

# Young stars in the periphery of the Large Magellanic Cloud

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## ABSTRACT

Despite their close proximity, the complex interplay between the two Magellanic Clouds, the Milky Way and the resulting tidal features, is still poorly understood. Recent studies have shown that the Large Magellanic Cloud (LMC) has a very extended disc strikingly perturbed in its outskirts. We search for recent star formation in the far outskirts of the LMC, out to  $\sim 30^\circ$  from its centre. We have collected intermediate-resolution spectra of 31 young star candidates in the periphery of the LMC and measured their radial velocity, stellar parameters, distance and age. Our measurements confirm membership to the LMC of six targets, for which the radial velocity and distance values match well with those of the Cloud. These objects are all young (10–50 Myr), main-sequence stars, projected between  $7^\circ$  and  $13^\circ$  from the centre of the parent galaxy. We compare the velocities of our stars with those of a disc model, and find that our stars have low to moderate velocity differences with the disc model predictions, indicating that they were formed *in situ*. Our study demonstrates that recent star formation occurred in the far periphery of the LMC, where thus far only old objects were known. The spatial configuration of these newly formed stars appears ring-like with a radius of 12 kpc and a displacement of 2.6 kpc from the LMC's centre. This structure, if real, would be suggestive of a star formation episode triggered by an off-centre collision between the Small Magellanic Cloud and the LMC's disc.

**Key words:** stars: early-type – Magellanic Clouds.

## 1 INTRODUCTION

The Large and Small Magellanic Clouds (LMC and SMC, respectively) form a binary pair of gas-rich, star-forming galaxies that are satellites of the Milky Way, located within  $\sim 50$  kpc of their host. This is a configuration that is rarely found around Milky Way type galaxies (James & Ivory 2011; Liu et al. 2011). Interactions between the Clouds themselves and with the Galaxy have generated a complex ensemble of gaseous structures such as: the extended Magellanic Stream (Mathewson, Cleary & Murray 1974), encompassing  $\sim 200^\circ$  across the sky along the Clouds' orbit (Nidever et al. 2010); the Bridge between the Clouds (Kerr 1957; Hindman, Kerr & McGee 1963); and the Leading Arm (LA), a

fragmented, wide feature leading the orbit of the Clouds (Gardiner & Noguchi 1996, see their fig. 5). Observational data on the gas in this system have accumulated over the past decades and it is now realized that the Magellanic Stream, Bridge and LA bring to the halo of the Milky Way a mass of gas on the order of  $2 \times 10^9 M_\odot$ , or over twice as much gas as the Clouds themselves contain (see review by D'Onghia & Fox 2016). Are there any stars in these gaseous features? Tidal models predict the existence of stars; however, observations have not been decisive primarily due to the lack of deep, wide and contiguous photometric surveys. While young, recently formed stars have been found in the Bridge (Irwin, Demers & Kunkel 1990; Demers & Battinelli 1998; Harris 2007; Casetti-Dinescu et al. 2012; Skowron et al. 2014) and, more recently, in the LA (Casetti-Dinescu et al. 2014), a substantial, intermediate-age population is still elusive, with some hints of its presence in the Bridge (Bagheri, Cioni & Napiwotzki 2013; Noël et al. 2015;

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Saha & Olszewski 2015). This status is however improving, primarily due to the Dark Energy Survey (DES; Diehl et al. 2014), which has recently revealed a large, complex stellar extent to the LMC (Belokurov & Koposov 2016; Mackey et al. 2016) as well as new dwarf galaxies, some of which may be associated with the Magellanic system (see review by D’Onghia & Fox 2016).

In this contribution, we focus on the outskirts of the LMC. Earlier studies based on near-IR surveys agree that the LMC disc extends to some 9 kpc in radius (or about  $8^\circ$  on the sky), and is made of an intermediate to old stellar population. However, Muñoz et al. (2006) report the finding of a handful of giant stars, likely members of the LMC, in the field of Carina dwarf galaxy some  $20^\circ$  away from the LMC’s centre. This finding was later confirmed by McMonigal et al. (2014). Saha et al. (2010) present a photometric study in a series of fields that probes the outer regions of the LMC. In particular, the northward fields from a distance of  $7^\circ$ – $20^\circ$  from the LMC’s centre show an exponential decline in the number of stars; this indicates that the intermediate-age/old disc extends out to  $16^\circ$  from the LMC’s centre (see Saha et al. 2010, and references therein). Finally, the recent study by Mackey et al. (2016) finds in the same northern region of the LMC a 10-kpc tidal arm at some  $16^\circ$  from the LMC’s centre. Clearly, the disc of the LMC is very extended and perturbed in its outer parts.

Given that the disc extends so far out, and that it shows evidence of tidal interaction in its outskirts, we seek to know just how far from LMCs centre star formation can occur. This will help us better understand star formation in an environment that was subject to a recent ( $\sim 200$  Myr) interaction with the SMC, i.e. a turbulent and tidally stressed environment. To this end, we search for young, newly formed stars that are LMC members. Casetti-Dinescu et al. (2012, hereafter CD12) listed 567 OB star candidates in a wide area including the periphery of the Clouds, the Bridge, the Leading Arm and part of the Magellanic Stream. These stars were selected from cuts in magnitude, colours and proper motions, after merging the UV photometry from the *Galaxy Evolution Explorer* (GALEX; Bianchi et al. 2011), the infrared data of the Two Micron All-Sky Survey (2MASS; Skrutskie et al. 2006), the optical photometry of the AAVSO Photometric All-Sky Survey (APASS; Henden et al. 2011) and proper motions and photometry from the Southern Proper Motion program (SPM4; Girard et al. 2011). The selection criteria were adjusted to isolate distant, hot, young main-sequence (MS) stars. However, hot subdwarf stars (sdBs; Heber 1986) and white dwarfs (WDs) should inevitably be present in the sample, because they are photometrically indistinguishable from hot MS stars. These objects are intrinsically fainter than MS stars by more than five magnitudes, hence any such star in the sample will be a foreground Galactic object. Galactic runaway stars could also contaminate the sample to a lesser extent.

We started an extensive spectroscopic campaign to investigate the CD12 young star candidates in more detail (Moni Bidin et al. 2015). In Casetti-Dinescu et al. (2014), we identified some 19 young OB stars in the direction of the LA, where no stellar component was previously known; five of these have kinematics consistent with LA membership and are located at the edge of the Galactic disc. This discovery suggests that the interaction between the Clouds and our Galaxy is strong enough to trigger recent star formation in certain regions of this gaseous structure. However, our analysis also revealed that a relevant fraction (55 per cent) of the candidate sample consists of foreground, more compact Galactic objects. In this paper, we focus on the young star candidates found in the periphery of the LMC, i.e. out to a projected radius of about  $33^\circ$  from the LMC centre, while excluding the inner  $6^\circ$ .

## 2 OBSERVATIONS AND MEASUREMENTS

### 2.1 Observations and data reduction

We selected 31 young star candidates with  $V < 16$  between  $6^\circ$  and  $33^\circ$  from the centre of the LMC, from the catalogue of Casetti-Dinescu et al. (2012). The IDs, coordinates and  $V$  magnitudes of target stars, drawn from the SPM4 catalogue (Girard et al. 2011), are given in the first four columns of Table 1. Their intermediate-resolution spectra ( $R \approx 3500$ ) were collected during three half-nights of observations between 2015 January 28 and February 22 at Las Campanas Observatory, with the IMACS spectrograph at the focus of the Baade 6.5 m telescope. The instrument was used at  $f/4$  in longslit mode, and the 0.75 arcsec wide slit was employed. The 1200–17.5 grating was tilted by an angle of  $16:8$  to cover the spectral range 3660–5250 Å on the CCD. Exposure times varied between 150s and 1500s, according to target magnitude and sky conditions, to yield spectra with signal-to-noise ratio (S/N)  $> 50$ . A comparison lamp for wavelength calibration was acquired after each exposure. A twilight solar spectrum was collected at the beginning of each night, and the radial-velocity standards LTT2415 (Latham et al. 2002) and GJ273 (Chubak et al. 2012) were also observed.

The spectra were split on four chips with short gaps between them. Each spectral section was independently reduced, extracted with an optimum algorithm (Horne 1986) and normalized, with standard IRAF<sup>1</sup> routines (see e.g. de la Fuente Marcos et al. 2015). Before merging them, the absolute offsets of the four independent wavelength calibration solutions were measured on the standard stars and corrected. Some reduced spectra are shown in Fig. 1 as an example.

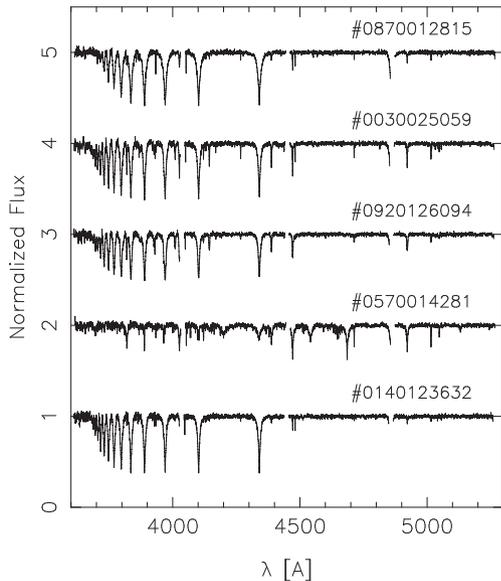
### 2.2 Radial velocities

We measured the radial velocity (RV) of target stars using IRAF task *fxcor*, cross-correlating (Simkin 1974; Tonry & Davis 1979) their spectra with synthetic templates drawn from the library of Munari et al. (2005). Previous investigations have shown that the RV measurements are not affected by the exact choice of the template, because a partial mismatch between object and template spectral types only enhances the formal uncertainties, without shifting the peak of the cross-correlation function (Morse, Mathieu & Levine 1991; Moni Bidin et al. 2011). All RVs were corrected to heliocentric values. The results are given in the fifth column of Table 1. The errors were computed as the quadratic sum of the most relevant sources. The wavelength calibration introduced a negligible uncertainty, smaller than  $1 \text{ km s}^{-1}$  for all stars. The cross-correlation error was in the range  $2$ – $8 \text{ km s}^{-1}$  (but up to  $15$ – $20 \text{ km s}^{-1}$  in four cases), and thus it usually dominated the error budget. Some systematic offsets are expected, caused by an imperfectly centred position of the star in the slit (see the discussion in Moni Bidin et al. 2006). These are hard to estimate and remove in the absence of deep telluric lines in the observed spectral range. Instead, they were estimated from the rms of the residual RV offsets of the standard stars, and thus approximated as a fixed  $4 \text{ km s}^{-1}$  for all targets. The spectra were eventually shifted to laboratory wavelength for the measurement of the stellar parameters.

<sup>1</sup> IRAF is distributed by the National Optical Astronomy Observatories, which are operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation.

**Table 1.** Data and measurements of target stars.

SPM ID	RA (h:mm:ss.s)	Dec ( $^{\circ}$ : $'$ : $''$ )	$V$	$RV$ ( $\text{km s}^{-1}$ )	$T_{\text{eff}}$ K	$\log(g)$ (dex)	$\log\left(\frac{N(\text{He})}{N(\text{H})}\right)$ (dex)	$V\sin(i)$ ( $\text{km s}^{-1}$ )	note
0030025059	3:32:14.6	-83:44:13	14.08	$183 \pm 6$	$17\,000 \pm 250$	$3.88 \pm 0.06$	-1.00	<50	MS
0870012815	3:42:24.1	-64:03:21	14.82	$19 \pm 7$	$18\,800 \pm 350$	$4.69 \pm 0.06$	$-2.20 \pm 0.09$	$\sim 50$	sdB/WD
0140123632	3:47:26.8	-78:41:54	15.90	$413 \pm 11$	$15\,900 \pm 500$	$3.81 \pm 0.09$	$-2.15 \pm 0.15$	$\sim 50$	post-HB?
0570014281	3:50:37.7	-69:20:57	14.70	$-24 \pm 15$	$42\,600 \pm 200$	$6.19 \pm 0.09$	$+0.15 \pm 0.09$	$\sim 50$	sdB/WD
0320149401	4:13:50.7	-72:18:52	15.95	$26 \pm 10$	$25\,300 \pm 500$	$5.55 \pm 0.06$	$-2.83 \pm 0.12$	<50	sdB/WD
0320022407	4:15:21.2	-72:51:41	13.93	$-4 \pm 7$	$24\,100 \pm 500$	$5.18 \pm 0.06$	$-2.56 \pm 0.12$	<50	sdB/WD
1660012832	4:15:30.2	-54:21:59	14.91	$-7 \pm 21$	$50\,600 \pm 700$	$6.50 \pm 0.09$	$+0.34 \pm 0.27$	<50	sdB/WD
0880011003	4:19:30.8	-65:22:13	13.98	$12 \pm 6$	$25\,300 \pm 500$	$5.02 \pm 0.06$	$-2.08 \pm 0.09$	<50	sdB/WD
0320115607	4:19:55.5	-73:51:45	16.04	$164 \pm 9$	$24\,500 \pm 400$	$4.03 \pm 0.06$	$-1.03 \pm 0.09$	$\sim 100$	MS
0320018938	4:19:58.5	-73:52:26	14.95	$193 \pm 8$	$22\,300 \pm 400$	$3.78 \pm 0.09$	$-1.02 \pm 0.09$	$\sim 100$	MS
0040202213	4:20:32.1	-82:06:35	15.89	$-27 \pm 21$	$47\,500 \pm 800$	$6.14 \pm 0.15$	$+0.78 \pm 0.18$	<50	sdB/WD
1250013582	4:26:29.2	-59:06:10	14.29	$-17 \pm 6$	$26\,000 \pm 500$	$5.41 \pm 0.06$	$-2.84 \pm 0.09$	<50	sdB/WD
1250015381	4:36:09.8	-58:37:07	14.90	$212 \pm 7$	$21\,100 \pm 400$	$3.50 \pm 0.09$	$-0.95 \pm 0.09$	$\sim 100$	MS
2740050272	4:42:20.2	-45:52:29	15.73	$50 \pm 12$	$59\,000 \pm 3000$	$6.01 \pm 0.12$	$-2.22 \pm 0.15$	<50	sdB/WD
1270012151	5:16:10.6	-60:57:37	14.81	$2 \pm 6$	$25\,700 \pm 400$	$5.40 \pm 0.06$	$-2.66 \pm 0.09$	<50	sdB/WD
1680007950	5:29:01.2	-56:33:10	15.06	$43 \pm 7$	$27\,400 \pm 400$	$5.50 \pm 0.06$	$-2.53 \pm 0.09$	<50	sdB/WD
0340108821	5:57:02.2	-75:40:23	16.04	$-52 \pm 11$	-	-	-	-	sdB/WD
1280025746	5:58:49.6	-58:56:27	15.25	$303 \pm 5$	$21\,700 \pm 350$	$3.65 \pm 0.06$	-1.00	<50	MS
1280143250	6:00:00.6	-59:01:03	16.01	$-34 \pm 9$	$42\,200 \pm 300$	$5.62 \pm 0.15$	$+0.58 \pm 0.09$	<50	sdB/WD
1280031295	6:00:22.1	-57:54:20	15.54	$368 \pm 8$	$16\,000 \pm 250$	$3.33 \pm 0.06$	-1.00	$\sim 50$	MS
0910042645	6:05:34.6	-62:55:06	15.66	$-12 \pm 7$	$27\,800 \pm 400$	$5.48 \pm 0.06$	$-2.89 \pm 0.12$	<50	sdB/WD
3380011095	6:15:47.1	-41:32:06	14.65	$424 \pm 6$	$16\,100 \pm 200$	$4.05 \pm 0.06$	-1.00	<50	MS
1280035162	6:20:38.5	-57:05:38	14.72	$-11 \pm 5$	$34\,100 \pm 500$	$5.60 \pm 0.06$	$-0.89 \pm 0.06$	<50	sdB/WD
0350036197	6:46:44.1	-72:10:37	13.64	$23 \pm 5$	$29\,400 \pm 900$	$5.14 \pm 0.12$	$-1.16 \pm 0.09$	<50	sdB/WD
0600320649	6:48:41.5	-68:30:36	16.14	$-4 \pm 16$	$42\,400 \pm 300$	$5.66 \pm 0.15$	$+0.55 \pm 0.09$	<50	sdB/WD
0600047219	6:53:13.1	-66:59:51	14.64	$55 \pm 9$	$37\,500 \pm 900$	$5.38 \pm 0.12$	$-3.16 \pm 0.18$	<50	sdB/WD
0920126094	6:59:58.0	-66:18:28	16.38	$284 \pm 7$	$19\,100 \pm 400$	$3.78 \pm 0.06$	-1.00	$\sim 100$	MS
0350232289	7:04:25.2	-72:38:43	15.89	$280 \pm 8$	$22\,400 \pm 600$	$3.69 \pm 0.09$	-1.00	$\sim 200$	MS
0350031043	7:10:55.1	-73:27:21	15.18	$-35 \pm 7$	$27\,100 \pm 600$	$5.39 \pm 0.09$	$-2.64 \pm 0.12$	<50	sdB/WD
1310017952	7:48:37.5	-61:50:03	14.91	$8 \pm 6$	$27\,000 \pm 500$	$5.45 \pm 0.06$	$-2.18 \pm 0.09$	<50	sdB/WD
0950048906	8:39:21.8	-64:49:35	15.55	$27 \pm 8$	$32\,900 \pm 400$	$5.46 \pm 0.06$	$-3.33 \pm 0.12$	<50	sdB/WD



**Figure 1.** Examples of reduced spectra of our sample stars. The spectra are shifted vertically to avoid overlap. The SPM ID of each target is also indicated. From top to bottom, we show the spectrum of a helium-poor sdB, an MS star at the same temperature, a fast-rotating MS object, a hot WD and the uncertain SPM4 object 0140123632 discussed in Section 3.1.

### 2.3 Stellar parameters

The temperature, gravity and surface helium abundance of the target stars were measured by fitting the observed hydrogen and helium lines with synthetic spectra. The model spectra were computed with Lemke's version<sup>2</sup> of the Linfor program (developed originally by Holweger, Steffen and Steenbock at Kiel University), fed with model atmospheres computed with ATLAS9 (Kurucz 1993). The metallicity was set to solar, as appropriate for both recently formed stars and hot foreground subdwarfs, whose atmospheres are enriched with heavy elements due to diffusion processes (Moehler et al. 2000). The grid of templates covered the range  $8000 \leq T_{\text{eff}}(\text{K}) \leq 35\,000$ ,  $2.5 \leq \log(g) \leq 6.0$ ,  $-3.0 \leq \log\left(\frac{N(\text{He})}{N(\text{H})}\right) \leq -1.0$ . The synthetic spectra were convolved with a Gaussian during the fitting procedure, to match the instrumental resolution of the data. The spectra of the few evolved objects hotter than 35 000 K were fitted using a grid of metal-free non-LTE models described in Moehler et al. (2004), calculated as in Napiwotzki (1997).

The best fit and the derivation of the stellar parameters were established through the routines developed by Bergeron, Saffer & Liebert (1992) and Saffer et al. (1994), as modified by Napiwotzki, Green & Saffer (1999). The code normalizes both the model and the observed spectra simultaneously using the same points for the continuum definition, and makes use of a  $\chi^2$  test to establish the best fit. The noise in the continuum spectral regions is used to estimate the  $\sigma$  for the calculation of the  $\chi^2$  statistics, which the routines use

<sup>2</sup> <http://a400.sternwarte.uni-erlangen.de/~ai26/linfit/linfor.html>

to estimate the errors of the parameters (see Moehler, Sweigart & Catelan 1999). However, they thus neglect other sources of errors, such as those introduced by the normalization procedure, the sky subtraction and the flat-fielding. Therefore, the resulting uncertainties were multiplied by three to obtain a more realistic estimate of the true errors (Napiwotzki 2005, private communication). The lines of the Balmer series from  $H_\beta$  to  $H_{12}$  were simultaneously fitted, along with four  $\text{He I}$  lines (4026 Å, 4388 Å, 4471 Å, 4922 Å). Two  $\text{He II}$  lines (4542 Å, 4686 Å) were also included when visible in the spectra of hot stars. Results for all the targets are given in Table 1. We could not find a good solution for the spectrum of star SPM4 0340108821, nor could we determine if this was due to problems related to its spectrum, or to limitations of our model grid. This object will therefore be excluded from further analysis. However, its strong and wide  $\text{He II}$  lines reveal that it is most likely a foreground WD ( $\log(g) > 6$ ,  $T_{\text{eff}} > 40\,000$  K).

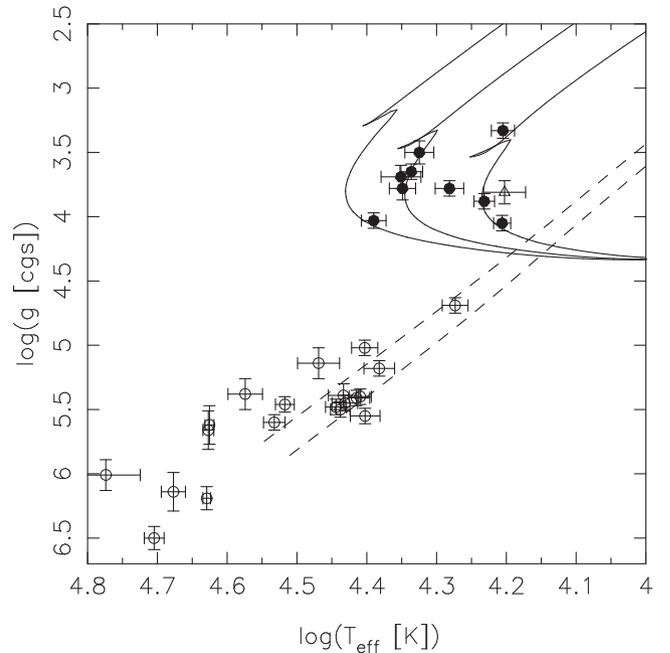
The lines of some stars were clearly broadened due to high projected rotational velocity. The employed routine did not include  $V\sin(i)$  as a fit parameter, but as a fixed input quantity beforehand. We then estimated  $V\sin(i)$  varying this value manually in such a way as to minimize the resulting  $\chi^2$  of the fit. A precise estimate of  $V\sin(i)$  is not possible due to the relatively low resolution. In addition, the wide Balmer lines dominated the fit statistics, but these features were intrinsically too wide for this purpose. As a consequence, our results are reliable only for a rough distinction between stars showing no relevant evidence of rotation ( $V\sin(i) < 50$  km s $^{-1}$ ), stars showing hints of rotation ( $V\sin(i) \sim 50$  km s $^{-1}$ ) and fast rotators ( $V\sin(i) \sim 100$  or 200 km s $^{-1}$ ). The results are given in the ninth column of Table 1.

The surface helium abundance was assumed as a fit parameter in the procedure, and the model grid spanned a wide range from solar value ( $\log(\frac{N(\text{He})}{N(\text{H})}) = -1$ ) down to one-thousandth of solar abundance. This is appropriate for foreground subdwarfs, and this is, in fact, a key parameter to identify them in our sample (see Section 3.1). However, diffusion processes are not active in MS stars, and their surface helium abundance is in general close to solar. This parameter is therefore not very informative for these objects. On the other hand, we found that fixing  $\log(\frac{N(\text{He})}{N(\text{H})}) = -1$  led to better fits of all the lines for some MS stars. We adopted the results obtained with helium abundance fixed to solar for these objects, and this parameter has no associated error in Table 1. We nevertheless verified that our final results were not affected by this choice as, for example, distance moduli and ages changed by less than  $1-1.5\sigma$ .

### 3 RESULTS

#### 3.1 MS stars and foreground objects

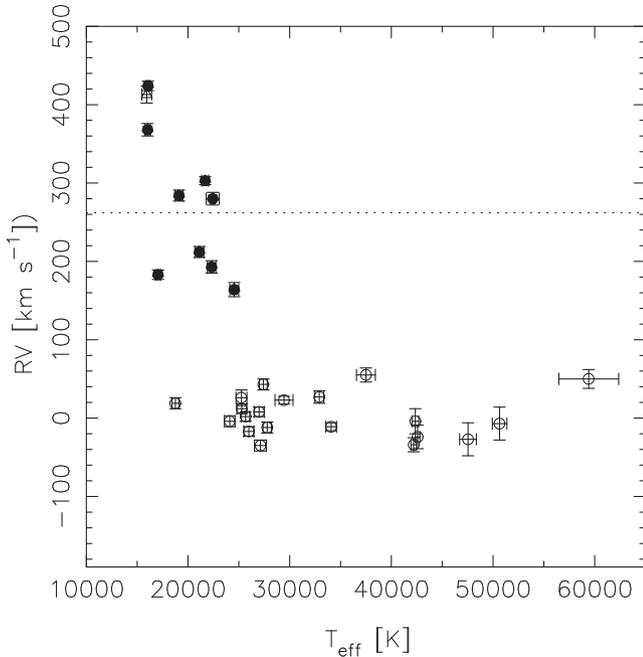
The RV distribution of target stars is clearly bimodal, with a huge gap between 10 high-velocity stars ( $RV > 160$  km s $^{-1}$ ) and the rest of the sample ( $RV < 55$  km s $^{-1}$ ). This dichotomy suggests an easy separation between LMC members and foreground stars, because the RV of the LMC ( $262.1 \pm 3.4$  km s $^{-1}$ ; van der Marel et al. 2002) is very high compared to the typical RV of Galactic stars along the same line of sight. However, our single-epoch measurements alone are not enough to uniquely identify LMC stars from Galactic interlopers, since the incidence of close binaries is very high among both massive OB stars (Sana et al. 2012) and subdwarfs (Maxted et al. 2001), although the binary fraction for old halo sDBs could be much lower than in the disc (Han 2008; Moni Bidin, Catelan & Altmann 2008). Galactic runaway stars, on the other



**Figure 2.** Location of the programme stars in the temperature-gravity plane. Filled and open circles represent MS and foreground compact objects, respectively. The open triangle indicates the star whose classification is uncertain. Solid curves show solar-metallicity isochrones of 10, 20 and 50 Myr. The dashed curves indicate the zero-age and terminal-age horizontal branch.

hand, are predominantly single objects (Mason et al. 1998), but they can show anomalously high RV due to their intrinsic nature of being high-velocity stars. However, the measured stellar parameters provide a straightforward identification of sDBs and WDs. In fact, their surface gravity is higher than that of MS stars at the same temperature for  $T_{\text{eff}} > 15\,000$  K (see e.g. Salgado et al. 2013), and their loci in temperature-gravity space are well separated. The atmospheres of sDBs with  $T_{\text{eff}} < 35\,000$  K are also depleted of helium (typically  $\log(\frac{N(\text{He})}{N(\text{H})}) < -1.5$ , Moni Bidin et al. 2009), at variance with the close-to-solar values expected for MS stars. In addition, while fast rotators are common in the MS, they are extremely rare among sDB's (Geier & Heber 2012).

Fig. 2 shows the position of target stars in the temperature-gravity plane. Solar-metallicity isochrones from Bressan et al. (2012) are also shown as solid curves. Dashed lines indicate the zero-age and terminal-age horizontal branch, which enclose the loci where stars spend 90 per cent of their helium core-burning stage. 20 stars lie on the loci of sDBs and WDs, and their surface gravity is too high for MS objects ( $\log(g) > 4.7$ ). All but two of the high-gravity stars with  $T_{\text{eff}} < 40\,000$  K are strongly helium-depleted ( $\log(\frac{N(\text{He})}{N(\text{H})}) < -2$ , while surface helium abundance increases for the hottest stars up to supersolar values, as is usual for O-type subdwarfs and hot WDs (e.g. Edelman et al. 2003). Only two of these stars show hints of rotation (SPM4 0870012815 and SPM4 0570014281). All these objects are therefore classified as foreground stars. The gravity of the coolest star of our sample (SPM4 0140123632) is too low to be an sDB. However, its very low helium abundance is incompatible with it being an MS star. While its classification should be regarded as dubious, we interpret it as a post-horizontal branch star, evolving at lower gravities towards the asymptotic giant branch after exhaustion of helium in the core. It should be a foreground object even in this case because its absolute magnitude would have to be too bright



**Figure 3.** RV of target stars as a function of their effective temperature. The symbols are as in Fig. 2. The dotted line indicates the systemic RV of the LMC.

for a post-HB star at its temperature to place it at the distance of the LMC ( $(m - M)_0 \approx 18.5$ ; Freedman et al. 2001). The remaining nine targets are clearly MS stars, as deduced from their position in the temperature-gravity plane and their close-to-solar surface helium abundance. Six of them also show evidence of non-negligible rotation. The classification of each target (‘sdB/WD’, ‘post-HB?’ or ‘MS’) is given in the last column of Table 1.

The RVs of our targets are shown in Fig. 3 as a function of effective temperature, where we indicate foreground compact objects and MS stars with open and filled symbols, respectively. The figure reveals that the RV of all foreground sdBs/WDs is lower than  $55 \text{ km s}^{-1}$ , while the aforementioned high-velocity stars are all MS stars, with the addition of the uncertain post-HB object. The foreground stars will not be considered further, and we will focus on the nine MS stars hereafter.

### 3.2 Distances and ages

The age of the MS stars was estimated comparing their position in the temperature-gravity plane with solar-metallicity isochrones, as shown in Fig. 2. Varying the stellar parameters within their error bars caused small variations in the derived age, of the order of 10 per cent. A larger uncertainty was introduced by the assumption of the isochrone metallicity, as a change of  $\pm 0.5$  dex affected the results by  $\mp 20$  per cent or more. We therefore assumed this uncertainty as the final error.

The observed  $V$  magnitudes were de-reddened by means of the line-of-sight values given by the Schlegel, Finkbeiner & Davis (1998) maps, as corrected by Bonifacio, Monai & Beers (2000), assuming a standard reddening law ( $R = \frac{A_V}{E(B-V)} = 3.1$ ). An estimate of the distance modulus was thus obtained, by comparison with the absolute magnitude in the  $V$  band obtained from the same solar-metallicity isochrones. The assumption of  $[\text{Fe}/\text{H}] = \pm 0.5$  for the theoretical isochrones shifted the absolute magnitude by

$\mp 0.2$  mag at most ( $\mp 0.12$  mag on average). A larger but comparable uncertainty on  $(m - M)_0$  was induced by the observational errors of the stellar parameters (0.23 mag on average), and the two values were quadratically summed to obtain the final errors. The results for reddening, age, distance modulus and line-of-sight velocity with respect to the Local Standard of Rest  $V_{\text{lsr}}$  are given in Table 2, where we also add an internal ID (first column) and sky coordinates (see Section 4).

Fig. 4 shows the RV of our MS objects as a function of their distance modulus, compared to the LMC values  $RV_{\text{LMC}} = 262.1 \pm 3.4 \text{ km s}^{-1}$  (van der Marel et al. 2002) and  $(m - M)_{0, \text{LMC}} = 18.5$  (distance  $d = 50$  kpc; Freedman et al. 2001). Two stars are clearly foreground objects ( $(m - M)_0 \approx 15.5$ ,  $d \approx 12.5$  kpc). A third more distant star ( $d \approx 33$  kpc), whose  $(m - M)_0$  is incompatible with the LMC at the  $3.3\sigma$  level, is also likely not an LMC member. These three stars are probably Galactic runaway stars, as also suggested by their high recession velocity. The distance modulus of the six remaining stars agrees with the LMC value within  $1\sigma$ . Therefore, we consider these six young stars as being members of the LMC.

## 4 ANALYSIS

In what follows, we adopt the Magellanic coordinate system introduced by Nidever et al. (2010), in which  $\Lambda_M$  is the longitude along the Magellanic Stream, positive towards the direction of motion of the LMC (i.e. towards the Galactic plane) and  $B_M$  is the Magellanic latitude. Our adopted centre for the LMC is (RA, Dec) = (81:90, -69:86), from van der Marel et al. (2002). This centre, however, does not correspond to the origin of the Magellanic coordinate system, but is slightly offset:  $(\Lambda_{M, \text{LMC}}, B_{M, \text{LMC}}) = (0:215, 2:330)$ . We adopt this system since it reflects the orbital motion of the LMC and – at the location of the LMC – it is closely aligned with the equatorial system (a  $\sim 2^\circ$  offset). Finally, the spherical  $(\Lambda_M, B_M)$  coordinates are projected on to a plane tangent to the sky at  $(\Lambda_M, B_M) = (0^\circ, 0^\circ)$ , a gnomonic projection with coordinates  $(\xi_{\Lambda_M}, \eta_{B_M})$ . Units for sky coordinates are degrees throughout the paper.

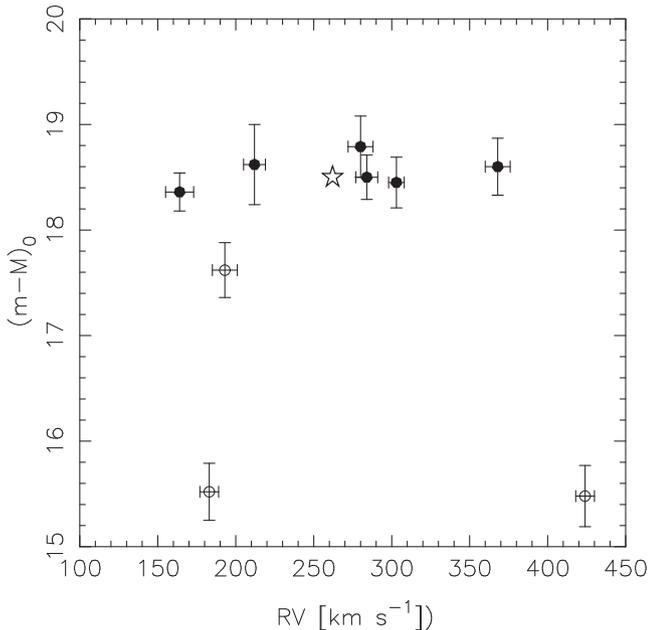
### 4.1 Matching with existing catalogues of clusters and associations

We have cross-matched our six stars with the catalogues of Bica et al. (2008). In Fig. 5, we show the spatial distribution of our six stars, and of the catalogue of clusters (left-hand panel), young associations (middle panel) and emission nebulae (right-hand panel). Target no. 403 is 15 arcmin from the centre of the globular cluster Reticulum. However, this star is not a cluster member, since it is a young star, and its RV of  $212 \pm 7 \text{ km s}^{-1}$  is different from that of the cluster ( $241.5 \pm 1.5 \text{ km s}^{-1}$  Suntzeff et al. 1992). In Fig. 5, we also mark the newly found, loose cluster by Piatti (2016). This 280 Myr open cluster is at a distance some 11.3 kpc larger than that of the LMC. However, its sky location is well within the distribution of the clusters catalogued by Bica et al. (2008).

The target no. 292 is within the stellar association ICA76 (aka IDK6) as given in Bica et al. (2008), Irwin et al. (1990) and Battinelli & Demers (1992). Specifically, no. 292 is at 2 arcmin from the centre of ICA76, which has a radius of  $\sim 7$  arcmin (Bica et al. 2008). We note that star no. 290 is also located within the projected radius of ICA76, but the distance modulus determined here (Table 2) places it in the foreground. Demers & Battinelli (1998) determine a distance modulus of  $18.63 \pm 0.10$  to ICA76 and an age of 10–25 Myr. This agrees reasonably well with the distance and age of star no. 292 (see Table 2). Demers & Battinelli (1998) also quote previous RV

**Table 2.** Reddening, age, distance modulus,  $V_{\text{lsr}}$  and Magellanic coordinates of the target MS stars.

ID	SPM ID	$E(B - V)$ (mag)	Age (Myr)	$(m - M)_0$ (mag)	$V_{\text{lsr}}$ ( $\text{km s}^{-1}$ )	$\xi_{AM}$ ( $^\circ$ )	$\eta_{BM}$ ( $^\circ$ )
005	0030025059	0.10	$50 \pm 20$	$15.52 \pm 0.27$	174	-2.281	-12.611
292	0320115607	0.10	$11 \pm 6$	$18.36 \pm 0.18$	153	-4.247	-2.510
290	0320018938	0.10	$20 \pm 5$	$17.62 \pm 0.26$	182	-4.240	-2.519
403	1250015381	0.01	$20 \pm 5$	$18.62 \pm 0.38$	198	-7.065	12.871
405	1280025746	0.04	$20 \pm 5$	$18.45 \pm 0.24$	288	3.910	13.412
406	1280031295	0.05	$45 \pm 10$	$18.60 \pm 0.27$	353	4.220	14.491
487	3380011095	0.07	$55 \pm 30$	$15.48 \pm 0.29$	407	9.457	33.541
390	0920126094	0.11	$35 \pm 8$	$18.50 \pm 0.21$	271	9.310	4.555
307	0350232289	0.16	$18 \pm 5$	$18.79 \pm 0.29$	268	7.447	-1.609

**Figure 4.** Distance modulus of the target MS stars as a function of their RV. Filled and open circles indicate the LMC members and the foreground runaway stars, respectively. The star symbol indicates the value of the LMC.

measurements of these two stars: they are  $194$  and  $270 \text{ km s}^{-1}$  for no.290 and no.292, respectively. The RV for no.290 agrees with our measurement ( $193 \pm 8 \text{ km s}^{-1}$ ), however, that for no.292 does not ( $164 \pm 4 \text{ km s}^{-1}$ ). While we do not have an explanation for this discrepancy, one reason may be binarity.

We have cross-matched our six stars with the recent DES cluster catalogue by Pieres et al. (2016), and found no matches. We have also inspected  $15 \text{ arcmin} \times 15 \text{ arcmin}$  *GALEX* and Digital Sky Survey B images centred on each of our six stars. Excepting no.292, where the stellar association ICA76 is clearly visible, the remaining five stars appear isolated. To conclude, our six young stars are much further out in projected distance from the LMC's centre than the catalogued clusters, associations and emission nebulae. Only one star in our sample is a likely member of a known young stellar association located in the Bridge, while the other five objects cannot be identified as being associated with any known stellar aggregate.

## 4.2 Spatial distribution

In Fig. 6, we show the area coverage of the CD12 study to select OB-type candidates in the region of the Clouds. Gaps in the coverage are primarily due to the *GALEX* GR5 survey. Candidates from CD12

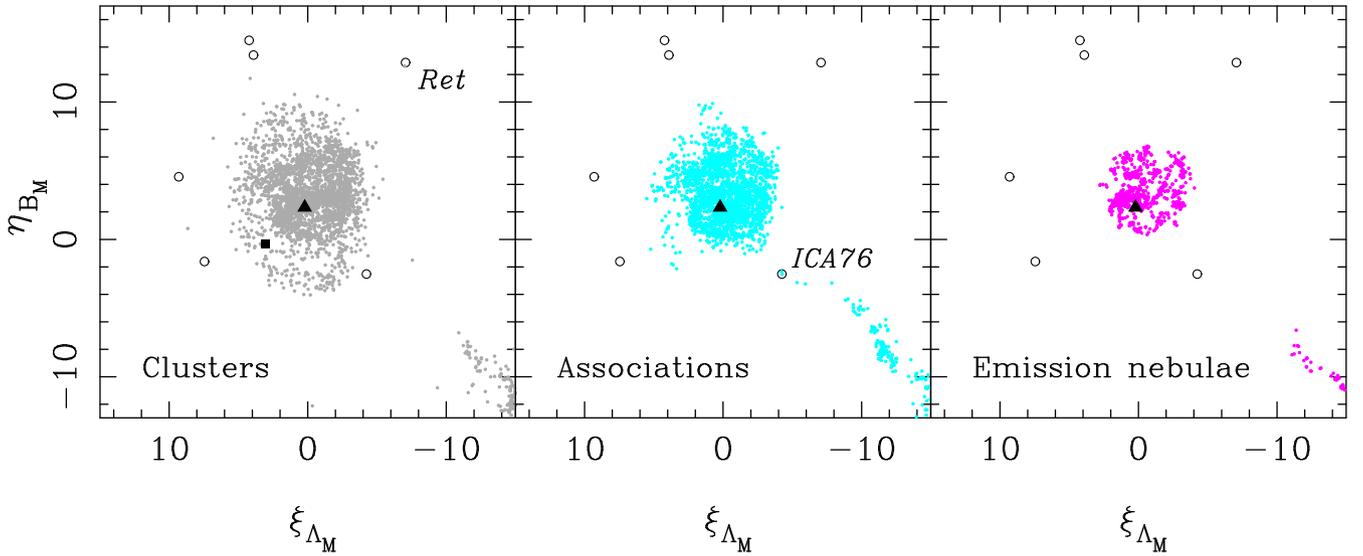
are shown with crosses: clearly, only the outer regions of the Clouds were probed by CD12. The Galactic plane is towards the left of this figure. We also mark two useful locations in this plot: the centre of Carina dwarf galaxy and the first detection of  $\text{HCO}^+$  absorption at the leading edge of the Magellanic Bridge by Murray et al. (2015). Muñoz et al. (2006) found giant stars, likely members of the LMC in the Carina foreground, thus hinting at the large spatial extent of the Cloud, while the detection of  $\text{HCO}^+$  is indicative of star formation. Fig. 6 is meant to caution the reader as to the areal incompleteness.

In Fig. 7, we show the map of the  $\text{H I}$  column density in our region centred on the LMC. The  $\text{H I}$  data are from the GASS III cleaned data described in Kalberla & Haud (2015), within the velocity range  $V_{\text{lsr}} = 100\text{--}450 \text{ km s}^{-1}$ . OB-type candidates from CD12 are shown with grey crosses. Stars observed spectroscopically in this study are shown with circles, young stars being highlighted in green and the six LMC members are marked with a square. As in Fig. 6, Carina's centre and the location of  $\text{HCO}^+$  detection are indicated. The thick blue line marks the newly found 10-kpc tidal feature/arm by Mackey et al. (2016) in the DES. The thick blue dotted line marks the western extension of this arm, a feature less conspicuous than the arm.

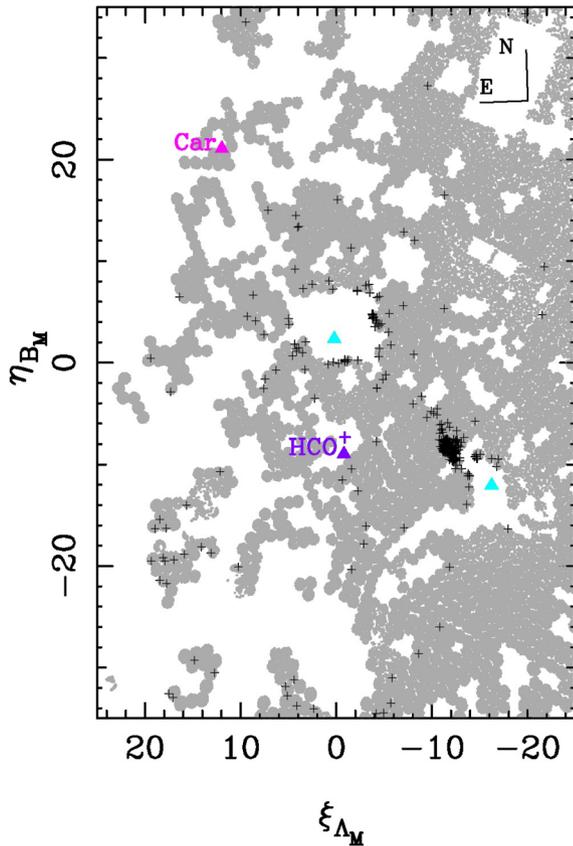
Our six stars are surrounding the LMC out to  $\sim 13^\circ$  from its centre. Star no.292 that belongs to association ICA76 is clearly located in the Bridge, and it is the closest in projected distance to LMC's centre. Two stars, no.405 and no.406, are very close to each other (within  $\sim 1^\circ$ ), while the remaining three are at various position angles. Except for no.292 that is in the Bridge, the stars are in regions of low  $\text{H I}$  column density,  $N_{\text{HI}} < 10^{19} \text{ cm}^{-2}$ . It is intriguing that the two close together stars, no.405 and no.406, are only some  $5^\circ$  behind the stellar tidal arm found by Mackey et al. (2016). Mackey et al. (2016) also indicate the existence of a western extension of this tidal arm in their fig. 2 (the blue dotted line in our Fig. 7). This extension, if real, is very near to our star no.403, with the star slightly closer to the LMC centre than the extension. We also note that Besla et al. (2016) found stellar arcs and spiral arms in the northern periphery of the LMC. These features are within  $\sim 8^\circ$  of the LMC's centre, specifically in region C of Besla et al. (2016), their fig. 4. Our study does not properly sample these structures, since we focused on more distant objects (see Fig. 7). None of our candidates lying close to the  $\text{HCO}^+$  detection proved to be LMC members, but that region of the sky is affected by a large areal gap in the survey data (see Fig. 6).

## 4.3 Testing a ring-like configuration

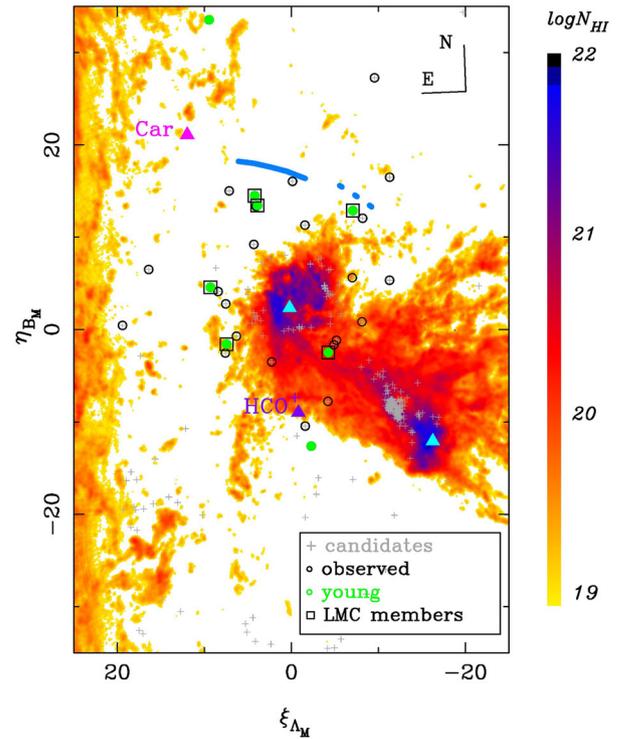
Our five stars (excluding the ICA76 member) appear to be in a half-ring configuration, with the ring's centre some  $3^\circ$



**Figure 5.** Spatial distribution of our six young stars (open circles) compared to that of clusters (left), young associations (middle) and emission nebulae (right) from Bica et al. (2008). The centre of the LMC is marked with a triangle. The positions of the globular cluster Reticulum and the stellar association ICA76 are also labelled. The dark square marks the newly discovered, loose cluster by Piatti (2016).



**Figure 6.** Area coverage of the CD12 study to select candidates (grey background). The round footprint of GALEX is easily visible; the gaps in area coverage are mainly due to the GALEX GR5 coverage. Young, OB-type star candidates from CD12 are shown with crosses. The centres of the LMC and SMC are shown with light-blue triangles. Two notable sky locations are also indicated and labeled: the Carina dwarf galaxy and the first detection of HCO<sup>+</sup> absorption in the Magellanic System by Murray et al. (2015). North and east directions are also indicated.



**Figure 7.** Map of the H I column density from the GASS survey within  $V_{\text{lsr}} = 100$  to  $450 \text{ km s}^{-1}$  in the region of interest. The centres of the Clouds are shown with light blue triangles. OB-type candidates from CD12 are shown with grey crosses. Circles represent the stars observed spectroscopically in this study. Filled green circles indicate the young stars. Squares indicate the six young stars that are members of the LMC. The thick blue curve indicates the newly found 10-kpc tidal feature/arm by Mackey et al. (2016) in the Dark Energy Survey. The thick dotted blue curve shows the western extension of this arm.

northwards of LMC's centre, and spanning some  $\sim 140^\circ$  in position angle (see Fig. 7). Note that some 10–12 candidates at moderate  $\sim 7^\circ$ – $10^\circ$  distance from the LMC's centre are foreground stars (see open circles in Fig. 7). This leaves a gap between a region of inner star formation, as traced by emission nebulae and young associations (Fig. 5), and the off-centre, ring-like configuration of our five stars. Also, the five stars coincide with the 'near' part of the LMC disc (see Fig. 9). Alas, our distance uncertainties are inadequate to explore the 'nearness' of these five stars.

Our targets were selected from the CD12 candidates located between  $6^\circ$  and  $33^\circ$  from the centre. Hence, the spatial distribution of the observed sample has already some degree of circular symmetry about the LMC centre. One might wonder to what extent the ring-like feature traced by the five young LMC stars could be an effect induced by the underlying distribution of the selected targets. That is, while the stars found to be young LMC members appear to lie in the form of a ring, are they really that much more 'ringlike' than other, random groupings of five stars that might have been drawn from the observed sample? To answer this question, we perform a Monte Carlo test to quantify the statistical significance of the observed configuration.

From the 31 spectroscopically observed stars, we discard four that are located in the Bridge. From this set of 27 stars, we randomly select five stars and project their sky positions into the LMC disc plane. The geometry of the disc is adopted from Olsen et al. (2011), with an inclination angle of  $34.7^\circ$ , and a position angle of the line of nodes of  $142.0^\circ$  (see also Section 4.4). We search in a grid around the LMC's centre for a ring-centre position that gives the five stars a minimum in  $\delta r/r_{\text{ave}}$ , where  $r_{\text{ave}}$  is the average radius of a circle they define, and  $\delta r$  is the root-mean-square scatter of the stars' positions from a perfect circle. The optimal centre and  $r_{\text{ave}}$  define the 'best-fitting' circular ring for this group of five stars. We also calculate the arc length that the five stars subtend along the best-fitting circle. The search for the ring centre is done in steps of  $0.25^\circ$  within  $7^\circ$  of the LMC's centre. We perform 1000 such random, five-star trials and store the resulting values of  $\delta r/r_{\text{ave}}$  and arc length. Then, we repeat this 1000-trial run a total of 10 times. For each run, we determine how many groups of five stars form a 'better' circle than the actual five LMC stars, i.e. have  $\delta r/r_{\text{ave}}$  smaller than that of the five LMC stars. We find that the angular arc length of our five stars is fairly typical of randomly drawn ones. However, on average, only 1.35 per cent of randomly drawn groups form an arc of a circle with  $\delta r/r_{\text{ave}}$  lower than that of the five LMC stars. This result suggests that an accidental hypothesis can be ruled out at the 98.65 per cent confidence level. The ringlike configuration of these five stars is therefore unlikely to be due to chance, and a physical basis is suggested. We speculate as to the possible origin of such a structure in Section 4.5.

The ring parameters found for the five LMC stars are: average radius  $r_{\text{ave}} = 13.1^\circ$ , distance from the LMC's centre  $\rho = 3.0^\circ$  and position angle arclength  $\Delta\text{PA} = 145^\circ$ . At a distance of 50 kpc for the LMC, the average radius is 12 kpc and the distance from LMC's centre is  $2.6 \pm 0.3$  kpc. All of these parameters are in the LMC's disc plane.

While our statistical test shows it is unlikely that the ring is mere coincidence – and we proceed to derive the geometry in the plane of the LMC as if this were a real structure – this is all based on five stars situated over a rather large portion of sky. Determining whether these stars actually reside in the LMC disc will require additional observations (primarily astrometric) that are not available at this time.

#### 4.4 Velocity distribution

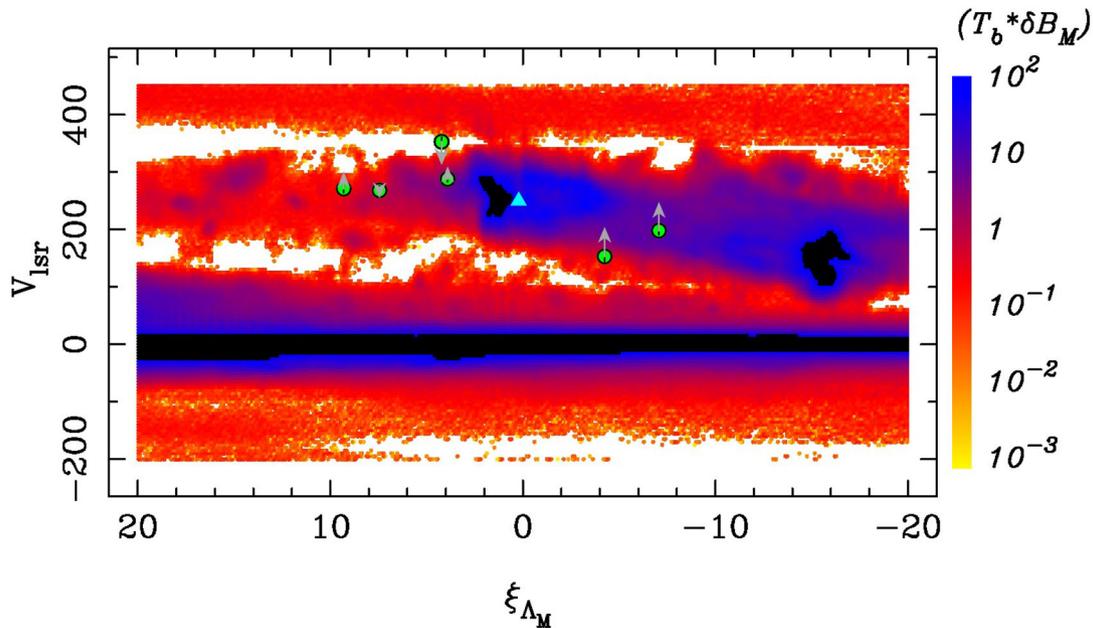
We compare our measured radial velocities for the LMC members with that of the H I gas (Kalberla & Haud 2015) in Fig. 8. Plotted is the local standard of rest radial velocity  $V_{\text{lsr}}$  as a function of  $\xi_{\Lambda_M}$ , the coordinate along the orbit of the LMC. The colour scale represents the temperature brightness integrated along  $\eta_{B_M}$  for the range  $-30^\circ$  to  $30^\circ$ . The locations of the Galactic disc and the central parts of the Clouds are clearly visible. Our young LMC members are shown with green symbols. They follow the overall pattern of the gas motions, with the two most deviant points being star no.292 at  $\xi_{\Lambda_M} = -4.247^\circ$  and star no.406 at  $\xi_{\Lambda_M} = 4.220^\circ$ . Star no.292 is a member of ICA76, located in the Bridge; and in Section 4.1 we noted a previous velocity measurement for this star ( $\text{RV} = 270 \text{ km s}^{-1}$ , or  $V_{\text{lsr}} = 259 \text{ km s}^{-1}$ ), which would bring it in better agreement with the gas motion.

We also compare our velocity measurements with the predictions of a disc model for the LMC. We adopt the geometry and rotation curve of the disc from Olsen et al. (2011); these were determined from the analysis of red supergiants, a relatively young population. The adopted distance to the LMC is 50 kpc and its RV is  $263 \text{ km s}^{-1}$ . The inclination angle of the disc is  $34.7^\circ$ , and the position angle of the line of nodes is  $142.0^\circ$ . The model assumes circular motions only, and we use the formalism from van der Marel et al. (2002) to calculate velocities at a given position angle and angular distance from the LMC's centre. The arrows in Fig. 8 point to the velocity model predictions at each star's location in  $(\xi_{\Lambda_M}, \eta_{B_M})$ , i.e. the length of the arrow is the velocity difference between observation and model. These differences are rather modest, indicating that the velocities of the stars are not too deviant from simple, circular motion in the disc of the LMC.

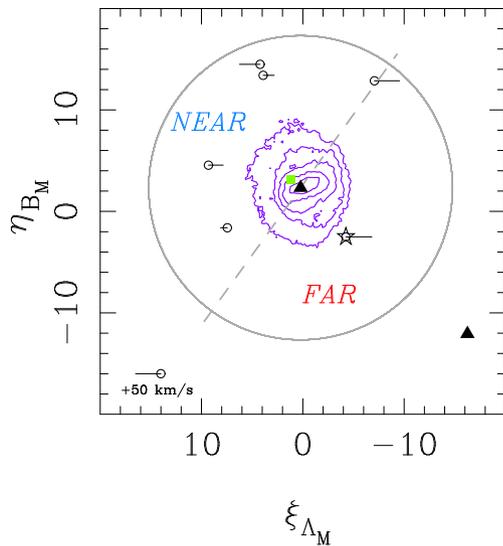
To explore whether the velocity differences have any correlation with the distance from LMC centre, or with the position angle, we plot these velocity differences at the sky location of each star in Fig. 9. For reference, we show isodensity contours of M giants in the LMC selected from 2MASS as in CD12. The LMC's bar is easily visible in the inner regions. We also mark the location of the star-forming region 30 Doradus with a green square. The line of nodes of the disc is shown with a dashed line, and the 'near' and 'far' halves of the disc are labelled. The six young LMC members are shown with circles, and the member of ICA76 is highlighted with a star symbol. A circle with a radius of  $15^\circ$  centred on the LMC is also shown. Velocity differences do not show a correlation with either projected distance from LMC centre or with position angle. The average of these differences is  $-16 \pm 15 \text{ km s}^{-1}$ , with a standard deviation of  $36 \text{ km s}^{-1}$ . If we exclude star no.292, the member of ICA76, we obtain an average velocity difference of  $-9 \pm 16 \text{ km s}^{-1}$ , with a similar standard deviation of  $36 \text{ km s}^{-1}$ .

#### 4.5 Origin of the young stars

Inspecting Table 2, we see that star no.292 is 11 Myr old, the youngest in our sample of six stars. The ages of the remaining five stars range between 18 and 45 Myr, with an average of  $\sim 30$  Myr. If we assume that the standard deviation of the line-of-sight velocities is representative of the deviations from simple LMC disc rotation for the tangential components as well, we can estimate the maximum distance in the sky these stars will travel in 30 Myr. For  $36 \text{ km s}^{-1}$  (Section 4.4), this distance amounts to 1 kpc, or  $\sim 1^\circ$  at the distance of the LMC. Thus, in the absence of precise proper motions for these stars, we argue that line-of-sight velocities of the young LMC members indicate that they have not travelled far



**Figure 8.**  $V_{\text{lsr}}$  as a function of  $\xi_{\Lambda_M}$  along the LMC's orbit. The colour scale shows the temperature brightness integrated along  $\eta_{B_M}$  for the range  $-30^\circ$  to  $30^\circ$ . Black regions indicate values larger than  $10^2$  K deg, while white regions indicate values less than  $10^{-3}$  K deg. Green symbols show our young LMC members. The LMC's centre is represented with a triangle. The arrows indicate the difference in velocity between our measurements and the predictions of a model of the LMC disc rotation, at each star's position.



**Figure 9.** The sky distribution of the six young LMC members with  $V_{\text{lsr}}$  differences (observed – model) shown as horizontal vectors. A vector of  $+50 \text{ km s}^{-1}$  is shown in the bottom left corner for scale. Triangles mark the centres of the LMC and SMC. The star symbol shows the stellar member of ICA76. The green square indicates the location of 30 Doradus. Contours represent isodensity curves of M giants in the LMC selected from 2MASS. The line of nodes of the disc is shown with a dashed line, with the near and far sides of the disc labelled. The circle highlights a radius of  $15^\circ$  from LMC's centre.

during their lifetimes, and their current locations are representative of their birthplace. This points to *in situ* formation rather than being runaways from star-forming regions and stellar associations in the inner LMC. Moreover, stars no. 405 and no. 406 being so close to one another further argues against the likelihood of a runaway

mechanism, since one would assume random directions for the ejecting stars.

Our findings are rather unexpected, since they point to star formation in the far outskirts of the LMC where gas density is currently low,  $N_{\text{HI}} < 10^{19} \text{ cm}^{-2}$  (Fig. 7). Fig. 5 clearly shows this as well; notice the small, inner LMC region occupied by emission nebulae and young stellar associations compared to the distribution of our stars. That the gaseous disc of the LMC is truncated when compared to the intermediate-age/old stellar disc is no longer a novelty since the work of Saha et al. (2010). Various theoretical models (see Hammer et al. 2015, and references therein) attribute this morphology primarily to truncation caused by ram pressure exerted by the Galactic hot halo gas. Traditional star formation studies of the LMC disc, based on colour–magnitude diagrams, indicate that the young star-forming epoch continues to the present time only in the inner  $\sim 5^\circ$  of the disc (Meschin et al. 2014).

Whence these  $\sim 20$ – $50$  Myr stars in the far outskirts? One intriguing clue as to a possible origin is their half-ring spatial distribution, which was shown in Section 4.3 to be unlikely a mere random configuration. If these are, in fact, in the form of a circular ring in the plane of the LMC disc, it suggests an interesting possible formation scenario.

*N*-body/SPH simulations of collisions between a barred disc galaxy and a small (20 per cent mass) satellite (Berentzen et al. 2003) show that an off-centre collision produces expanding ring structures, off-centre bars and other asymmetries in the stars and gas. Typically, the impact produces a density wave in the host's disc, which first becomes apparent in the gas. If the impact is off centre, the symmetry of the gaseous ring is broken when it encounters spiral arms. In our case, the SMC had a recent collision with the LMC's disc. The particulars of this collision – impact parameter and angle – will dictate the morphology of the outskirts of the LMC's disc (see e.g. Besla et al. 2016, and references therein). It is thus conceivable that the formation of our five young stars that appear in a ring-like

configuration was triggered by the recent collision between the SMC and the LMC's disc. At the time the Mackey et al. (2016) study was completed, the DES did not cover the eastern portion of the outskirts of the LMC where our stars no.307 and no.390 lie; it would be very informative to have that area covered by deep, contiguous photometric surveys.

The impact parameter of the collision between the SMC and the LMC disc appears to be moderate to small. Based on a simple orbit integration, CD12 finds an impact parameter of  $\sim 12$  kpc as an upper limit, and an impact angle of  $\sim 35^\circ$  (where  $0^\circ$  is coplanar). From dynamical models of the LMC–SMC interaction, Besla et al. (2016) find the impact parameter to be  $< 10$  kpc in order to reproduce the arcs found in their study, while Besla et al. (2012) find it to be 2–5 kpc in order to warp the bar out of the LMC disc plane. The centre of our half-ring structure can be conjectured as the location of the impact between the SMC and LMC's disc (see e.g. Mapelli & Mayer 2012). In Section 4.3, we determine the physical distance between the LMC's centre and that of the five-star ring to be  $2.6 \pm 0.3$  kpc, a value that is towards the low end of the estimated impact parameter from models of the interaction. The uncertainty in this value is derived from  $\delta r$  (see Section 4.3), i.e. it is based on the fit of the five LMC stars. Nevertheless, our observations bring to light an as yet unknown aspect of the LMC–SMC interaction: the formation of young stars in a half-ring configuration in the outskirts of LMC's disc. These observations should be useful in constraining the geometry of the collision from a totally new viewpoint, as well as provoking models of star formation in low gas environments.

## 5 SUMMARY

We have measured radial velocities and stellar parameters for 31 stars from the CD12 sample, candidates for being young OB stars in the periphery of the LMC; specifically between  $6^\circ$  and  $33^\circ$  from the LMC's centre. 21 of our targets are foreground subdwarfs or white dwarfs, and one is an object of uncertain classification, likely a post-HB star. The remaining nine targets are young, hot, MS stars, whose distances and ages are determined in our study. We find that three of them are foreground to the LMC, probably runaway stars from the Galactic disc, while six of our targets are recently formed stars at the distance of the LMC. The youngest of these (11 Myr) is a member of a known stellar association, ICA76, located in the Bridge. The remaining five stars span a narrow age range between 18 and 45 Myr. These five MS objects are located between  $\sim 8^\circ$  and  $13^\circ$  from LMC's centre in regions of low H I density, and appear to have formed in isolation. Their radial velocities follow the H I gas motions in the Magellanic system, and are not too deviant from the predictions of an LMC disc model. Specifically, velocity differences with the model are on average  $36 \text{ km s}^{-1}$ , which imply a travel distance of at most 1 kpc ( $\sim 1^\circ$  at the LMC's distance) during their lifetime. We argue that these stars were formed *in situ* rather than being runaways from the inner disc of the LMC. Their spatial configuration appears ring-like, and is suggestive of a formation process triggered by the recent collision of the SMC with the LMC disc at a modest, 4-kpc impact parameter. Highly accurate trigonometric parallaxes and absolute proper motions for these young stars, as is expected within the next few years from the *Gaia* mission, should readily confirm or refute this possible explanation. Regardless of their apparent geometric configuration, these newly identified LMC members represent a feature that, unlike previously discovered structures in the periphery of the disc, is traced by young stars.

## ACKNOWLEDGEMENTS

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