2''$ . In the North field 160 sources (90%) have a single IRAC counterpart, while in the south field 118 sources (94%) have a single counterpart; the rest have multiple counterparts. *This means that 100% of the X-ray sources are also bright in the near-infrared.* Most of the X-ray sources have an infrared counterpart separated by less than  $1''$ , and in many cases this distance is smaller than  $0''.5$ , as shown in Figure 1. Assuming a Poissonian spatial distribution of infrared sources, the median chance of a random association is 0.74%, and therefore we have high confidence that the correct infrared counterpart was identified in all but one or two cases. The full catalog of AGNs detected in the GOODS field used in this work is presented in Table 1.

Our sample only considers AGNs selected based on the X-ray information. Because of this, highly obscured AGNs in which most of the X-ray emission is absorbed will be missed in our sample. The effects of this incompleteness are discussed in § 4. Recently, AGN selection methods based only on IRAC and MIPS colors (e.g., Sajina et al. 2005) or a combination of *Spitzer* infrared and radio properties (e.g., Martínez-Sansigre et al. 2005) were presented. In the work of Sajina et al. (2005), AGNs are found in the region defined by  $0.0 \lesssim \log(S_{4.5}/S_{3.6}) \lesssim 0.3$ ,  $0.2 \lesssim \log(S_{5.8}/S_{3.6}) \lesssim 0.5$ ,  $0.3 \lesssim \log(S_8/S_{3.6}) \lesssim 0.7$ ,  $0.1 \lesssim \log(S_{5.8}/S_{4.5}) \lesssim 0.3$ ,  $0.3 \lesssim \log(S_8/S_{4.5}) \lesssim 0.5$ , and  $0.1 \lesssim \log(S_8/S_{5.8}) \lesssim 0.3$  using only the IRAC bands (Fig. 8 in Sajina et al. 2005).

<sup>9</sup> More details about data reduction can be found at [http://data.spitzer.caltech.edu/popular/goods/Documents/goods\\_dataproducts.html](http://data.spitzer.caltech.edu/popular/goods/Documents/goods_dataproducts.html).

<sup>10</sup> See <http://www.noao.edu/noao/staff/ecm/simulations.html> for more details and calculation of aperture corrections and completeness simulations.

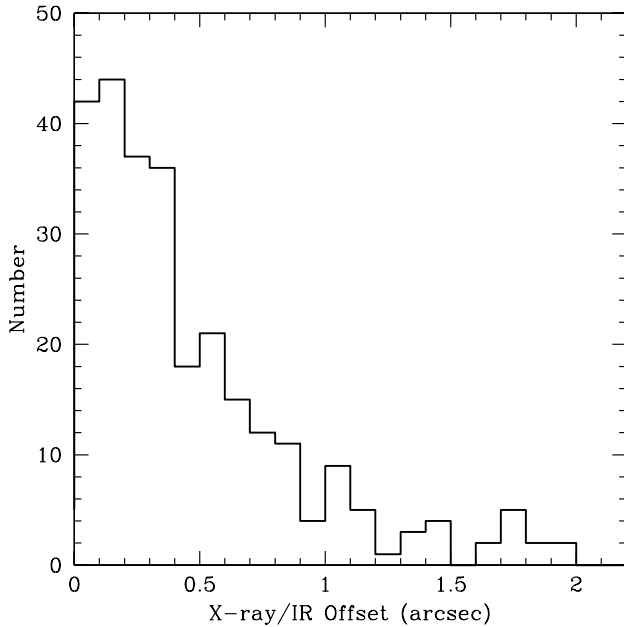


FIG. 1.— Distribution of offsets between the X-ray and infrared positions for sources in our AGN sample. Most of the X-ray sources have an infrared counterpart closer than  $1''$ , while for a large number of AGNs, the distance between the X-ray and infrared positions is smaller than  $0''.5$ . Therefore, we are very confident that the right infrared counterpart for the X-ray sources in our sample was used in all but one or two cases.

Although this criterion was defined only for AGNs at  $z < 1$ ,  $K$ -corrections are very small, so this method can be extended to higher redshifts. Only 79 sources satisfy the AGN criterion of Sajina et al. (2005) in both the GOODS-North and South fields. Of these, only 18 ( $\sim 6\%$ ) sources are detected in hard X-rays. This means that using this selection method misses a large fraction of the AGNs in the field, but it is useful to separate AGN continuum-dominated sources from normal galaxies. This can be explained because a power-law continuum is assumed for the infrared emission in AGNs, while it is possible that most sources are dominated by the torus thermal emission in the infrared (J. Van Duyne et al. 2006, in preparation).

The selection method of Martínez-Sansigre et al. (2005) is based mostly in radio emission, and it is effective for detecting luminous obscured quasars. Using the infrared criterion of Martínez-Sansigre et al. (2005) on the GOODS-North field, where MIPS data were available, only three sources have  $S_{24} > 300 \mu\text{Jy}$  and

$S_{3,6} < 45 \mu\text{Jy}$  and are detected in hard X-rays. Unfortunately, neither spectroscopic nor photometric redshifts are available for those sources, so we cannot (yet) classify them as obscured quasars. The fact that only a small number of sources are selected using this criterion is not surprising, since only very luminous sources are found in this way and such sources are very rare and thus not easily detected in narrow-field surveys such as GOODS.

### 3. RESULTS AND DISCUSSION

#### 3.1. Infrared Luminosity and Colors

When spectroscopic redshifts were not available, we used photometric redshifts computed by Mobasher et al. (2004) for the South field and Barger et al. (2003) for the North field. Using both spectroscopic and photometric redshifts, we were able to calculate luminosities for 93 sources (74%) in the South field and 105 sources (59%) in the North field. In order to study the quality of photometric redshifts we compared them with spectroscopic redshifts for sources in which both measurements are available. The average value of the quantity  $(z_{\text{spec}} - z_{\text{phot}})/(1 + z_{\text{spec}})$  is 0.02, which indicates that photometric redshifts can be used in a statistical sense to study the properties of a sample of sources. However, errors in the photometric redshifts for individual sources can be large. About 30% of the sources compared here had a difference between photometric and spectroscopic redshift determination greater than 0.3. Therefore, photometric redshifts are used only to calculate statistical properties of a sample, not for individual sources.

The observed-frame luminosity distributions for GOODS AGNs in the IRAC  $3.6 \mu\text{m}$  and MIPS  $24 \mu\text{m}$  bands are shown in Figure 2. The vast majority of the sources can be classified as moderate-luminosity AGNs, in the range  $10^{42}$  to  $\sim 10^{44} \text{ ergs s}^{-1}$ . No  $K$ -correction was attempted. However, from the theoretical AGN spectral energy distribution library of Treister et al. (2004), which spans unobscured to fully obscured AGNs over 4 decades in luminosity, we estimate that infrared luminosities in the  $3\text{--}20 \mu\text{m}$  range would change by at most a factor of 2 due to  $K$ -corrections. This is smaller than the bin size used in these histograms, so we do not expect the figure to change significantly were we able to do accurate  $K$ -corrections.

For comparison, in Figure 2 we also show the observed-frame  $3.6$  and  $24 \mu\text{m}$  luminosity distributions for normal galaxies in the GOODS fields, randomly selected to follow the same redshift distribution as the AGN sample. Redshifts for the normal galaxies were obtained from the spectroscopic surveys of Wirth

TABLE 1  
CATALOG OF AGNs IN THE GOODS FIELDS

FIELD	ID	PROB. <sup>a</sup>	X-RAY FLUX ( $\text{ergs cm}^{-2} \text{s}^{-1}$ )			HARDNESS RATIO <sup>b</sup>	INFRARED FLUX ( $\mu\text{Jy}$ )					REDSHIFT		CLASSIFICATION <sup>d</sup>
			Soft (0.5–2 keV)	Hard (2–8 keV)	3.6 $\mu\text{m}$		4.5 $\mu\text{m}$	5.7 $\mu\text{m}$	8 $\mu\text{m}$	24 $\mu\text{m}$	REDSHIFT	TYPE <sup>c</sup>		
S.....	34	0.011	3.58E–15	7.03E–15	–0.42	24.89	26.55	30.76	33.73	–1.00	1.037	S	B	
S.....	39	0.007	2.77E–15	5.62E–15	–0.41	8.24	8.47	9.55	14.32	–1.00	2.726	S	B	
S.....	41	0.004	4.35E–15	8.61E–15	–0.42	52.97	60.26	59.70	53.95	–1.00	1.616	S	B	
S.....	42	0.032	2.04E–16	6.74E–16	–0.24	5.06	6.49	9.46	5.81	–1.00	9.999	N	U	
S.....	43	0.015	6.84E–16	8.99E–16	–0.54	7.18	7.94	6.25	6.73	–1.00	2.573	S	N	

NOTE.—Table 1 is published in its entirety in the electronic edition of the *Astrophysical Journal*. A portion is shown here for guidance regarding its form and content.

<sup>a</sup> Probability of random association assuming a Poissonian distribution of sources. The expression  $1 - \exp(-n\pi d^2)$  was used, where  $n$  is the density of sources and  $d$  is the distance between the X-ray and infrared sources.

<sup>b</sup> Defined as  $(H - S)/(H + S)$ , where  $H$  and  $S$  are the background-subtracted counts in the hard and soft X-ray bands, respectively.

<sup>c</sup> S: Spectroscopic Redshift; P: Photometric Redshift; N: No redshift measured.

<sup>d</sup> B: Broad-line AGN; N: Narrow-line AGN; G: Galaxy; U: Unknown.

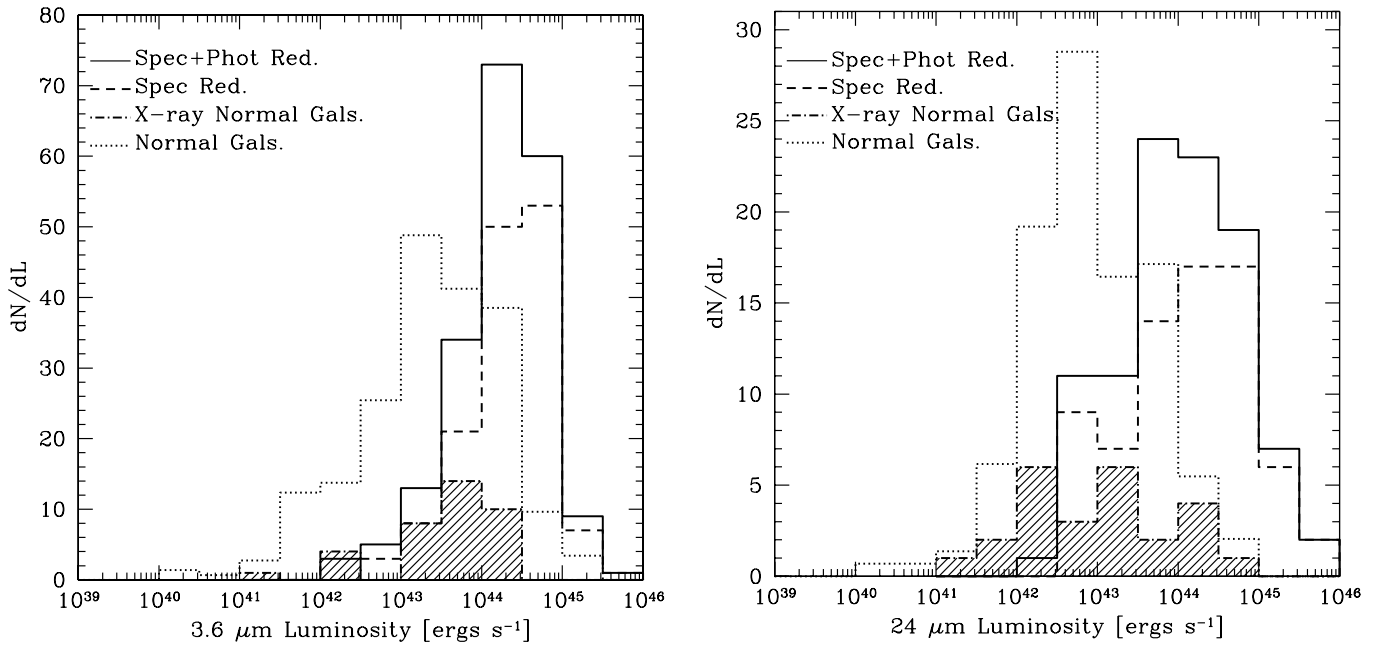


FIG. 2.—Infrared luminosity distribution for AGNs with spectroscopic (*dashed lines*) and both spectroscopic and photometric redshifts (*solid lines*) in the IRAC 3.6  $\mu\text{m}$  (*left*) and MIPS 24  $\mu\text{m}$  (*right*) bands. Most of the sources in our sample can be classified as moderate-luminosity AGNs, in the  $10^{42}$ – $10^{44}$   $\text{ergs s}^{-1}$  range. *Dotted lines*: Luminosity distribution for random GOODS field galaxies with either photometric or spectroscopic redshifts selected to have the same redshift distribution as the AGN sample. Sources classified as AGNs in X-rays are  $\sim 10$ – $15$  times brighter in the infrared than normal galaxies. *Dot-dashed/hatched histograms*: Luminosity distribution for X-ray sources with  $L_X < 10^{42}$   $\text{ergs s}^{-1}$ , and thus not included in our AGN sample. These sources have distributions similar to normal galaxies and much fainter on average than the AGN luminosity distribution. Therefore, we conclude that the X-ray emission in these sources does not come from AGNs.

et al. (2004) and Cowie et al. (2004) in the North field and from the Visible Multiobject Spectrograph Very Large Telescope (VIMOS VLT) Deep Survey (Le Fèvre et al. 2004) in the South field. Comparing the median infrared luminosities of AGNs and normal galaxies, AGNs are brighter by a factor of  $\sim 10$  in the IRAC bands and  $\sim 15$  in the MIPS 24  $\mu\text{m}$  band. There is a fair amount of overlap between the two populations, indicating that for low-luminosity AGNs ( $\sim 10^{42}$ – $10^{43}$   $\text{ergs s}^{-1}$ ) the host galaxy can provide a significant contribution and perhaps even dominate the infrared emission. In addition, the luminosity distribution for sources with hard X-ray luminosity lower than  $10^{42}$   $\text{ergs s}^{-1}$  (and thus removed from our AGN sample) is also shown in Figure 2. These sources have a luminosity distribution similar to that of normal galaxies, and therefore we conclude that the X-ray emission in those sources is not of AGN origin and most likely from star formation activity.

In order to understand the relative contributions of infrared emission from active nucleus and galaxy, we performed a simple simulation in which we randomly paired a normal galaxy from the redshift-equivalent distribution (infrared emission from stars only) to each AGN in the sample (infrared emission from stars plus active nucleus). We then calculated the distribution of the ratio of galaxy to nuclear light, assuming that the assigned galaxy correctly represented the emission from the host galaxy. This process was repeated 150 times with different random pairings. If indeed these normal galaxies correctly represent the host galaxy emission, which is of course uncertain but which is certainly true at optical wavelengths (e.g., Urry et al. 1999; Dunlop et al. 2003), the ensuing distribution indicates the relative importance of the two components. The results of a set of these simulations are shown in Figure 3 for the 8 and 24  $\mu\text{m}$  bands. We find that at 8  $\mu\text{m}$ ,  $\sim 87\%$  of the sample is dominated by the AGN (i.e., the luminosity of the active nucleus is greater than the galaxy luminosity). Separating the sample in three luminosity bins, each

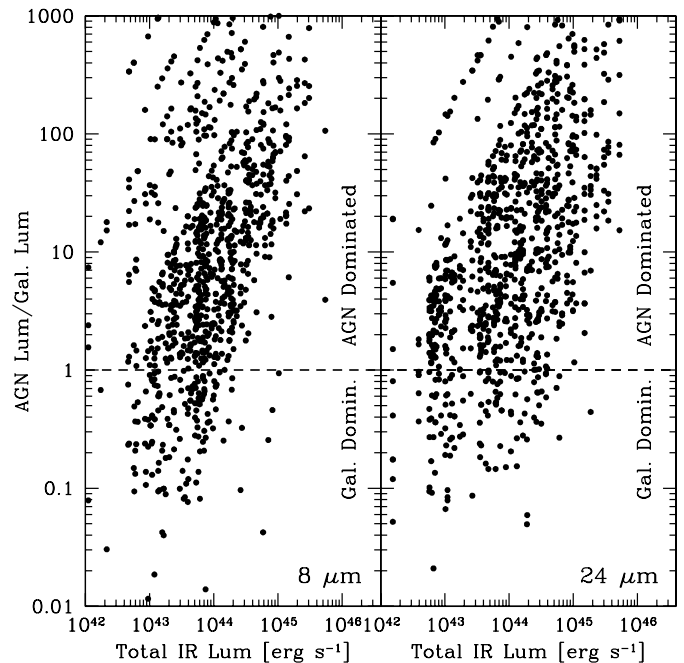


FIG. 3.—Simulated fraction of AGN to host galaxy infrared luminosity as a function of total infrared luminosity in the 8  $\mu\text{m}$  (*left*) and 24  $\mu\text{m}$  (*right*) bands. The AGN and host galaxy emission are separated by associating the luminosity of random galaxies in the field to AGNs in our sample. Here the results for five random associations for each source are shown. As can be seen, most of the sources are AGN dominated in the infrared with the fraction increasing with increasing total luminosity. This result does not depend much on the observed wavelength.

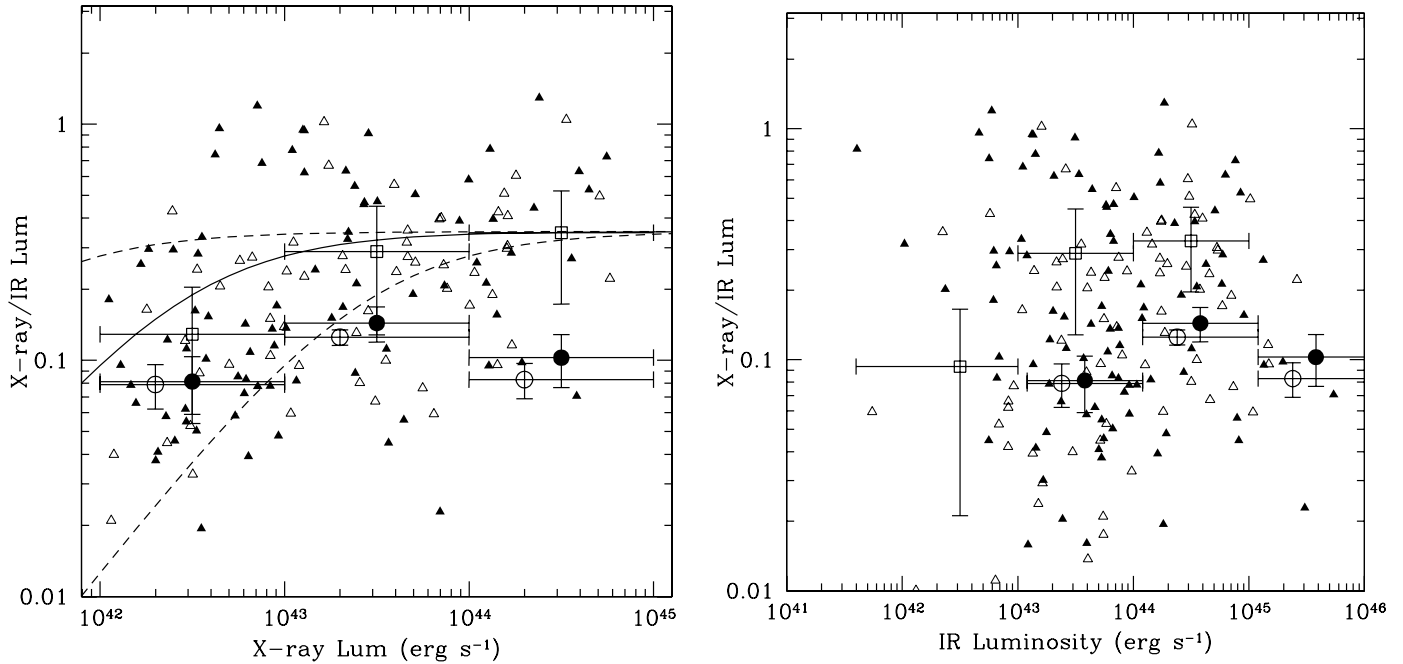


FIG. 4.—*Left*: Hard X-ray (2–8 keV) to infrared ( $8\ \mu\text{m}$ ) luminosity ratio as a function of hard X-ray luminosity for obscured (*filled triangles*) and unobscured (*open triangles*) AGNs (using  $N_{\text{H}} = 10^{22}\ \text{cm}^{-2}$  as the separating point) with measured spectroscopic redshifts in the GOODS fields. Average values (excluding the top and bottom 10%) in three X-ray luminosity bins are shown by the open squares, with error bars corresponding to the rms dispersion. For comparison, the average values for this ratio obtained from a compilation of local AGNs by Silva et al. (2004) are also shown (*filled circles*: obscured AGNs; *open circles*: unobscured AGNs). GOODS AGNs, the majority of them at  $z \sim 0.5\text{--}1$  (higher redshift AGNs are too faint for spectroscopy), and local sources have similar values for this ratio, possibly diverging at higher luminosity. *Solid line*: X-ray-to-infrared luminosity ratio assuming an intrinsic ratio of 0.35 and a host galaxy with a luminosity of  $7.67 \times 10^{42}\ \text{ergs s}^{-1}$  (median luminosity for our random sample of GOODS normal galaxies). The dashed lines show the effects of changing the luminosity of the host galaxy by 1 order of magnitude (*lower line*: brighter host galaxy). The small observed trend with luminosity can be explained by the relatively more important contribution from the host galaxy for low-luminosity AGNs. *Right*: Hard X-ray-to-infrared luminosity ratio as a function of infrared luminosity. Symbols are the same as in the left panel. In this case, sources below the X-ray luminosity threshold are also included in the diagram. The same conclusions are still valid when the X-ray-to-infrared luminosity ratio is measured as a function of infrared or hard X-ray luminosity.

with equal numbers of AGNs, we find that for total AGN luminosities below  $4 \times 10^{43}\ \text{ergs s}^{-1}$ , roughly  $\frac{3}{4}$  of the sample is AGN dominated. For  $4 \times 10^{43}\ \text{ergs s}^{-1} < L_{\text{IR}} < 1.6 \times 10^{44}\ \text{ergs s}^{-1}$ , the fraction of AGN-dominated sources is roughly 90%, and for higher luminosities it rises to 97%. At shorter wavelengths and at  $24\ \mu\text{m}$ , the fractions are similar. At  $24\ \mu\text{m}$ ,  $\sim 80\%$  of the sources are AGN dominated in the lowest luminosity bin and  $\sim 90\%$  at the highest luminosities. Of course, because of the limited spatial resolution, we cannot say from the infrared data alone whether the AGN host galaxies are similar to normal galaxies in the same redshift range. Given the similarity of the optical properties, however, it would require a fairly contrived scenario to have the infrared properties be massively mismatched. The correlation of infrared and X-ray flux (next section) also supports the idea that the infrared emission from AGNs is dominated by the active nucleus.

### 3.2. X-Ray-to-Infrared Ratios

In Figure 4 we show the ratio of hard X-ray (2–8 keV) to infrared ( $8\ \mu\text{m}$ ) luminosity versus hard X-ray and infrared luminosity for GOODS AGNs with spectroscopic redshifts and thus reliable luminosities. Using the X-ray spectrum to separate obscured and unobscured AGNs (the separation was set at  $N_{\text{H}} = 10^{22}\ \text{cm}^{-2}$ , with  $N_{\text{H}}$  inferred from the X-ray hardness ratio, assuming an intrinsic photon slope of  $\Gamma = 1.7$ ), we find no differences between the two classes of AGNs (Fig. 4, *filled and open triangles*). This is consistent with the AGN unification paradigm, since both hard X-ray and infrared emission are *relatively* isotropic and unaffected by the amount of dust in the line of sight. Similar results are found if the  $24\ \mu\text{m}$  fluxes are used

instead; however, these fluxes are only available for half of our sample, and thus we decided to use the  $8\ \mu\text{m}$  fluxes.

In order to compare with previously known sources, in Figure 4 we also show the average observed-frame hard X-ray-to-infrared luminosity ratio for both local Seyfert 1 and 2 galaxies in three X-ray luminosity bins, from the compilation of Silva et al. (2004). The average ratio for the local sources is  $0.10 \pm 0.02$  (standard deviation for the sample, including  $K$ -correction), considering both obscured and unobscured AGNs and excluding the top and bottom 10% values of the distribution. For the GOODS sample this ratio is  $0.18 \pm 0.12$ , also excluding the top and bottom 10% values, and there is also no significant difference in the average values for the obscured and unobscured AGNs. The average value for the GOODS sample in three X-ray luminosity bins is shown by the open squares in Figure 4. There is no statistically significant discrepancy between the GOODS and local samples, although there is a hint that the values might diverge at high X-ray luminosity ( $< 2\ \sigma$  significance). However, it should be noted that Malmquist bias, which would make the average luminosity appear to be higher at higher redshifts, can play a role in this difference.

The trend for the average ratio of X-ray-to-infrared luminosity to increase with X-ray luminosity in the GOODS sample can be explained by the contribution of the host galaxy to the infrared light at low luminosities. Assuming that this ratio has a value of 0.35 for the highest luminosity sources, which are clearly AGN dominated, and adding to the infrared luminosity the contribution from an average host galaxy (estimated from the median  $8\ \mu\text{m}$  luminosity for our random sample of galaxies,  $7.67 \times 10^{42}\ \text{ergs s}^{-1}$ ), we were able to reproduce the observed trend of X-ray-to-infrared luminosity ratio with X-ray luminosity (Fig. 4, *solid line*).

The spread in the values of the X-ray-to-infrared luminosity for the GOODS sample can be explained by the combined effects of ignoring the  $K$ -correction and the intrinsic spread in the luminosity of the host galaxy. Using the AGN spectrum library of Treister et al. (2004), we estimate that the  $K$ -correction can account for variations of  $\sim 50\%$  around the mean observed ratio. The dashed lines in Figure 4 show the effects of a spread in host galaxy luminosity, which is particularly important for low-luminosity AGNs and can account for most of the observed spread in the X-ray-to-infrared luminosity ratio at these luminosities.

Performing a partial correlation analysis, we found no significant correlation between the hard X-ray-to-infrared luminosity ratio and either hard X-ray or infrared luminosity, or with redshift or absorbing column density ( $N_{\text{H}}$ ). We found only a mildly significant correlation between hard X-ray and infrared luminosities, with a Pearson coefficient of  $\sim 0.604$  and a correlation of similar significance between hard X-ray and infrared fluxes, indicating that the fluxes in these bands are directly related to the luminosity of the central engine. *The lack of correlation of the ratio with redshift or luminosity suggests that there is little dependence of the emission mechanism on either redshift or luminosity.*

### 3.3. Infrared Counts of AGNs

The cumulative distributions of infrared fluxes in the 3.6, 5.7, 8, and 24  $\mu\text{m}$  bands for the GOODS AGNs are presented in Figure 5. The observed counts were corrected to account for the differences in covered area depending on the X-ray flux (since the sample is X-ray selected) based on the relation given by Figure 5 in Treister et al. (2004), obtained by adapting the sensitivity versus area function for the Chandra Deep Fields of Alexander et al. (2003) to the portion covered by GOODS (D. Alexander & F. Bauer 2004, private communication). No correction for incompleteness in the X-ray or infrared catalogs is necessary, as these are almost complete in the flux ranges sampled in the number counts determination. Error bars in these distributions were calculated following the prescription of Gehrels (1986) for 84% confidence levels on the number of sources. The number counts in the 3.6, 4.5, and 5.7  $\mu\text{m}$  bands can be well described by a power law with integral slope  $1.5 \pm 0.2$  for fluxes higher than 50  $\mu\text{Jy}$  from a maximum likelihood analysis of the unbinned data (Crawford et al. 1970). In the 8  $\mu\text{m}$  band the integral slope is slightly shallower,  $1.2 \pm 0.2$ , at fluxes higher than 70  $\mu\text{Jy}$ , with a density of  $\simeq 1000$  AGNs per square degree at a depth of  $\sim 35 \mu\text{Jy}$ . In the MIPS 24  $\mu\text{m}$  band the slope is shallower still,  $0.9 \pm 0.2$ , at all the fluxes sampled by the GOODS sources, and there are 1000 AGNs per square degree at a flux density of  $\sim 200 \mu\text{Jy}$ . The resulting best fits are shown in Figure 5. In order to test the goodness of these fits we performed a K-S test comparing the observed distribution to the fitting function, finding that the null hypothesis (that both distributions are drawn from the same parent distribution) cannot be rejected at confidence levels higher than  $\sim 40\%$  in all the bands. The values of the K-S probability for each band are 45%, 50%, 75%, 85%, and 40% on the 3.6, 4.5, 5.7, 8, and 24  $\mu\text{m}$  bands, respectively.

In the unification paradigm for AGNs, most of the energy absorbed by the circumnuclear optically thick matter is reemitted at infrared wavelengths. It is not surprising, then, that AGNs are luminous infrared sources, even when the optical, ultraviolet, and soft X-ray emission is faint. Population synthesis models for the X-ray background suggest that a substantial fraction of AGNs are obscured along the line of sight, and thus will be bright only at

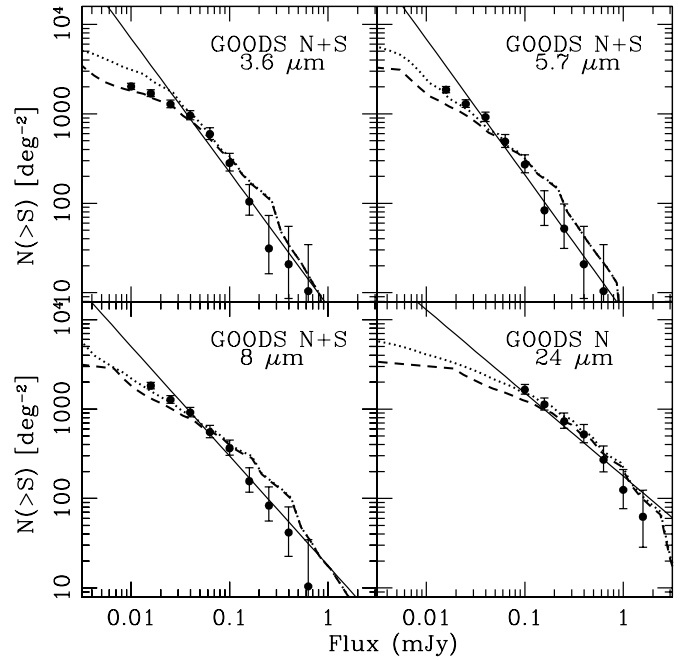


Fig. 5.—Filled circles: Observed cumulative infrared flux distributions for X-ray-detected AGNs in the GOODS-North and South fields in three IRAC bands (3.6–8  $\mu\text{m}$ ) and for the North field only in the MIPS 24  $\mu\text{m}$  band. Error bars correspond to the 84% confidence level in the number of sources on that bin (Gehrels 1986). Note that because this is a cumulative plot, the bins are not independent. Solid line: Power-law fit to the observed number counts. Dashed line: Expected infrared flux distribution based on the models of Treister et al. (2004), taking into account the X-ray flux limit in the GOODS fields and correcting for the effects of undersampling of the bright end caused by the small volume of the GOODS fields. Dotted line: Expected infrared flux distribution not considering the X-ray flux limit (i.e., all sources expected from extrapolating the AGN luminosity function). In general, the model follows the normalization and shape of the counts reasonably well, especially considering the poor statistics at the bright end. At the faint end, where the host galaxy light dominates, the model diverges more significantly, since the model assumptions are clearly far too simple.

mid- to far-infrared and hard X-ray wavelengths. For a simple unification model, in which the material is distributed in a torus and the ratio of obscured to unobscured AGNs is set to  $\sim 3:1$ , as observed in the local universe (Risaliti et al. 1999), we can calculate the expected infrared counts. Model details are as described in Treister et al. (2004) and modified by Treister & Urry (2005) to include a decrease in the number of obscured AGNs with increasing luminosity, with the exception that we no longer interpolate between the near-infrared and optical spectra because this artificially overpredicts the mid-infrared flux.

The obscured-to-unobscured AGN ratio for sources with spectroscopic redshifts (and thus  $N_{\text{H}}$  determinations) in our sample is  $\sim 1.5:1$ , but this is consistent with the assumed ratio in the model, given that most of the obscured AGNs do not have spectroscopic redshift, since they are optically faint, and some obscured AGNs are even missed in X-rays. Therefore, the observed value of the ratio only represents a lower limit to the intrinsic ratio, which can easily be higher by a factor of 2 and therefore consistent with the average value assumed in our model.

This simple AGN unification model agrees within a factor of 2 with the observed distribution in the GOODS fields over the IRAC 3.6–8  $\mu\text{m}$  and MIPS 24  $\mu\text{m}$  bands, as shown in Figure 5. Compared to the model, which assumes only an elliptical host galaxy, with a single ( $L_*$ ) luminosity and no star formation, and fixed parameters for the dust torus, there are slightly fewer observed AGNs at the bright end ( $>100 \mu\text{Jy}$ ). The small area of the

GOODS survey causes bright AGNs to be poorly sampled, hence the large error bars and the marginal significance of the discrepancy. In addition, it is important to point out that in a cumulative plot the bins are not independent. Wide-area surveys such as the *Spitzer* Wide-Area Infrared Extragalactic Survey (SWIRE; Lonsdale et al. 2003) and the MIPS GTO<sup>11</sup> programs will provide a better estimate of the AGN counts at the bright end. At the faint end ( $<10 \mu\text{Jy}$ ), the observed number of AGNs is higher than the model predictions except in the  $3.6 \mu\text{m}$  band, where there is very good agreement. These discrepancies are  $\sim 13\%$  at  $4 \mu\text{m}$  and rising to  $36\%$  at  $8 \mu\text{m}$  and  $\sim 30\%$  at  $24 \mu\text{m}$ . These discrepancies may suggest that for fainter AGNs, the host galaxy plays a relatively more important role in some AGNs, thus making them brighter. At  $3.6 \mu\text{m}$  the stellar component in the host galaxy dominates; the good agreement with our simple model means that the host galaxy must be reasonably well represented by our assumed  $L_*$  elliptical galaxy with no recent star formation.

### 3.4. Extragalactic Infrared Background Light

Using the *Spitzer* data we can estimate the integrated contribution of AGNs to the extragalactic infrared background light at IRAC and MIPS wavelengths. A representative lower limit can be obtained simply by summing the infrared emission from the AGNs detected in the GOODS fields, weighted by the effective area at that X-ray flux. In Figure 6, we show the resulting contribution measured from the GOODS sources, compared to the observational limits for the total extragalactic infrared background compiled by Hauser & Dwek (2001; the lower limit is from the integrated extragalactic light, while the upper limit comes from fluctuation analysis) and updated measurements at the IRAC and MIPS  $24 \mu\text{m}$  wavelengths (Matsumoto et al. 2005; Fazio et al. 2004b; Kashlinsky et al. 1996; Stanev & Franceschini 1998; Papovich et al. 2004). AGNs and their host galaxies contribute approximately  $2.4_{-0.7}^{+1.7}\%$ ,  $3.0_{-1.3}^{+5.7}\%$ ,  $5.6_{-1.9}^{+5.3}\%$ , and  $4.7_{-1.7}^{+5.4}\%$  in the IRAC 3.6, 4.5, 5.7, and  $8 \mu\text{m}$  bands, respectively, and  $10.7_{-4.5}^{+11.0}\%$  in the MIPS  $24 \mu\text{m}$  band. (The uncertainties are dominated by the range in the total infrared background light.) These are lower limits to the total AGN contribution, since they exclude the contribution from faint sources below the GOODS X-ray detection limit (including a large expected population of Compton-thick AGNs; Treister & Urry 2005) and from bright sources that are not well sampled in the relatively small GOODS fields. Using our AGN unification model, we can calculate the relative contribution of these X-ray-faint sources and add it to the total emission from X-ray-detected sources, in order to estimate the total contribution of AGNs to the infrared background. We find that sources missed in X-rays can add  $\sim 45\%$  more flux in the  $3\text{--}24 \mu\text{m}$  range, with  $\sim 30\%$  extra flux coming from obscured AGNs that are too faint to be detected in X-rays and the remaining  $\sim 15\%$  coming from bright sources that are not well sampled in the GOODS fields. Correcting for this factor, the AGN contribution is increased to  $\sim 3\%$ – $8\%$  in the IRAC bands and  $\sim 15\%$  in the MIPS  $24 \mu\text{m}$  band. Given that these values are obtained from the summed emission of AGNs and their host galaxies, these are upper limits to the AGN contribution to infrared background or, equivalently, can be seen as the contribution to the infrared background by AGNs plus their host galaxies.

We also investigate the flux dependence of the AGN contribution to the integrated infrared emission. This was computed by comparing the total flux from AGNs with the integrated flux from all sources for several flux bins. At  $3.6 \mu\text{m}$  the AGN

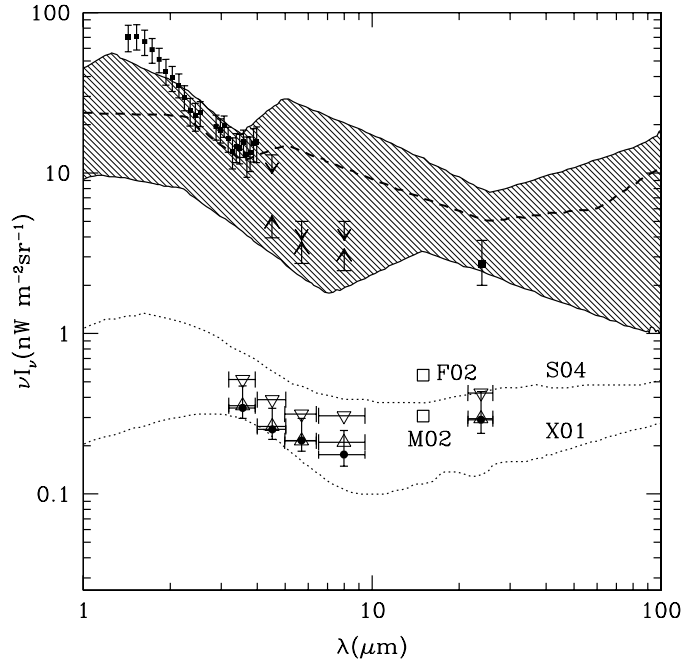


FIG. 6.—Extragalactic infrared background intensity as a function of wavelength. *Shaded region*: Allowed values compiled by Hauser & Dwek (2001).  $1\text{--}4 \mu\text{m}$ : Measurements reported by Matsumoto et al. (2005).  $4\text{--}8 \mu\text{m}$ : Lower limits from Fazio et al. (2004b), upper limits from Kashlinsky et al. (1996) at  $4.5 \mu\text{m}$  and Stanev & Franceschini (1998) at  $5.8$  and  $8 \mu\text{m}$ . Measurement at  $24 \mu\text{m}$  from Papovich et al. (2004). *Solid circles*: Integrated infrared emission from X-ray-detected AGNs in the GOODS fields. Error bars in these measurements are estimated based on the number of sources. *Triangles*: Integrated AGN emission from the models of Treister et al. (2004) and Treister & Urry (2005) if only X-ray-detected AGNs are considered (*lower*) and including all AGNs (*upper*). *Dotted lines*: Expected AGN integrated emission from the AGN models of Silva et al. (2004; *upper line*) and Xu et al. (2001; *lower line*). *Open squares*: Expected AGN integrated emission at  $15 \mu\text{m}$  from the extrapolation to fainter fluxes of IRAS and ISO (Matute et al. 2002 [M02]; corrected by a factor of 5 to include the contribution from obscured AGNs) and ISO only (Fadda et al. 2002 [F02]) observations.

contribution is  $2.1\%$  above  $600 \mu\text{Jy}$ , rising to  $\sim 4\%$  above  $10 \mu\text{Jy}$ . At  $8 \mu\text{m}$ , the AGN contribution is  $\sim 6\%$  and remains mostly constant, with a slight increase to  $\sim 10\%$  at  $40 \mu\text{Jy}$ . Only in the MIPS  $24 \mu\text{m}$  band do we see a decrease in the AGN contribution with decreasing flux, going from  $30\%$  above  $1000 \mu\text{Jy}$  to  $\sim 8\%$  above  $100 \mu\text{Jy}$ , as shown in Figure 7. In other words, at lower fluxes the number of galaxies rises much more steeply than the number of AGNs. While this can also mean that many faint AGNs are missed by the X-ray selection, in Figure 7 we use the previously described AGN model to estimate the AGN fraction as a function of infrared flux, correcting for AGNs missed by the X-ray selection. As can be seen in Figure 7, the decrease in the AGN fraction with decreasing infrared flux is still observed, even after correcting for AGNs missed in X-rays.

Previous reports of the AGN contribution in general estimated a higher contribution to the integrated emission. Earlier infrared observations with ISO (e.g., Fadda et al. 2002) in the  $15 \mu\text{m}$  band at brighter fluxes give higher values for the AGN contribution ( $17\%$  at  $15 \mu\text{m}$ ) but depend on large extrapolations to lower fluxes. (The GOODS MIPS observations are  $\sim 25$  times deeper than the ISO data on the Lockman Hole.) Matute et al. (2002) made an estimate based on ISO observations of bright unobscured AGNs in the European Large-Area ISO Survey (ELAIS) S1 field. If, in order to calculate the total AGN contribution from their measured value for unobscured AGNs, we assume an obscured-to-unobscured ratio of  $\sim 4 : 1$ , their results are consistent with the observed values obtained from the GOODS

<sup>11</sup> More information can be found at <http://lully.as.arizona.edu/>.

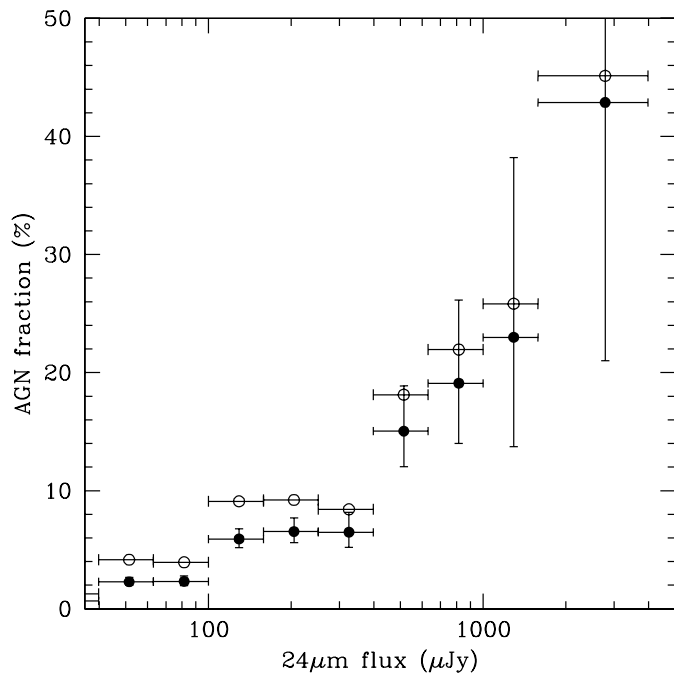


FIG. 7.—Fraction of sources classified as AGNs increases with infrared flux in the  $24\ \mu\text{m}$  band (filled circles). Horizontal error bars show the size of the flux bins, while vertical error bars show the  $1\ \sigma$  Poissonian errors on the number of sources. In order to test whether this effect can be caused by the X-ray selection, we also show the contribution corrected by the AGNs expected to be missed by X-ray selection, as estimated using the AGN population synthesis model (open circles). As can be seen, this effect is still observed after correcting for the AGNs not detected in X-rays.

observations. An estimate based on X-ray samples and an average X-ray-to-infrared ratio by Silva et al. (2004) is  $\sim 50\%$ – $80\%$  higher than the values obtained directly from the *Spitzer* observations. Using a backward evolution model based on *IRAS* and *ISO* observations and infrared luminosity functions, the AGN contribution to the infrared background was also calculated by Xu et al. (2001). In this case only an infrared color criterion was used to identify AGNs. This prediction is similar to the observed values in the  $3.6$  and  $4.5\ \mu\text{m}$  bands but estimates a lower contribution at longer wavelengths,  $\sim 55\%$  lower at  $24\ \mu\text{m}$ . A recent study of AGNs detected by *Chandra* using *Spitzer* performed by Franceschini et al. (2005) reported a value for the AGN contribution to the infrared background of  $10\%$ – $15\%$  at  $24\ \mu\text{m}$  in good agreement with our results.

The earlier estimates of the AGN contribution to the integrated extragalactic infrared light are therefore largely consistent with

our results; the earlier numbers differ mainly because of large extrapolations to lower fluxes. In summary, the more precise values from our deep *Chandra* and *Spitzer* observations are  $2.4\%$  at  $3.6\ \mu\text{m}$ ,  $3.0\%$  at  $4.5\ \mu\text{m}$ ,  $5.6\%$  at  $5.7\ \mu\text{m}$ ,  $4.7\%$  at  $8\ \mu\text{m}$ , and  $10.7\%$  at  $24\ \mu\text{m}$ . These numbers should be multiplied by a factor of  $\sim 1.45$  in order to account for the contribution by Compton-thick AGNs not detected in X-rays by *Chandra*, namely,  $3.5\%$  at  $3.6\ \mu\text{m}$ ,  $4.4\%$  at  $4.5\ \mu\text{m}$ ,  $8.1\%$  at  $5.7\ \mu\text{m}$ ,  $6.8\%$  at  $8\ \mu\text{m}$ , and  $15.5\%$  at  $24\ \mu\text{m}$ .

#### 4. CONCLUSIONS

We here present *Spitzer* observations of the moderate-luminosity, moderate- to high-redshift AGNs in the GOODS fields, which include the deepest infrared and X-ray observations to date. In general, these sources have X-ray-to-infrared luminosity ratios similar to those of AGNs in the local universe. AGNs are  $\sim 10$ – $15$  times brighter in the infrared than normal galaxies, although there is some overlap between the luminosity distribution of AGNs and normal galaxies, indicating that for low-luminosity AGNs the contribution to the infrared light from the host galaxy can be very important and may well dominate the observed emission. This can explain the observed trend of the X-ray-to-infrared luminosity ratio to increase with X-ray luminosity.

We report the observed infrared flux distribution for GOODS AGNs, which has a slope of  $\simeq 1.5$ – $0.9$  for fluxes higher than  $\sim 100\ \mu\text{Jy}$ , with a source density of  $\simeq 1000$  AGNs per square degree at a depth of  $\sim 50\ \mu\text{Jy}$  in the IRAC bands. Comparing to the expectations of a unification model for AGNs, we find reasonable agreement, showing that the basic properties of the moderate- to high-redshift AGN population are well explained by the same unification paradigm that fits local sources very well, but with a factor of 2 less dust in the circumnuclear region.

From the *Spitzer* observations, we calculate the minimum AGN contribution to the extragalactic near-mid-infrared background, obtaining a lower value than previously estimated, ranging from  $2\%$  to  $10\%$  in the  $3$ – $24\ \mu\text{m}$  range. Accounting for heavily obscured AGNs that are not detected in X-rays, the AGN contribution to the infrared background increases by  $\sim 45\%$ , to  $\sim 3\%$ – $15\%$ .

We thank the GOODS team for very useful comments and the anonymous referee for suggestions that improved the paper. E. T. acknowledges the support of Fundación Andes, Centro de Astrofísica FONDAF, MECESUP UCH-118, and the Sigma Xi Foundation through a grant-in-aid for research. This work was supported in part by NASA grant HST-GO-09425.13-A.

#### REFERENCES

- Alexander, D. M., et al. 2003, *AJ*, 126, 539  
 Barger, A. J., et al. 2003, *AJ*, 126, 632  
 Bauer, F. E., Alexander, D. M., Brandt, W. N., Schneider, D. P., Treister, E., Hornschemeier, A. E., & Garmire, G. P. 2004, *AJ*, 128, 2048  
 Bauer, F. E., Condon, J. J., Thuan, T. X., & Broderick, J. J. 2000, *ApJS*, 129, 547  
 Bertin, E., & Arnouts, S. 1996, *A&AS*, 117, 393  
 Brandt, W. N., et al. 2001, *AJ*, 122, 2810  
 Cowie, L. L., Barger, A. J., Hu, E. M., Capak, P., & Songaila, A. 2004, *AJ*, 127, 3137  
 Crawford, D. F., Jauncey, D. L., & Murdoch, H. S. 1970, *ApJ*, 162, 405  
 Dickinson, M., & Giavalisco, M. 2003, in *The Mass of Galaxies at Low and High Redshift*, ed. R. Bender & A. Renzini (Berlin: Springer), 324  
 Dunlop, J. S., McLure, R. J., Kukula, M. J., Baum, S. A., O’Dea, C. P., & Hughes, D. H. 2003, *MNRAS*, 340, 1095  
 Fabian, A. C., & Iwasawa, K. 1999, *MNRAS*, 303, L34  
 Fadda, D., Flores, H., Hasinger, G., Franceschini, A., Altieri, B., Cesarsky, C. J., Elbaz, D., & Ferrando, P. 2002, *A&A*, 383, 838  
 Fazio, G. G., et al. 2004a, *ApJS*, 154, 10  
 ———. 2004b, *ApJS*, 154, 39  
 Franceschini, A., Fadda, D., Cesarsky, C. J., Elbaz, D., Flores, H., & Granato, G. L. 2002, *ApJ*, 568, 470  
 Franceschini, A., et al. 2005, *AJ*, 129, 2074  
 Gehrels, N. 1986, *ApJ*, 303, 336  
 Giacconi, R., et al. 2001, *ApJ*, 551, 624  
 Giavalisco, M., et al. 2004, *ApJ*, 600, L93  
 Gilli, R., Salvati, M., & Hasinger, G. 2001, *A&A*, 366, 407  
 Hauser, M. G., & Dwek, E. 2001, *ARA&A*, 39, 249  
 Kashlinsky, A., Mather, J. C., & Ohenwald, S. 1996, *ApJ*, 473, L9  
 Le Fèvre, O., et al. 2004, *A&A*, 428, 1043  
 Lira, P., Ward, M., Zezas, A., Alonso-Herrero, A., & Ueno, S. 2002, *MNRAS*, 330, 259  
 Lonsdale, C. J., et al. 2003, *PASP*, 115, 897  
 Martínez-Sansigre, A., Rawlings, S., Lacy, M., Fadda, D., Marleau, F. R., Simpson, C., Willott, C. J., & Jarvis, M. J. 2005, *Nature*, 436, 666



- Matsumoto, T., et al. 2005, ApJ, 626, 31  
Matute, I., et al. 2002, MNRAS, 332, L11  
Mobasher, B., et al. 2004, ApJ, 600, L167  
Papovich, C., et al. 2004, ApJS, 154, 70  
Rieke, G. H., et al. 2004, ApJS, 154, 25  
Risaliti, G., Maiolino, R., & Salvati, M. 1999, ApJ, 522, 157  
Rush, B., Malkan, M. A., & Spinoglio, L. 1993, ApJS, 89, 1  
Sajina, A., Lacy, M., & Scott, D. 2005, ApJ, 621, 256  
Setti, G., & Woltjer, L. 1989, A&A, 224, L21  
Silva, L., Maiolino, R., & Granato, G. L. 2004, MNRAS, 355, 973  
Stanev, T., & Franceschini, A. 1998, ApJ, 494, L159  
Szokoly, G. P., et al. 2004, ApJS, 155, 271  
Treister, E., & Urry, C. M. 2005, ApJ, 630, 115  
Treister, E., et al. 2004, ApJ, 616, 123  
Urry, C. M., Falomo, R., Scarpa, R., Pesce, J. E., Treves, A., & Giavalisco, M. 1999, ApJ, 512, 88  
Wirth, G. D., et al. 2004, AJ, 127, 3121  
Xu, C., Lonsdale, C. J., Shupe, D. L., O'Linger, J., & Masci, F. 2001, ApJ, 562, 179