

# Optical spectroscopy of blazars for the Cherenkov Telescope Array<sup>★</sup>

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## ABSTRACT

**Context.** Blazars are the most numerous class of high-energy (HE;  $E \sim 50$  MeV–100 GeV) and very high-energy (VHE;  $E \sim 100$  GeV–10 TeV) gamma-ray emitters. Currently, a measured spectroscopic redshift is available for only about 50% of gamma-ray BL Lacertae objects (BL Lacs), mainly due to the difficulty in measuring reliable redshifts from their nearly featureless continuum-dominated optical spectra. The knowledge of the redshift is fundamental for understanding the emission from blazars, for population studies and also for indirect studies of the extragalactic background light and searches for Lorentz invariance violation and axion-like particles using blazars.

**Aims.** This paper is the first in a series of papers that aim to measure the redshift of a sample of blazars likely to be detected with the upcoming Cherenkov Telescope Array (CTA), a ground-based gamma-ray observatory.

**Methods.** Monte Carlo simulations were performed to select those hard spectrum gamma-ray blazars detected with the *Fermi*-LAT telescope still lacking redshift measurements, but likely to be detected by CTA in 30 hours of observing time or less. Optical observing campaigns involving deep imaging and spectroscopic observations were organised to efficiently constrain their redshifts. We performed deep medium- to high-resolution spectroscopy of 19 blazar optical counterparts with the Keck II, SALT, and ESO NTT telescopes. We searched systematically for spectral features and, when possible, we estimated the contribution of the host galaxy to the total flux.

**Results.** We measured eleven firm spectroscopic redshifts with values ranging from 0.1116 to 0.482, one tentative redshift, three redshift lower limits including one at  $z \geq 0.449$  and another at  $z \geq 0.868$ . Four BL Lacs show featureless spectra.

**Key words.** galaxies: active – BL Lacertae objects: general – gamma rays: galaxies – galaxies: distances and redshifts

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## 1. Introduction

Blazars, a subclass of radio-loud active galactic nuclei (AGN), display the most extreme observational properties of all types of AGN. They show unique characteristics such as beamed non-thermal emission from the radio to gamma-rays, strong ( $\geq 3\%$ ) polarised optical (see e.g. Angel & Stockman 1980; Angelakis et al. 2016) and radio (Lister et al. 2011) emission, variability from a few percent up to a few orders of magnitude on different timescales at all wavelengths (see e.g. Wagner & Witzel 1995; Falomo et al. 2014), and for some, on the Very Long Baseline Interferometry scale, the presence of superluminal radio blobs (see e.g. Vermeulen & Cohen 1994). These characteristics are generally explained by strong Doppler amplification of the jet emission with Lorentz factors of up to  $\sim 40$  (e.g. Jorstad et al. 2017), the latter being seen at a very small angle ( $\theta \leq 5$  deg) with respect to the jet axis.

Blazars are classified into flat-spectrum radio quasars (FSRQs) and BL Lacertae objects (BL Lacs), an important distinctive feature being the presence of broad (a few thousand  $\text{km s}^{-1}$ ) and luminous ( $\geq 10^{42} \text{ erg s}^{-1}$ ) emission lines in the optical spectra of the former and their weakness or even absence in the latter. The strong emission lines in FSRQs are produced by ionised gas in the vicinity of the black hole. Their presence may be an indication that the accretion regime in FSRQs is different with respect to that in BL Lacs (e.g. Ghisellini et al. 2011). Both blazar subtypes present two broad distinct components in their spectral energy distribution (SED), the low-energy component peaking in the infrared to X-ray band and the high-energy component peaking in the MeV to TeV band. The lower energy component is due to synchrotron radiation, while the higher energy one is generally ascribed to inverse Compton upscattering of lower energy photons off the population of accelerating electrons in the jet, though a hadronic component may also be responsible for the second SED peak (see e.g. Mücke & Protheroe 2001; Cerruti et al. 2015). The location of the low-energy peak can be used to subdivide BL Lacs into four different classes (see e.g. Padovani & Giommi 1995; Costamante et al. 2001). If the peak is at a frequency lower than  $10^{14}$  Hz, the BL Lac object is a low-frequency peaked BL Lac object (LBL); if it is between  $10^{14}$  and  $10^{15}$  Hz it is an intermediate-frequency peaked BL Lac object (IBL); if it is between  $10^{15}$  and  $10^{17}$  Hz it is a high-frequency peaked BL Lac object (HBL); and finally if it is at a frequency higher than  $10^{17}$  Hz, it is an extreme high-frequency peaked BL Lac object (EHBL). A slightly different classification scheme, defined in Abdo et al. (2010), is used in the Third Fermi High Energy Catalog (3FHL) (Ajello et al. 2017) where both FSRQs and BL Lacs are divided into low-synchrotron peaked (LSP, equivalent to LBL for BL Lacs), intermediate-synchrotron peaked (ISP, equivalent to IBL for BL Lacs), and high-synchrotron peaked (HSP, equivalent to HBL and EHBL for BL Lacs).

In the HE regime, covered since 2008 by the Large Area Telescope (LAT) on board the *Fermi* satellite, blazars account for more than 3400 of  $\sim 5000$  ( $\sim 68\%$ ) detected sources, as reported in the Data Release 2 of the Fourth *Fermi* Large Area Telescope catalogue of Gamma-Ray sources (4FGL-DR2) (Ballet et al. 2020); 730 of them are FSRQs, 1190 are BL Lacs, and 1517 are blazar candidates of uncertain types (BCUs). In the VHE regime the current generation of Imaging Atmospheric Cherenkov Telescopes (IACTs) (H.E.S.S.<sup>1</sup>, MAGIC<sup>2</sup>, and VERITAS<sup>3</sup>), has

detected 75 blazars, about one-half of the total number of identified sources<sup>4</sup>. Most of these blazars, 64 out of 75, are BL Lacs. Of these 64 TeV BL Lacs, we note that 13 still lack spectroscopic redshift values.

In the next few years a new facility, the Cherenkov Telescope Array (CTA<sup>5</sup>) will become operational with a northern site in the Canary Islands (Spain) and a southern site in the Atacama desert (Chile). With a 20 GeV–300 TeV energy range and a sensitivity approximately ten times better than the current generation of IACTs, it is expected to detect hundreds of blazars according to current estimates, thus opening the possibility of population studies with a significant sample size (Cherenkov Telescope Array Consortium 2019).

The CTA will enable substantial progress on gamma-ray population studies by deepening existing surveys, and will permit more detailed studies of the VHE gamma-ray emission and its origin (Sol et al. 2013). Furthermore, the stellar optical and near-infrared (NIR) radiation, called the extragalactic background light (EBL) (see e.g. Hauser & Dwek 2001; Biteau & Williams 2015; H.E.S.S. Collaboration 2017; Fermi-LAT Collaboration 2018; Abeyssekara et al. 2019; Acciari et al. 2019), acts as a source of opacity for the gamma rays from blazars. This effect can be used to derive within the same fitting process the properties of the EBL from its imprint on the VHE spectrum of blazars and the intrinsic VHE spectra of these sources (see e.g. Domínguez & Ajello 2015). The propagation of VHE gamma-ray radiation can be used to investigate topics related to cosmology or fundamental physics: the properties of the intergalactic magnetic field (IGMF) (see e.g. Ackermann et al. 2018; Aharonian et al. 1994; Alves Batista et al. 2019), the possible existence of axion-like particles (see e.g. Mirizzi et al. 2007; de Angelis et al. 2011; Abramowski et al. 2013), an independent measurement of the Hubble constant  $H_0$  (Salamon et al. 1994), or the search for Lorentz Invariance Violation (Kifune 1999). More details on these subjects can be found in (Cherenkov Telescope Array Consortium 2019, and references therein).

The capabilities of CTA to advance these subjects are presented in Cherenkov Telescope Array Consortium (2021). Finally, another important science case for the measurement of the redshifts of BL Lacs is the first evidence of neutrino emission from these sources (IceCube Collaboration 2018a,b; Franckowiak et al. 2020; Giommi et al. 2020; Paliya et al. 2020). A precise estimation of the total luminosity is necessary to fully understand the role of hadrons in the jet, which implies the knowledge of the redshift (see e.g. Paiano et al. 2018).

For BL Lacs these exciting possibilities are hampered by the difficulty in measuring reliable redshifts from their nearly featureless, continuum-dominated optical spectra. In optical spectroscopy one of the distinctive properties of BL Lacs is that they are objects with weak emission lines. The limit is usually set at an equivalent width (EW) of 5 Å (Urry & Padovani 1995), but it is known that the emission lines of BL Lacs can sometimes be brighter than that (see e.g. Stickel et al. 1991). High signal-to-noise (S/N) spectra are usually needed to detect these weak lines. For this reason the measurement of the redshifts of BL Lacs is a challenging task; spectroscopic observations are often unsuccessful and a large fraction of BL Lacs lack redshifts. Recognising this, several spectroscopic campaigns

<sup>1</sup> <https://www.mpi-hd.mpg.de/hfm/HESS/>

<sup>2</sup> <https://magic.mpp.mpg.de>

<sup>3</sup> <https://veritas.sao.arizona.edu>

<sup>4</sup> <http://tevcat.uchicago.edu>

<sup>5</sup> <https://www.cta-observatory.org>

to measure the redshifts of gamma-ray BL Lacs detected by *Fermi*-LAT have been organised. An early extensive effort was made by [Shaw et al. \(2013\)](#), who reported on rather deep observations from the southern and northern sites of 372 BL Lacs. Adding previous literature results, they obtained a sample with 44% redshift completeness whose median redshift is  $z_{\text{med}} = 0.33$ . Since then new *Fermi*-LAT source catalogues have been published, the latest being the 4FGL-DR2 ([Ballet et al. 2020](#)), and new associations with blazars have been produced (see e.g. [Acero et al. 2013](#); [Arsioli et al. 2015](#); [Arsioli & Chang 2017](#); [Kaur et al. 2019](#)). Several groups have performed spectroscopic campaigns often focused on these new *Fermi*-LAT BL Lacs and BCUs. One group (see e.g. [Paiano et al. 2017a](#); [Landoni et al. 2018](#); [Paiano et al. 2020](#)) has pursued high S/N observations mainly from the Gran Telescopio Canarias (GTC) of different subsamples of gamma-ray blazars selected for being unidentified or for being likely to be detected at very high energies from simple estimations. A very extensive series of papers (see e.g. [Paggi et al. 2014](#); [Massaro et al. 2013, 2015a,b](#)) has pursued the identification and redshift measurement of *Fermi*-LAT BCUs with low- and medium-sensitivity observations from both hemispheres. Other less extensive campaigns with similar aims, which have targeted smaller samples, include those of [Masetti et al. \(2013\)](#), [Marchesini et al. \(2016\)](#), and [Klindt et al. \(2017\)](#). A synthesis of these recent observations has been presented by [Peña-Herazo et al. \(2020\)](#) for a total of 416 BL Lacs and BCUs, 311 taken from their own observations and 105 from recent literature including the ones cited above. Only about 30% of these objects have spectroscopic redshift values, and their median redshift is  $z_{\text{med}} = 0.285$ .

This incompleteness in redshift determination implies that it is very difficult to determine the properties of blazars as a population. A fundamental quantity such as luminosity is not determined for more than half of them. As a consequence, the blazar sequence (i.e. the observation that the peak frequency of the blazar SED becomes redder with increasing peak luminosity) is still a very controversial subject. It has been interpreted as being due to differences in radiative cooling among blazar classes ([Ghisellini et al. 2017](#)) or to selection effects. [Giommi & Padovani \(2015\)](#) postulate that blazars missing redshifts should be mostly high-luminosity HSP objects, in contrast to the fact that the low-energy SED peak correlates with luminosity. Therefore, measuring the redshift of a sizeable fraction of them could test this hypothesis by allowing the luminosity to be measured.

BL Lac redshifts can be estimated under the assumption that the host galaxy is a standard candle. The studies of BL Lac host galaxies have shown that BL Lacs are hosted in giant elliptical galaxies with absolute magnitude distribution well fitted by a Gaussian peaked at  $M_R \sim -22.8$  with FWHM of 1 mag (see e.g. [Sbarufatti et al. 2005](#), and references therein). Therefore the redshift can be estimated either from optical images (see e.g. [Falomo 1996](#); [Falomo & Kotilainen 1999](#); [Nilsson et al. 2003](#)) or from optical spectra ([Sbarufatti et al. 2006](#)). The non-detection of the host galaxy or of its absorption features allows a lower limit to be set on the redshift of the source (see review by [Falomo et al. 2014](#), for further discussion and references). A photometric method to derive limits on the redshifts of BL Lacs ([Rau et al. 2012](#)) is based on the absorption of UV photons from a BL Lac object by the neutral hydrogen along our line of sight causing a clear attenuation in the flux at the Lyman limit (912 Å). This dropout can be successfully used to measure the redshift of the BL Lac object if it is located at high redshift ( $\geq 1.3$ ). More frequently for spectroscopy, if no intrinsic spectral fea-

ture is detected, a firm lower limit on the redshift may be set by the detection of an absorption system (usually MgII doublets with wavelengths  $\lambda_1 = 2796.3 \text{ \AA}$  and  $\lambda_2 = 2803.5 \text{ \AA}$ ) along the line of sight towards the source. It should also be noted that in certain cases conflicting redshift values are reported in the literature even for objects for which high S/N spectra have been obtained. For example, the redshift of 1ES 0502+675 has been reported as  $z = 0.416$  ([Landt et al. 2002](#)) and as  $z = 0.314$  ([Scarpa et al. 1999](#)). Similarly, the redshift of PMN J0816–1311 has been reported as  $z = 0.046$  ([Jones et al. 2004, 2009](#)) and as  $z > 0.288$  ([Pita et al. 2014](#)). Details on these and other cases are reported in Appendix A.

Gamma-ray blazars are the main extragalactic targets for CTA and high-confidence spectroscopic redshifts are needed for them. The planning of CTA observations by the CTA Consortium is currently under way. It is therefore of great importance to start acquiring highly reliable redshifts for a large fraction of the AGN sources detected with *Fermi*-LAT that are likely to be detected with CTA. Such a redshift measurement campaign is recognised as necessary support for the CTA Key Science Programme (KSP) on AGN ([Cherenkov Telescope Array Consortium 2019](#)). We thus initiated this redshift-measuring campaign by carrying out observations at different facilities to which we have access, and we report the first results here.

This paper is organised in the following way: the sample selection is presented in Sect. 2; the observing strategy in Sect. 3; the observations, data reduction, and analysis in Sects. 4–6; and the discussion and conclusions in Sect. 7. For all calculations, we used a cosmology with  $\Omega_M = 0.27$ ,  $\Omega_\Lambda = 0.73$ , and  $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ . All wavelengths are in air. All magnitudes are in the AB system.

## 2. Sample selection

The CTA will detect several hundreds of blazars in the VHE band ([Cherenkov Telescope Array Consortium 2019](#)). It is expected that many of these blazars also emit gamma rays at lower energies, in the energy range currently covered by the *Fermi*-LAT. Therefore, it is possible to use the *Fermi*-LAT catalogues to identify a population of blazar candidates for CTA. The 3FHL catalogue ([Ajello et al. 2017](#)) is particularly interesting because it contains the spectral information, averaged over 7 years of its all-sky survey, for the harder and brighter sources detected by *Fermi*-LAT. The catalogue considers only photons above 10 GeV, which is very close to the energy threshold of CTA, and contains 1556 sources. The vast majority of them (1212) are blazars. The 3FHL blazars comprise 172 FSRQs, 750 BL Lacs, and 290 BCUs. A redshift value, if available, is provided for each source, but the information on its origin is not given. Conversely spectroscopic lower limits are not part of the catalogue. Among blazars in the 3FHL catalogue 95% of FSRQs have a known redshift. Conversely 46% of the BL Lacs and only 10% of BCUs have a known redshift. We therefore focused on the 1040 BL Lacs and BCUs, of which only 373 (36%) have a redshift in the 3FHL catalogue.

We performed Monte Carlo simulations using the Gammapy<sup>6</sup> software ([Deil et al. 2017](#); [Nigro et al. 2019](#)) to estimate the minimal observation time necessary to detect at  $5\sigma$  each of these 1040 3FHL BL Lacs and BCUs with the North or South CTA array, depending on the declination of the

<sup>6</sup> <https://gammapy.org>

source. We used publicly available CTA performance files<sup>7</sup>. For each source, the average energy spectrum reported in 3FHL was extrapolated to very high energies and an intrinsic exponential cutoff at 3 TeV in the comoving frame was assumed in order to simulate the spectral curvature expected at these energies (Cherenkov Telescope Array Consortium 2021)<sup>8</sup>. To take into account the energy and redshift-dependent absorption of gamma rays due to their interaction with the EBL, the spectral model was multiplied by  $\exp(-\tau(E, z))$ , where  $\tau(E, z)$  is the gamma-gamma optical depth provided by Domínguez et al. (2011),  $E$  is the gamma-ray energy, and  $z$  the source redshift from the 3FHL catalogue. For sources without a reported redshift in 3FHL, a value of  $z=0.3$ , similar to  $z_{\text{med}}=0.33$  (Shaw et al. 2013) and to  $z_{\text{med}}=0.285$  (Peña-Herazo et al. 2020) for BL Lacs, was considered.

Sources were selected in a two-step process. In the first step a literature review or the analysis of publicly available archived spectra was performed for the 221 sources expected to be detected with CTA in less than 50 h from the simulations defined above. This condition allowed us to reduce the number of sources for which the literature review was needed. During this check we examined published results and publicly available spectra for the selected sources. Among the sources having a reported redshift in 3FHL, 13 incorrect or unreliable redshift values were identified. These values were discarded because either we could not identify the features in the publicly available spectra or because they were contradicted or not confirmed by later spectra with much higher S/N (see Appendix A for details). In 2 out of the 13 cases, we associated instead published spectroscopic lower limits. Following the same procedure, among sources with no redshift reported in the 3FHL catalogue, a reliable spectroscopic redshift was assigned for 7 sources and a spectroscopic lower limit was associated with 12 sources. This process resulted in the revision of 32 redshift values (see the list of sources in Appendix A). We note that the new lower limit values we obtained were adopted as redshifts whenever they were greater than our chosen value  $z=0.3$ ; conversely, for smaller values we used  $z=0.3$ . In the second step, using the revised redshift values, simulations were reprocessed. We then selected the sample of 165 sources without redshift measurement that are expected to be detectable in less than 30 h if they are in the average spectral state reported by 3FHL, and in a significantly lower time if they are in a flaring state.

As an early effort to determine the redshift of the sources in this sample, we extracted the 19 sources that we observed for this paper. The criteria used for this selection are explained in the next section.

### 3. Observing strategy

We discuss here the strategy employed in deriving our observing campaign, the results of which are reported for the first time in this paper. The goal is to obtain spectroscopic redshifts or lower limits for the highest possible number of sources in the sample. We plan to release our results as they become available so that they can be used to update the CTA Consortium observing programme (Cherenkov Telescope Array Consortium 2019). These

<sup>7</sup> <https://www.cta-observatory.org/wp-content/uploads/2019/04/CTA-Performance-prod3b-v2-FITS.tar.gz>

<sup>8</sup> However, in most cases the exact value of this cutoff has only a marginal effect on the estimation of the observation time required to reach a  $5\sigma$  detection as the detection significance is dominated by low-energy events.

observations will also serve the astronomical community at large as having confirmed redshifts for a larger sample of sources will help scientists who are investigating the properties of blazars and of their emission.

To pursue this goal we have devised an observing programme aimed at constraining the redshift of these sources through deep imaging and spectroscopic observations. The goal of the imaging observations is to search for the extension of the source profile due to the host galaxy, and the first results will be reported in a follow-up paper (Fallah Ramazani et al., in prep.). The goal of the spectroscopic observations is to search for stellar absorption features of the host galaxy that are usually overwhelmed by the non-thermal continuum of the jet. As the host galaxies are usually luminous ellipticals (Urry et al. 2000), the main features that we expect are the CaHK doublet, Mgb and NaID. Emission lines (especially [OII], [OIII], H $\alpha$ , and N[II]) are only rarely detected. In all cases, EWs of about 5 Å or less are expected. To reach this goal we require that each spectrum has a spectral resolution  $\lambda/\Delta\lambda$  of at least of a few hundred (if possible  $\sim 1000$ ), and an average S/N of  $\sim 100$  per pixel. The combination of these two constraints is extremely powerful. On the one hand, the imaging detection of the host galaxy is a clear indication of the likelihood of obtaining a redshift measurement. It has been shown for a sample of 100 X-ray detected BL Lacs that to date 90% of the 62 targets with a detected host galaxy (Nilsson et al. 2003) have a spectroscopic redshift from spectroscopic programmes (see e.g. Peña-Herazo et al. 2020; Paiano et al. 2020), while more than 80% of the unresolved sources still do not have redshift values. On the other hand, spectra at  $S/N \sim 100$  and resolution  $\sim 1000$  allow the detection of weak host-galaxy features with EWs smaller than 5 Å and of intervening absorption systems (see e.g. Pita et al. 2014). If the instrument we are using cannot provide us with spectra having both of these properties, we choose configurations that allow us to obtain at least one of them.

We also performed a comprehensive literature search on our targets looking for previous spectroscopic results and for evidence of extension in archival and published data (e.g. the Two Micron All sky Survey (2MASS) Extended Source Catalogue; Skrutskie et al. 2006). The results were classified in terms of reliability based on available information (images, plots, data). A source with low S/N spectroscopy and a tentative redshift value is a high-priority target. This allows us to concentrate our early efforts on promising and relatively uninvestigated sources. Conversely, if we find that a source already has at least one deep and featureless spectrum and/or is not extended, it is classified as a low-priority target. A possible option for these sources is to trigger a spectroscopic observation during an epoch of low optical activity in order to take advantage of the improved S/N due to the lower non-thermal foreground.

### 4. Observations and data reduction

Data were collected on 19 blazars using three different instruments at three facilities for a total observation time of about 17.5 h between May 2018 and November 2019. Observations were performed using the Echelle Spectrograph and Imager (ESI; Sheinis et al. 2002) installed on the Keck II telescope at the Keck observatory, with the Robert Stobie Spectrograph (RSS; Burgh et al. 2003) on the Southern African Large Telescope (SALT) at the South African Astronomical Observatory and with the ESO Faint Object Spectrograph and Camera (EFOSC2; Buzzoni et al. 1984) on the New Technology Telescope at La Silla Observatory. The Keck II and SALT telescopes respectively

**Table 1.** List of observed sources and parameters of the observations.

3FHL name	4FGL name	Source name	Ext.	RA	Dec	Telescope/ instrument	Slit (")	Start time UTC	Exp. (s)	Airm. (11)	Seeing (")
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)
3FHL J0114.9–3359	4FGL J0114.9–3400	1RXS J011501.3–340008 <sup>†</sup>	N	01 15 01.6	–34 00 27	NTT/EFOSC2	1.5	2019-06-23 08:45:38	5400	1.23	0.8
3FHL J0156.7–5302	4FGL J0156.9–5301	1RXS J015658.6–530208 <sup>†</sup>	N	01 56 58.0	–53 01 60	SALT/RSS	2.0	2019-11-24 23:01:16	2250	1.33	1.2
3FHL J0156.7–5302	4FGL J0156.9–5301	1RXS J015658.6–530208 <sup>†</sup>	N	01 56 58.0	–53 01 60	SALT/RSS	2.0	2019-11-26 22:34:09	2250	1.27	1.4
3FHL J0209.3–5229	4FGL J0209.3–5228	1RXS J020922.2–522920 <sup>†</sup>	N	02 09 21.6	–52 29 23	SALT/RSS	2.0	2019-12-23 21:15:36	2220	1.28	1.4
3FHL J1443.9–3908	4FGL J1443.9–3908	PKS 1440–389 <sup>†</sup>	Y	14 43 57.2	–39 08 40	NTT/EFOSC2	1.5	2019-06-22 23:21:10	4500	1.10	1.1
3FHL J1457.8–4642	4FGL J1457.8–4642	PMN J1457–4642	Y	14 57 41.8	–46 42 10	NTT/EFOSC2	1.5	2019-06-23 23:15:26	2700	1.18	1.6
3FHL J1511.8–0513	4FGL J1511.8–0513	NVSS J151148–051345	N	15 11 48.5	–05 13 47	NTT/EFOSC2	1.5	2019-06-23 02:03:48	2700	1.10	0.9
3FHL J1520.7–0348	4FGL J1520.8–0348	NVSS J152048–034850 <sup>†</sup>	N	15 20 48.9	–03 48 51	NTT/EFOSC2	1.5	2019-06-23 03:00:00	2700	1.15	1.0
3FHL J1532.7–1319	4FGL J1532.7–1319	TXS 1515–273	Y	15 17 59.5	–27 32 51	NTT/EFOSC2	1.5	2019-06-23 00:51:57	3600	1.03	0.7
3FHL J1539.7–1127	4FGL J1539.7–1127	PMN J1539–1128	N	15 39 41.2	–11 28 35	NTT/EFOSC2	1.5	2019-06-24 00:19:35	5400	1.14	1.6
3FHL J1548.4+1456	4FGL J1548.3+1456	WISE J154824.39+145702.8	Y	15 48 24.4	+14 57 03	Keck/ESI	1.0	2018-05-13 10:34:13	3600	1.01	0.6
3FHL J1637.8–3448	4FGL J1637.8–3449	NVSS J163750–344915	N	16 37 51.0	–34 49 15	NTT/EFOSC2	1.5	2019-06-24 02:01:38	2700	1.04	0.7
3FHL J1838.8+4802	4FGL J1838.8+4802	GB6 J1838+4802 <sup>†</sup>	Y	18 38 49.2	+48 02 34	Keck/ESI	1.0	2018-05-13 12:28:22	3900	1.15	0.6
3FHL J1841.3+2909	4FGL J1841.3+2909	MITG J184126+2910	Y	18 41 21.7	+29 09 41	Keck/ESI	1.0	2018-05-13 11:41:46	2400	1.11	0.5
3FHL J1842.4–5841	4FGL J1842.4–5840	1RXS J184230.6–584202	N	18 42 29.8	–58 41 56	NTT/EFOSC2	1.5	2019-06-23 04:12:40	4500	1.17	1.0
3FHL J1958.3–3011	4FGL J1958.3–3010	1RXS J195815.6–301119 <sup>†</sup>	Y	19 58 14.9	–30 11 11	NTT/EFOSC2	1.5	2019-06-23 05:38:55	1800	1.03	1.0
3FHL J2001.2+4353	4FGL J2001.2+4353	MAGIC J2001+435 <sup>†</sup>	Y	20 01 12.9	+43 52 53	Keck/ESI	1.0	2018-05-13 13:38:54	3480	1.13	0.6
3FHL J2036.9–3328	4FGL J2036.9–3329	1RXS J203650.9–332817	N	20 36 49.5	–33 28 30	NTT/EFOSC2	1.5	2019-06-25 08:07:13	3600	1.04	0.8
3FHL J2131.0–2746	4FGL J2131.0–2746	RBS 1751 <sup>†</sup>	N	21 31 03.3	–27 46 58	NTT/EFOSC2	1.5	2019-06-25 09:26:18	3600	1.07	1.5
3FHL J2324.7–4040	4FGL J2324.7–4041	IES 2322–409 <sup>†</sup>	N	23 24 44.7	–40 40 49	NTT/EFOSC2	1.5	2019-06-23 06:18:39	7200	1.31	0.8

**Notes.** All sources were observed once with the exception of 1RXS J015658.6–530208, which was observed twice. The NTT/EFOSC2 spectra were all taken with Gr6, except for NVSS J151148–051345 (Gr 14) and NVSS J152048–034850 (Gr 8). The sources with a † symbol are listed in the BZCAT catalogue (Massaro et al. 2015c). The columns are (1) 3FHL name, (2) 4FGL name, (3) Source name, (4) Extension flag, (5) Right ascension (J2000), (6) Declination (J2000), (7) Telescope and instrument, (8) Slit width in arcsec, (9) Start time of the observations, (10) Exposure time, (11) Average airmass, and (12) Average seeing.

have 10 m and 11 m diameter primary mirrors, while the NTT primary mirror is substantially smaller (3.5 m in diameter). The list of observed sources together with the details of the observations are given in Table 1.

#### 4.1. Keck/ESI

The ESI spectrograph is a visible-wavelength faint-object imager and single-slit spectrograph; it has been in operation at the Cassegrain focus of the Keck II telescope since 1999. We used it in its main spectroscopy mode, the echellette mode, which has a single-shot wavelength coverage of 3900–10 000 Å, a throughput up to 28%, and spectral resolution  $\lambda/\Delta\lambda \sim 10\,000$ . Two short observations of the flux standard HD165459 with slit widths of 1 and 6 arcsec were performed at the end of the night to allow for flux calibrations.

The data reduction was performed using the XIDL pipeline<sup>9</sup> based on the Interactive Data Language (IDL<sup>10</sup>) software. We used the pipeline to perform bias subtraction, flat-field division, wavelength calibrations, cosmic ray subtraction and spectral extraction. We then performed order merging and flux calibration using our own procedures under IDL. Telluric corrections were performed using molecfit (Smette et al. 2015; Kausch et al. 2015). The spectra were dereddened using the maps of Schlafly & Finkbeiner (2011) and the extinction curve of Fitzpatrick (1999).

#### 4.2. SALT/RSS

The RSS is SALT’s main instrument; it is a complex multi-mode instrument with a wide range of capabilities. We used it in Long Slit Spectroscopy (LSS) mode with the PG0900

grating and a 2 arcsec slit. This configuration is sensitive between 4500 and 7500 Å and it has a throughput greater than 20% (Kobulnicky et al. 2003).

We reduced the spectra using PySALT (Crawford et al. 2010) accounting for cross-talk, bias, gain, and flat-field correction. Wavelength calibration was performed using standard IRAF routines (Tody 1986), while cosmic ray cleaning and flux calibration were performed under IDL. To estimate the spectral resolution we extracted the sky spectrum and fitted ten isolated sky lines across the spectrum with Gaussian functions. This analysis produced an approximate value of the spectral resolution  $\lambda/\Delta\lambda$  of about 1000. To perform flux calibration, we used observations of the standard star HILT600 taken on November 26, 2019, with a 4 arcsec slit. These data were reduced in the same way as the data of the target. Other standard star observations taken near the end of December 2019 were of much lower quality and were therefore discarded. If more than one high-quality observation was obtained, an average spectrum was produced. Telluric and reddening corrections were performed as described in Sect. 4.1.

SALT has a moving, field-dependent, and under-filled entrance pupil, which makes absolute flux calibration difficult to achieve to a good degree of accuracy (see e.g. Buckley et al. 2018). We therefore obtained near-contemporary photometric observations for both targets.

Our first target, 1RXS J015658.6–530208, was observed using the Ultraviolet/Optical Telescope (UVOT; Roming et al. 2005, 170–600 nm) on board the *Neil Gehrels Swift* Observatory (Gehrels et al. 2004) on November 30, 2019. The UVOT instrument observed in the optical ( $u$ ,  $b$ , and  $v$ ) photometric bands (Poole et al. 2008; Breeveld et al. 2010) with exposures of 305 s, 300 s, and 305 s for the  $u$ ,  $b$ , and  $v$  filter, respectively. We analysed the data using the uvotsource task included in the HEASoft package (v6.22). Source counts were extracted from a circular region of 5 arcsec radius centred on the source,

<sup>9</sup> <https://github.com/profxj/xidl>

<sup>10</sup> <http://www.harrisgeospatial.com>

**Table 2.** Equivalent widths in Å of the absorption features detected in the spectra at the measured redshift for each source.

Source name (1)	CaHK (2)	CaIG (3)	Mgb (4)	CaFe (5)	NaID (6)
1RXS J011501.3–340008	$6.4 \pm 1.6$	$2.5 \pm 0.7$	–	–	–
1RXS J015658.6–530208	$2.3 \pm 0.3$	$0.8 \pm 0.1$	$0.9 \pm 0.2$	–	–
1RXS J020922.2–522920	$1.9 \pm 0.3$	$0.5 \pm 0.1$	$0.5 \pm 0.2$	–	$0.5 \pm 0.1$
PKS 1440–389	$0.6 \pm 0.1$	$0.5 \pm 0.1$	–	–	$0.4 \pm 0.1$
PMN J1457–4642	$11.8 \pm 0.8$	$4.8 \pm 0.7$	$11.0 \pm 1.2$	–	$6.8 \pm 0.9$
TXS 1515–273	$1.2 \pm 0.2$	$0.3 \pm 0.1$	$0.3 \pm 0.1$	–	$0.7 \pm 0.2$
WISE J154824.39+145702.8	$25.0 \pm 0.9$	$7.2 \pm 0.3$	$6.5 \pm 0.3$	$2.5 \pm 0.3$	$4.8 \pm 0.2$
MITG J184126+2910	$1.6 \pm 0.3$	$0.8 \pm 0.2$	$0.7 \pm 0.2$	$0.6 \pm 0.2$	–
1RXS J184230.6–584202	$3.6 \pm 1.0^*$	–	–	–	–
1RXS J195815.6–301119	$3.3 \pm 0.3$	$1.2 \pm 0.3$	$5.3 \pm 0.6$	$2.4 \pm 0.2^*$	$5.5 \pm 0.4$
MAGIC J2001+435	$2.2 \pm 0.2$	–	$1.0 \pm 0.2$	$0.8 \pm 0.1$	$0.4 \pm 0.2$
1ES 2322–409	$0.6 \pm 0.1$	$1.2 \pm 0.2$	$2.3 \pm 0.3$	–	$1.2 \pm 0.3$

**Notes.** The CaFe feature of 1RXS J195815.6–301119 is likely contaminated by Galactic NaID. The columns are (1) Source name, (2) Equivalent width of the CaHK feature with errors, (3) Equivalent width of the CaIG feature with errors, (4) Equivalent width of the Mgb feature with errors, (5) Equivalent width of the CaFe feature with errors, (6) Equivalent width of the NaID feature with errors. If the feature is not detected, the legend is ‘–’. The detection of CaHK in 1RXS J184230.6–584202 is uncertain and it is flagged with an asterisk.

while background counts were derived from a circular region of 20 arcsec radius in a nearby source-free region. The results are presented in Table B.1.

The second target, 1RXS J020922.2–522920, was observed using the REM Optical Slitless Spectrograph (ROSS2) at the REM telescope (Zerbi et al. 2001; Covino et al. 2004), a robotic telescope located at the ESO Cerro La Silla observatory (Chile). With the ROSS2 instrument we obtained two 240 s integration images of the target in *g*, *r*, and *i* filters on five separate dates between the end of December 2019 and the beginning of January 2020. All raw optical frames obtained were reduced following standard procedures. Instrumental magnitudes were obtained via aperture photometry and absolute calibration was performed by means of secondary standard stars in the field reported by the American Association of Variable Star Observers Photometric All-Sky Survey (APASS) catalogue<sup>11</sup>. The results are shown in Table B.2.

#### 4.3. NTT/EFOSC2

The EFOSC2 is a versatile multimode instrument that is particularly efficient in low-resolution spectroscopy. In order to obtain a wide wavelength coverage with good sensitivity and reasonable resolving power we selected Grism 6 for 11 of our 13 targets. Grism 6 is sensitive in the range 3860–8070 Å, but with low spectroscopic resolution  $\lambda/\Delta\lambda \sim 400$ . For the remaining two, NVSS J151148–051345 and NVSS J152048–034850, we selected Grism 14 (3095–5085 Å;  $\lambda/\Delta\lambda \sim 550$ ) and Grism 8 (4320–6360 Å;  $\lambda/\Delta\lambda \sim 660$ ), respectively, which are more adapted to investigating previously reported detections of MgII absorbers in their spectra. The throughputs of the grisms<sup>12</sup> are between 20% and 30%. In all cases we used a 1.5 arcsec slit; the same slit was also used for standard stars.

The observations were performed during three nights from June 22 to June 25, 2019 (see Table 1), with variable atmospheric conditions. Data reduction was performed using the ESO/EFOSC2 pipeline version 2.3.3 and esorex version 3.13.2.

<sup>11</sup> <https://www.aavso.org/apass>

<sup>12</sup> <https://www.eso.org/sci/facilities/lasilla/instruments/efosc/images-/GrismAllEfficiency2004.jpg>

The pipeline performs bias subtraction, flat-field correction, and wavelength calibration using daytime calibration files. Cosmic ray subtraction was performed under IDL. Subsequently flux calibration was performed using the standard stars observed during the run. These steps were performed for each independent frame; we then averaged the extracted spectra to obtain the final flux calibrated spectra. During flux calibration we discovered distortions in the spectral shape of the sources, likely due to the presence of clouds. We corrected them using spectra of stars that were put in the slit of some of our targets. Telluric and reddening corrections were performed as described in Sect. 4.1.

## 5. Redshift measurement and estimation of the blazar total emission

The optical spectrum of a blazar is a combination of non-thermal jet emission, AGN activity (thermal and non-thermal), and stellar emission of the host galaxy, usually an elliptical (Urry et al. 2000). The jet emission has the form of a featureless power law  $f_\lambda \propto \lambda^\alpha$ , which, as discussed above, is often much stronger than the host galaxy emission, making the host spectral features undetectable. Simulations (Landt et al. 2002; Piranomonte et al. 2007) have shown that when the rest frame jet-to-galaxy ratio at 5500 Å is around 10 the features are already very difficult to detect.

For each source we carefully searched for absorption or emission features that could be used to measure the redshift. When a possible feature was found we checked for the presence of other possible features at the same redshift. We then analysed the features in the following way. The spectra were normalised with cubic splines and the flux of each pixel was integrated to determine the total EW of each line. The uncertainties were estimated by taking the square root of the quadratic sum of the error spectrum and taking into account the errors of the continuum placement (see the appendix in Sembach & Savage 1992). The results are shown in Tables 2 and 3. Only for WISE J154824.39+145702.8 did we measure the EWs of the rich emission line spectrum by Gaussian fitting (see Table 4).

We considered two factors to estimate the uncertainty on redshift measurements: uncertainties in wavelength calibration and

**Table 3.** Equivalent width in Å of the main emission features detected in the spectra at the measured redshift.

Source name	[OII]	[OIII] $\lambda 5007$	H $\alpha$ -[NII]
(1)	(2)	(3)	(4)
PMNJ1457-4642	–	–	$-5.0 \pm 0.9$
TXS 1515-273	$-0.8 \pm 0.1$	$-0.8 \pm 0.1$	NA
1RXSJ195815.6-301119	–	–	$-4.8 \pm 1.0$
MAGIC J2001+435	–	–	$-0.4 \pm 0.1$

**Notes.** The EWs of the emission features of WISEJ1548.24.39+145702.8 are in Table 4. The columns are (1) Source name, (2) Equivalent width of the [OII] feature with errors, (3) Equivalent width of the [OIII]  $\lambda 5007$  feature with errors, (4) Equivalent width of the H $\alpha$ -[NII] complex with errors. If the feature is not detected, the legend is ‘–’.

uncertainties in the position of the detected features. The dispersion of the wavelength calibration in our spectra is always smaller than  $0.5 \text{ \AA}$  from  $\sim 4000$  to  $\sim 8000 \text{ \AA}$ , which translates into a relative precision smaller than  $6-12 \times 10^{-5}$  ( $18-36 \text{ km s}^{-1}$ ). Once the redshift was determined, we fitted Gaussian functions at the positions of the features found in each source listed in Tables 2 and 3, and took the variance of the fitted positions as the uncertainty. We then summed these with uncertainties in wavelength calibration and we obtained the total uncertainty estimates between  $2$  and  $8 \times 10^{-4}$  listed in Table 5.

After the redshift was determined, we modelled the spectrum with a combination of a power law describing the jet continuum and templates describing the elliptical galaxy emission (Mannucci et al. 2001; Bruzual & Charlot 2003), adding Gaussian emission features when needed (Pita et al. 2014). For simplicity we used only one template per spectrum. The fit was performed in the rest frame using the MPFIT software (Markwardt 2009) with two free parameters: the jet-to-galaxy ratio and the power-law slope. We estimated the goodness of fit from the value of the  $\chi^2_{\text{d.o.f.}}$ . There are systematic differences between the flux calibrated spectra and the models we used. This is a common occurrence as, in general, polynomials are added to the spectral models in order to obtain an acceptable match with the calibrated flux (e.g. Cappellari 2017). Given the weakness of our detected features in most cases we chose not to add polynomials as it would have led to overfitting. We thus added systematic errors to obtain error estimates. In some cases, the errors of the fit parameters were unphysically small, thus we independently fitted separate sections of the spectra and we estimated the errors from the differences between the resulting parameters. The results of these fits are presented in Table 5.

We also estimated the absolute magnitude of the detected host galaxies. To estimate the slit losses, we assumed the value of the effective radius of the host galaxy  $r_e$  as  $10 \text{ kpc}$  for a de Vaucouleurs profile. The only detected host for which we had a photometric estimate of the effective radius is MAGIC J2001+435 for which we used  $r_e = 6.8 \text{ kpc}$  (from the measured  $r_e = 2.4 \text{ arcsec}$ ). Within the uncertainties, the magnitude we obtain for this host is compatible with the photometric value quoted in Aleksić et al. (2014). Thus, we estimate that the uncertainties on the absolute magnitude are the same as the uncertainties on the measured magnitude. The K-corrections were computed from the template spectra and we did not apply evolutionary corrections. The absolute magnitudes can be found in Table 5.

**Table 4.** Equivalent widths (in Å) of the emission lines detected in WISEJ154824.39+145702.8.

Line	Equivalent width (Å)
(1)	(2)
[OII] $\lambda 3727$	$18.8 \pm 0.3$
H $\beta$	$\leq 1$
[OIII] $\lambda 4959$	$1.6 \pm 0.2$
[OIII] $\lambda 5007$	$4.8 \pm 0.2$
[OI] $\lambda 6300$	$2.2 \pm 0.7$
H $\alpha$	$4.5 \pm 0.4$
[NII] $\lambda 6548$	$4.4 \pm 0.1$
[NII] $\lambda 6583$	$13.3 \pm 0.1$
[SII] $\lambda 6716$	$4.1 \pm 0.3$
[SII] $\lambda 6731$	$4.4 \pm 0.5$

When the host galaxy was not detected, we fitted the spectrum with a power law with normalisation at the centre of the band and we estimated the errors fitting separate sections of the spectra as described above. The results are presented in Table 5.

## 6. Sources and results

In the following we discuss the results of our observations for each of the sources.

### 6.1. 1RXSJ011501.3-340008

1RXSJ011501.3-340008 was identified as an extreme HBL by Giommi et al. (2005). While its redshift is unknown in the 3FHL, Piranomonte et al. (2007) report its redshift to be  $z = 0.482$  on the basis of a one-hour EFOSC2 spectrum taken at the ESO 3.6 m telescope. The plot of the spectrum, shown in the appendix of their paper, shows a medium- to low-quality spectrum. To assess the reliability of this result we downloaded the public data of the observation to reduce them. The resulting spectrum has a general power-law shape and displays a spectral break around  $\lambda \sim 5900 \text{ \AA}$  that can be attributed to the CaHK feature at  $z \sim 0.482$ . However its S/N is only  $\sim 16$  per  $4 \text{ \AA}$  pixel. We tried to obtain a much longer integration with EFOSC2, but we were able to obtain only an integration of 1 h and 30 min (see Table 1). The resulting spectrum has a median  $S/N = 20$  and it has a general shape remarkably similar to the first with a spectral break at the same wavelength. Given the similarity of the two spectra, we averaged them to measure the properties of the total emission (Fig. 1, first row, left). We detect the CaHK feature at  $4\sigma$  and the CaIG feature at  $3.5\sigma$  both at redshift  $z = 0.4824 \pm 0.0007$ . Given the stability of the source spectrum and despite the relatively low S/N we consider that this is a firm redshift. The spectral fit gives a bright host galaxy magnitude  $M_R = -23.3 \pm 0.2$ . Although a fit with a local template (Mannucci et al. 2001) is satisfactory, the best results are obtained using a 2.5 Gyr Simple Stellar Population model (Bruzual & Charlot 2003). This suggests the possibility of an anomalous star formation history for this object; further analysis on this subject is beyond the scope of this paper.

**Table 5.** Analysis results for all the observed sources.

Source name	$S/N$	$R_c$ (BL Lac) (obs)	Redshift	Flux ratio	$R_c$ (gal) (fit)	$M_R$ (gal)	Slope
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)
1RXS J011501.3–340008	20	$19.4 \pm 0.1$	$0.4824 \pm 0.0007$	$0.9 \pm 0.2$	$19.7 \pm 0.3$	-23.3	$-1.8 \pm 0.4$
1RXS J015658.6–530208	100	$17.0 \pm 0.3$	$0.3043 \pm 0.0004$	$3.2 \pm 1.3$	$18.3 \pm 0.5$	-22.7	$-1.4 \pm 0.2$
1RXS J020922.2–522920	160	$15.4 \pm 0.2$	$0.2110 \pm 0.0002$	$6.0 \pm 1.8$	$17.1 \pm 0.4$	-23.2	$-1.5 \pm 0.2$
PKS 1440–389	230	$15.0 \pm 0.1$	$0.1385 \pm 0.0005$	$8.6 \pm 0.7$	$16.8 \pm 0.2$	-22.4	$-1.5 \pm 0.2$
PMN J1457–4642	45	$17.5 \pm 0.2$	$0.1116 \pm 0.0002$	$0.3 \pm 0.2$	$16.7 \pm 0.1$	-22.2	$-2.3 \pm 0.2$
NVSS J151148–051345	23	$17.5 \pm 0.1$ <sup>(†)</sup>	$\geq 0.4480 \pm 0.0003$	–	–	–	$-2.5 \pm 0.2$
NVSS J152048–034850	43	$17.2 \pm 0.1$	$\geq 0.8680 \pm 0.0002$	–	–	–	$-1.6 \pm 0.1$
TXS 1515–273	160	$15.6 \pm 0.1$	$0.1284 \pm 0.0003$	$3.0 \pm 0.3$	$16.6 \pm 0.2$	-22.4	$-1.6 \pm 0.1$
PMN J1539–1128	80	$17.8 \pm 0.1$	–	–	–	–	$-2.1 \pm 0.3$
WISE J154824.39+145702.8	42	$18.2 \pm 0.1$	$0.2308 \pm 0.0002$	$\leq 0.03$	$18.2 \pm 0.1$	-22.3	–
NVSS J163750–344915	80	$17.5 \pm 0.1$	–	–	–	–	$-2.3 \pm 0.2$
GB6 J1838+4802	250	$15.8 \pm 0.1$	–	–	–	–	$-1.1 \pm 0.1$
MITG J184126+2910	100	$17.4 \pm 0.1$	$0.2883 \pm 0.0003$	$2.6 \pm 0.5$	$18.2 \pm 0.3$	-22.9	$-1.5 \pm 0.1$
1RXS J184230.6–584202	35	$18.3 \pm 0.2$	0.421 <sup>(*)</sup>	$1.7 \pm 0.3$	$19.5 \pm 0.2$	-23.0	$-2.4 \pm 0.3$
1RXS J184230.6–584202	35	$18.3 \pm 0.2$	–	–	–	–	$-1.6 \pm 0.3$
1RXS J195815.6–301119	110	$15.8 \pm 0.1$	$0.1190 \pm 0.0003$	$1.2 \pm 0.2$	$16.7 \pm 0.2$	-22.1	$-1.9 \pm 0.2$
MAGIC J2001+435	105	$15.9 \pm 0.1$	$0.1739 \pm 0.0004$	$5.0 \pm 0.4$	$17.0 \pm 0.2$	-22.4	$-1.0 \pm 0.1$
1RXS J203650.9–332817	19	$17.8 \pm 0.1$	–	–	–	–	$-1.6 \pm 0.2$
RBS 1751	40	$16.9 \pm 0.1$	$\geq 0.618$ <sup>(*)</sup>	–	–	–	$-1.6 \pm 0.2$
1ES 2322–409	210	$16.0 \pm 0.2$	$0.1736 \pm 0.0008$	$9.5 \pm 0.6$	$17.7 \pm 0.2$	-22.1	$-1.6 \pm 0.4$

**Notes.** As the redshift of 1RXS J184230.6–584202 is based on a low-confidence detection of the CaHK feature, for this source we also present the results of a simple power-law fit. The spectral bin width is 4 Å for the sources observed with EFOSC2; 1 Å for MITG J184126+2910, MAGIC J2001+435, 1RXS J015658.6–530208, and 1RXS J020922.2–522920; and 2 Å for WISE J154824.39+145702.8. The columns are (1) Source name; (2) Median signal-to-noise ratio per spectral bin measured in continuum regions; (3)  $R_c$ , Cousins magnitude of the BL Lac spectrum corrected for reddening, telluric absorption, and slit losses with errors. Slit losses were estimated using an effective radius  $r_e = 10$  kpc for all sources except for MAGIC J2001+435 for which we used  $r_e = 6.8$  kpc (Aleksić et al. 2014); (4) Redshift or lower limit with error; (5) Flux ratio jet/galaxy at 5500 Å in rest frame; (6)  $R_c$ , Cousins magnitude of the galaxy with the same corrections as in Col. (3); (7) Absolute  $R$  magnitude of the galaxy, the errors are the same as those in Col. (6); (8) Power-law slope with errors. If the entry is unknown, the legend is ‘–’. <sup>(†)</sup> $U$  Magnitude. <sup>(\*)</sup>Uncertain redshift.

## 6.2. 1RXS J015658.6–530208

The BL Lac nature of 1RXS J015658.6–530208 was established by a low  $S/N$  featureless optical spectrum taken with the Goodman spectrograph at the Southern Astrophysical Research (SOAR) telescope (Landoni et al. 2015). It was later classified as an EBL (Costamante & Ghisellini 2002; Foffano et al. 2019).

Our SALT/RSS observations were performed on November 24 and 26, 2019. The transparency was good in both observations, and seeing was around 1.2 and 1.4 arcsec, respectively. The source was clearly detected in both observations at a median continuum  $S/N$  of 100. Inspection of the average spectrum shows the presence of clear CaHK and CaIG features at  $z \sim 0.304$  (Fig. 1, first row, right). This result is confirmed by the presence of a weaker MgB feature at the same redshift, while the CaFe feature falls into residual telluric absorption and is therefore undetectable. The fit of the features gives a precise redshift value  $z = 0.3043 \pm 0.0004$ .

We then compared the spectrum with the near-contemporaneous *Swift*/UVOT photometric points (see Sect. 4.2). The *Swift*/UVOT fluxes are higher than the SALT/RSS values, while the slopes are comparable (see Fig. B.1). We used the  $\nu$  *Swift*/UVOT filter, whose bandwidth is completely contained in our spectrum, to rescale it. The ratio of the  $\nu$  flux to the spectral flux in the corresponding range is  $1.3 \pm 0.2$ . We therefore multiplied the spectrum by 1.3 to match the UVOT photometry. A possible caveat to this analysis is the possibility of substantial variability (i.e. greater than our errors) between

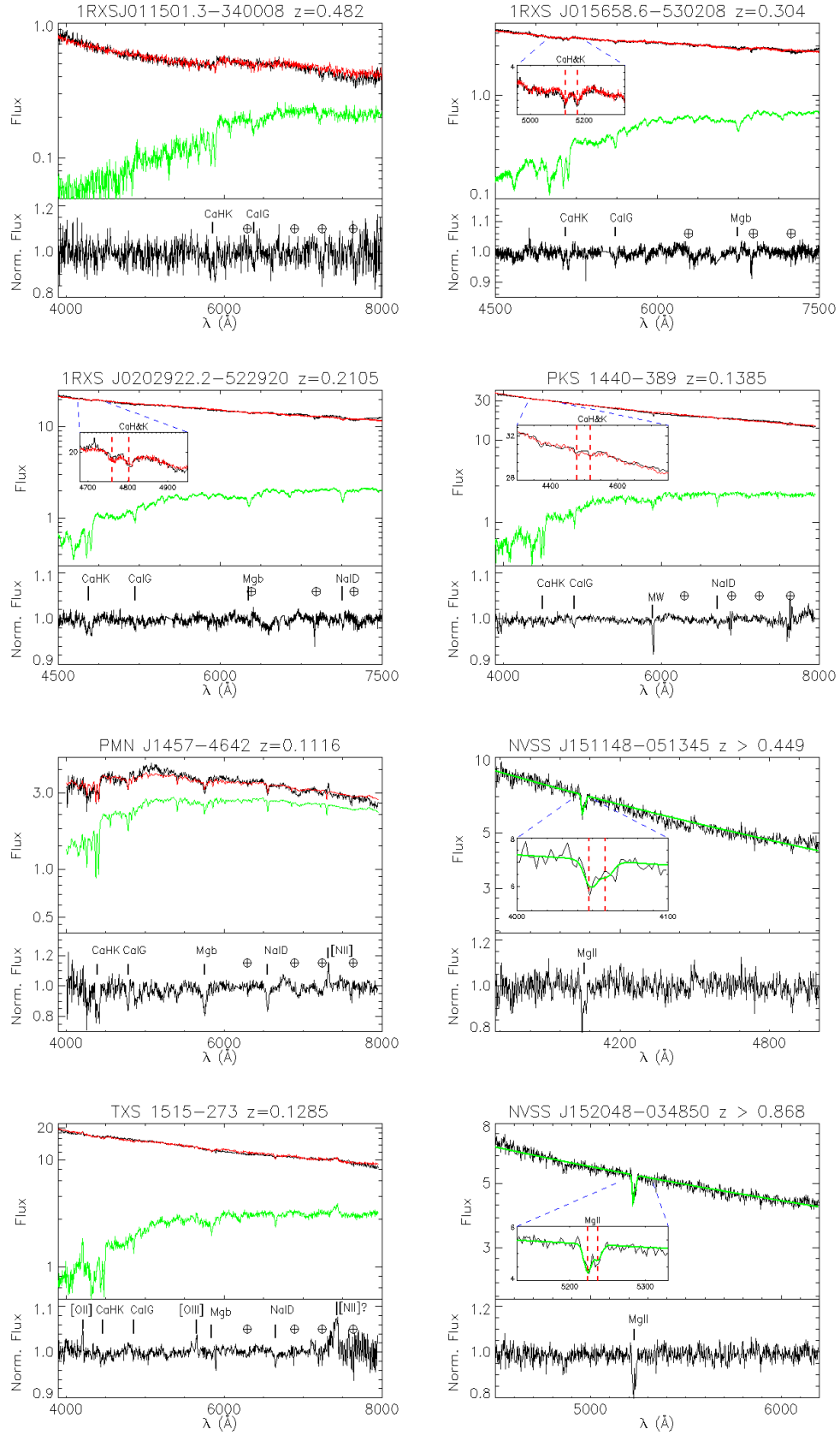
the spectroscopic and photometric observation. To estimate this effect we fitted the weekly variability distribution of the Catalina survey (Drake et al. 2009) photometry of the source with a Gaussian function. The  $2\sigma$  width of the distribution is 0.24 mag, which, summed in quadrature with the error of the flux ratio, pushes our errors to  $\pm 0.3$  which we use as the errors on the flux. The fit with a power law plus galaxy template model results in a bright host galaxy with  $M_R = -22.7 \pm 0.5$ .

## 6.3. 1RXS J020922.2–522920

Spectra of 1RXS J020922.2–522920 have been published in the six-degree Field Galaxy Survey (6dF) (Jones et al. 2004, 2009) and by Shaw et al. (2013). From the first spectrum a redshift  $z = 0.873$  has been derived, while from the second spectrum a statistical lower limit  $z \geq 0.31$  was proposed. We inspected both spectra and we could not detect any features in them; they also both have a low  $S/N$ .

We observed it with SALT/RSS on 2019 December 28. The source was clearly detected at a median continuum  $S/N$  of 160. Careful inspection of the spectrum revealed the presence of a clear CaHK feature at  $z \sim 0.211$  (Fig. 1, second row, left). This result is confirmed by the presence of weaker CaIG, MgB, and NaID features at the same redshift, while the CaFe feature falls into residual telluric absorption and is therefore undetectable. Our spectroscopic redshift value is  $z = 0.2110 \pm 0.0002$  and is at odds with both of the previously published values; however,





**Fig. 1.** Flux-calibrated and normalized spectra of the first eight sources in Table 1. Each panel contains the spectrum, continuum, and galaxy model for a given source. Each panel has two parts. *Upper:* flux-calibrated and telluric-corrected spectrum (black) alongside the best fit model (red). The flux is in units of  $10^{-16} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ \AA}^{-1}$ . The elliptical galaxy component is shown in green. *Lower:* normalised spectrum with labels for the detected absorption features. Atmospheric telluric absorption features are indicated by the symbol  $\oplus$  and Galactic absorption features are labelled ‘MW’.

assuming the same spectral state, the features we report here could not have been detected in the previous spectra.

We then compared the spectrum with the average ROSS photometric points (see Sect. 4.2). The average ROSS fluxes are all consistent with the SALT/RSS spectrum within 0.1 mag (see Fig. B.1). Concerning the effect of optical variability, we performed the same analysis as for 1RXSJ015658.6–530208 using the Catalina survey light curve of the source. In this case the expected variability in one week (the time between the SALT/RSS observation and the average time of the ROSS observations) is 0.14 mag ( $2\sigma$  level), again slightly increasing our errors to  $\pm 0.2$ . We therefore kept our original values. The host galaxy magnitude obtained from our fit is very bright at  $M_R = -23.2 \pm 0.4$ .

#### 6.4. PKS 1440–389

PKS 1440–389, in addition to its strong GeV gamma-ray emission detected by *Fermi*-LAT, is also a bright TeV source (Abdalla et al. 2020). A tentative redshift of  $z = 0.069$  has been reported from 6dF low S/N spectroscopy (Jones et al. 2004); we examined the spectrum, but could not find convincing spectral features at that redshift. Later, spectra with higher S/N (up to  $\sim 80$ ) taken by Shaw et al. (2013) and Landoni et al. (2015) were featureless and could not confirm this result. The magnitudes of the source at the times these spectra were taken are unknown. Very recently, combining optical (Shaw et al. 2013) and gamma-ray results, the redshift has been constrained to be in the range  $0.14 \leq z \leq 0.53$  at the 95% confidence limit (Abdalla et al. 2020). We were able to obtain a very high S/N EFOSC2 spectrum ( $S/N \geq 200$ , see Table 5) of PKS 1440–389 (see Fig. 1, second row, right). Careful inspection of the spectrum reveals the presence of CaHK, CaIG, and NaID features at  $z = 0.1385 \pm 0.0005$ . We note that at that redshift the Mgb feature of the galaxy falls into the strong Galactic NaID absorption, and is thus undetectable. This result is consistent with the range obtained by Abdalla et al. (2020) within slightly more than  $2\sigma$ , but it contradicts the tentative redshift published by 6dF galaxy survey.

The non-detection of these features in previous spectra is consistent with their low S/N assuming a similar optical spectral state. The host galaxy magnitude is average:  $M_R = -22.4 \pm 0.2$ .

#### 6.5. PMNJ1457–4642

PMNJ1457–4642 is one of the weakest and least studied of the gamma-ray sources in our sample, and no spectroscopic observation has been reported. Its EFOSC2 spectrum (see Fig. 1, third row, left) is clearly dominated by the emission of the host galaxy at redshift  $z = 0.1116 \pm 0.0002$ , with a measurable component from a non-thermal power law and a weak ( $EW \sim 5 \text{ \AA}$ )  $H\alpha$ -[NII] emission complex. The rest-frame power law-to-galaxy ratio at  $5500 \text{ \AA}$  is  $0.3 \pm 0.2$ . We note that the galaxy spectral shape cannot be well fitted with the local spectroscopic template of Mannucci et al. (2001). A better though not completely satisfying fit can be achieved with a Bruzual & Charlot (2003) template built with a 11 Gyr simple stellar population.

#### 6.6. NVSSJ151148–051345

Álvarez Crespo et al. (2016a) first reported a low S/N featureless optical spectrum for the HBL (Takeuchi et al. 2013) NVSSJ151148–051345. Recently, in an  $S/N \sim 200$ , resolution  $\lambda/\Delta\lambda \sim 250$ , GTC spectrum, a single feature around  $4053 \text{ \AA}$  with  $EW = 2.1 \text{ \AA}$  was detected (Paiano et al. 2017b). Interpreting this

feature as an unresolved MgII doublet implies that the redshift of the source is greater than 0.45. In order to resolve this feature we observed the source with EFOSC2 using Grism 14, which allows for  $\lambda/\Delta\lambda \sim 600$  in the  $3500\text{--}5000 \text{ \AA}$  wavelength range.

We obtained a  $S/N = 23$  spectrum in which we detect a double-peaked spectral feature at the position discussed above with total  $EW = 2.6 \pm 0.3 \text{ \AA}$  (see Fig. 1, third row, right). We fit the feature with an MgII doublet using vpfite (Carswell & Webb 2014) obtaining a  $\chi^2_{\text{d.o.f.}} \sim 1.07$  for a redshift  $z = 0.4480 \pm 0.0003$ . We consider it a firm result, therefore, that the redshift of NVSSJ151148–051345 is greater than 0.448. Finally the EWs of the two MgII components are about 1.7 and  $0.9 \text{ \AA}$ , their ratio is about 1.89, compatible with optically thin gas (see e.g. Raghunathan et al. 2016).

#### 6.7. TXS 1515–273

TXS 1515–273 has recently been recognised as a BL Lac (Lefaucheur & Pita 2017). An upper limit to its redshift,  $z \leq 1.1$  has been established from the lack of detection of the Ly  $\alpha$  break in UVOT and SARA photometry (Kaur et al. 2018). Moreover, a low S/N spectrum taken by Peña-Herazo et al. (2017) with the Goodman Spectrograph at the SOAR telescope did not show any features, leaving its redshift unknown.

In February 2019 it was detected in VHE gamma rays with the MAGIC telescope, triggered by a high state in HE gamma rays (Mirzoyan 2019). Observations from the Tuorla blazar monitoring programme (Nilsson et al. 2018) show that during that time TXS 1515–273 was in a high state also in the optical band (around  $R_c = 15.4$ ), but by the time of our spectroscopic observations (June 2019) the flux had decreased significantly (to  $R_c = 16.1$ )<sup>13</sup>. The Tuorla blazar monitoring magnitudes were derived with standard differential photometry and the comparison and control stars were calibrated using two nights of good weather data. We obtained a high S/N spectrum of TXS 1515–273 (see Fig. 1, fourth row, left) in which absorption and emission lines at  $z \sim 0.1285$  are visible along with a strong continuum. In absorption we detect with good confidence CaHK, CaIG, Mgb, and NaID. In emission [OII] and [OIII]  $\lambda 5007$  are detected with good confidence, while the [NII]/ $H\alpha$  complex is possibly present around  $7410 \text{ \AA}$ . The spectrum is very noisy at those wavelengths, however, due to instrumental effects; therefore, we do not analyse this possible feature in detail. The fit of the detected features gives a precise redshift value  $z = 0.1284 \pm 0.0003$ .

The magnitude of the spectrum after dereddening, telluric, and slit loss correction is  $R_c = 15.6 \pm 0.1$  (see Table 5); the observed magnitude without corrections is  $R_c = 16.2 \pm 0.2$ , fully compatible with the results of the Tuorla blazar monitoring. The host galaxy magnitude is  $M_R = -22.4 \pm 0.2$ . The rest-frame power law-to-galaxy ratio at  $5500 \text{ \AA}$  is  $3.0 \pm 0.3$ . These results are broadly consistent with those presented by Becerra González et al. (2021).

#### 6.8. NVSSJ152048–034850

NVSSJ152048–034850 is a BL Lac object, possibly an LBL or an IBL (Takeuchi et al. 2013). Shaw et al. (2013) report a low S/N detection of an MgII system at  $z = 0.867$  in a Palomar spectrum. Additionally, Kaur et al. (2017) report a redshift estimation of  $z = 1.46 \pm 0.1$  from a photometric detection of the Ly  $\alpha$  break using *Swift*/UVOT and Gamma-Ray burst Optical and Near-infrared (GROND) photometry. In order to

<sup>13</sup> [http://users.utu.fi/kani/1m/TXS\\_1515-273.html](http://users.utu.fi/kani/1m/TXS_1515-273.html)

assess this lower limit, we observed this source with EFOSC2 employing Grism 8 with wavelength range 4500–6200 Å and resolution 7.4 Å achieving  $S/N = 43$ . We detect a clearly asymmetric spectral feature at the position reported by Shaw et al. (2013) with total EW  $3.65 \pm 0.7$  Å. Single component fits yielded unacceptable chi-square values and line widths up to three times the spectral resolution. We then fit the feature with an MgII doublet using vpf (Carswell & Webb 2014) obtaining a  $\chi^2_{\text{d.o.f.}} \sim 1.0$  for a redshift  $z = 0.8680 \pm 0.002$  (see Fig. 1, fourth row, right). Therefore, we consider it a firm result that the redshift of NVSS J152048–034850 is greater than 0.8680. The EWs of the two components of the fit are roughly 2.3 and 1.2 Å. Their ratio is about 1.9, suggesting an optically thin gas (see e.g. Raghunathan et al. 2016). Finally, we note that NVSS J152048–034850 is a high-redshift BL Lac object with a possible redshift of  $z \sim 1.46$ , as suggested by Kaur et al. (2017), which makes it a very interesting target for future observations with CTA<sup>14</sup>.

### 6.9. PMN J1539–1128

A spectrum of PMN J1539–1128 was reported by Peña-Herazo et al. (2017). Using the Goodman spectrograph at the SOAR telescope, they obtained a featureless spectrum with  $S/N$  around 100. We observed the source with EFOSC2 using Grism 6 for 5400 s, obtaining a spectrum with median  $S/N = 80$  (see Fig. 2, first row, left). With the exception of the Galactic NaID, no spectral feature was detected and the redshift of PMN J1539–1128 remains unknown.

### 6.10. WISE J154824.39+145702.8

The optical counterpart of WISE J154824.39+145702.8 is classified as a galaxy in the Sloan Digital Sky Survey (SDSS) (Ahumada et al. 2020) with photometric redshift  $z = 0.217 \pm 0.02$ . Álvarez Crespo et al. (2016b) performed a 40 min observation with the Device Optimized for the LOw RESolution (DOLORES) spectrograph at the Telescopio Nazionale Galileo (TNG). The resulting 3700–8100 Å spectrum is dominated by the host galaxy emission at a redshift of  $z = 0.231$ . We took a one-hour spectrum of WISE J154824.39+145702.8 with ESI at Keck. We confirm the result of Álvarez Crespo et al. (2016b) that the optical emission of the source is dominated by the host galaxy at  $z \sim 0.231$  (see Fig. 2, first row, right). The rest-frame jet-to-galaxy ratio is lower than 0.07 ( $3\sigma$  limit). The best fit is obtained using the template by Mannucci et al. (2001). Thanks to our higher sensitivity and wider wavelength range, we detect several additional emission lines: [OII], [OIII]  $\lambda 4959$ , [OIII]  $\lambda 5007$ , [OI]  $\lambda 6300$ , H $\alpha$ , [NII]  $\lambda 6548$ , [NII]  $\lambda 6583$ , and the [SII] doublet (see Table 4). The fit of the detected features gives a precise redshift value  $z = 0.2308 \pm 0.0002$ . Some lines have an EW greater than 5 Å, the standard BL Lac limit, showing that this host galaxy is peculiarly active and gas-rich for a BL Lac object. In order to investigate this issue, we measured their EWs by fitting Gaussian functions; the results are listed in Table 4. Computing the ratios [OIII]/[OII] and [OI]/H $\alpha$  and using Fig. 5 in Kewley et al. (2006), the host galaxy can be classified as a LINER.

The SED shown in Fujinaga et al. (2016) shows the presence of a peak in the NIR–optical domain. The above results show that

<sup>14</sup> Photons at energy levels up to 100 GeV at  $z = 1.5$  and higher have been detected by *Fermi*-LAT; see Sect. 3.6 and Fig. 17 in Ajello et al. (2017). CTA is expected to have a better sensitivity than *Fermi*-LAT at those energies.

it is due to the galaxy emission. The very low jet-to-host ratio suggests an HBL or possibly EHBL nature of the source, but the lack of multiwavelength data (in particular X-ray data) makes a firm conclusion impossible at this moment. More observations are needed to settle this question.

### 6.11. NVSS J163750–344915

Peña-Herazo et al. (2017) reported on a spectrum of NVSS J163750–344915 obtained with the Goodman Spectrograph at SOAR (4000–7000 Å). The spectrum has  $S/N = 60$  and is featureless, confirming its BL Lac classification. We took an EFOSC2 spectrum of the object aiming to obtain a better  $S/N$ . However, we could only integrate for 45 min due to bad weather. The resulting spectrum has a  $S/N = 80$  and is featureless (see Fig. 2, second row, left). The redshift of NVSS J163750–344915 remains undetermined.

### 6.12. GB6 J1838+4802

No spectroscopic redshift measure is available for GB6 J1838+4802, but its redshift has been photometrically estimated at  $z \sim 0.3$  from the detection of the host galaxy at magnitude  $R_c = 19.12 \pm 0.05$  in imaging observations (Nilsson et al. 2003). Even if the redshift estimates depend on assumed host galaxy properties, and are therefore by definition less accurate than spectroscopic measurements, the detection of the host galaxy supports the possibility of also measuring the redshift spectroscopically. However, our observation of GB6 J1838+4802 resulted in a featureless spectrum (Fig. 2, second row, right), despite its very high  $S/N$  ( $\sim 250$ ). We note that this non-detection may be due to the brightness of the source; at  $R_c \sim 15.8$  the nuclear emission is about 3.3 mag (about 20 times) stronger than the host emission. In a case like this one the best option may be to trigger another spectroscopic observation during a minimum in the optical light curve.

### 6.13. MITG J184126+2910

The optical counterpart of MITG J184126+2910 was identified by Marchesini et al. (2016). The moderate  $S/N$ , featureless, optical spectrum they took with the DOLORES spectrograph at the TNG confirmed its classification as a BL Lac object. A later spectrum (Marchesi et al. 2018) taken with the Kitt Peak Ohio-State Multi-Object Spectrograph (KOSMOS) spectrograph at the Mayall Telescope is also featureless.

We note that the infrared counterpart, 2MASX J18412170 + 2909404, is extended, suggesting that the emission of the host galaxy may be detectable. Our Keck/ESI spectrum (see Fig. 2, third row, left) reaches an average  $S/N$  of 100. In the mostly featureless spectrum we identified the CaHK doublet and CaIG features at  $5\sigma$  and  $4\sigma$ , respectively, at redshift  $z \sim 0.288$ . We also possibly identify the Mgb and CaFe features at about  $3$ – $4\sigma$  at the same redshift. A precise redshift  $z = 0.2883 \pm 0.0003$  is obtained fitting these features. The absolute magnitude of the host galaxy is  $M_R = -22.9 \pm 0.3$ , bright but within the expected range.

### 6.14. 1RXS J184230.6–584202

Two spectra of 1RXS J184230.6–584202 have been reported recently. One spectrum (Desai et al. 2019) was taken on June 1, 2018, with KOSMOS at the 4 m Mayall telescope. The source had magnitude  $V = 17.5$ , and the spectrum was featureless with

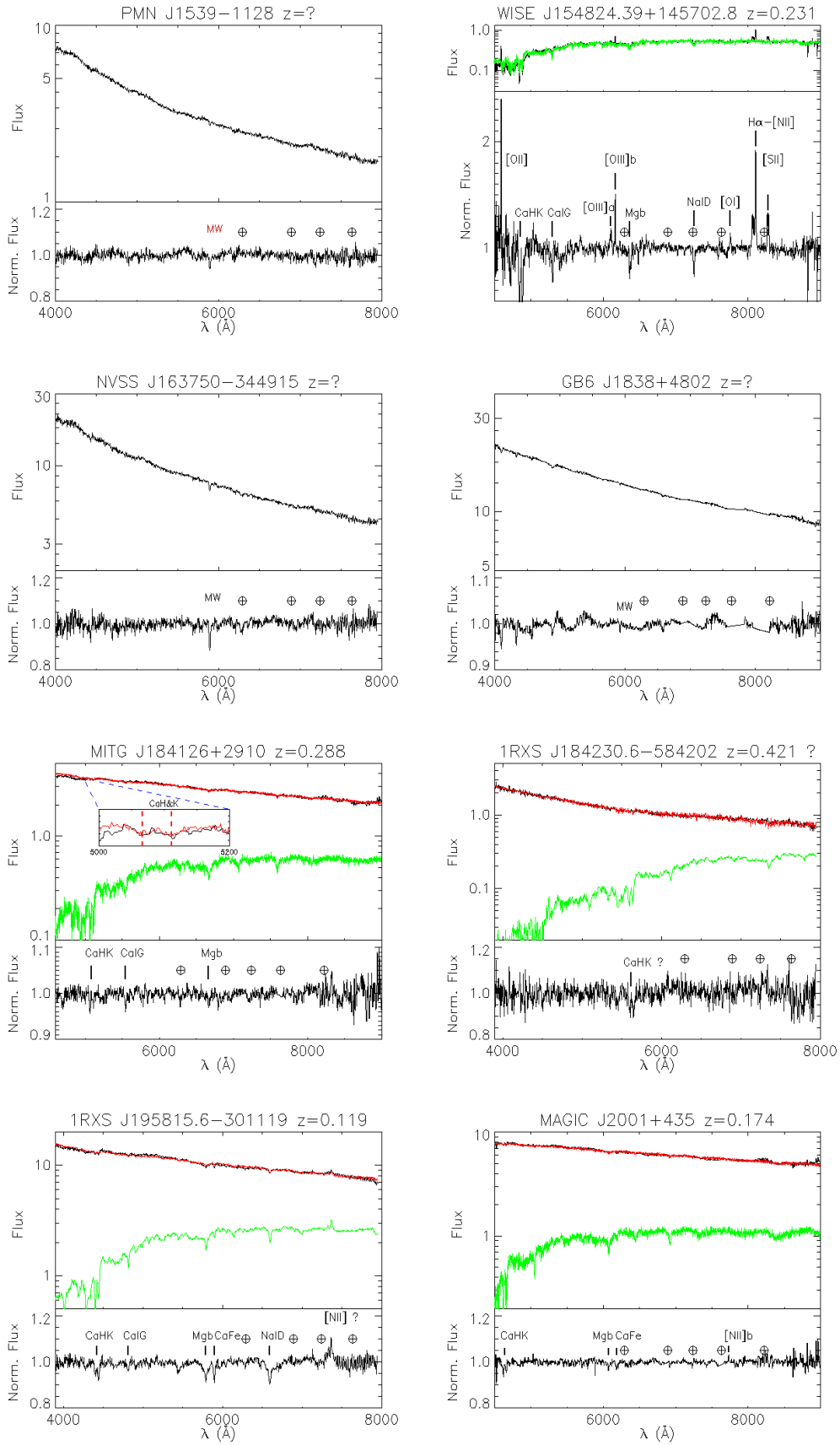


Fig. 2. Same as Fig. 1 for sources 9 to 16 in Table 1.

a S/N of 30. The other was taken with the Goodman spectrograph on the SOAR telescope (Marchesini et al. 2019) on May 9, 2017. The S/N is similar to the first one, but in this case the authors report the detection of the CaHK and CaIG features at redshift  $z = 0.421$ . The magnitude of the source is not reported by these authors.

In order to verify the presence of these features, we observed 1RXS J184230.6–584202 for 1 h and 15 min obtaining a S/N of 35. The spectrum has a power-law shape with magnitudes  $V = 18.8 \pm 0.2$  and  $R_c = 18.3 \pm 0.2$  (see Fig. 2, third row, right). After careful examination we possibly detect at the  $3.6\sigma$  level the CaHK feature at  $z \sim 0.421$  (the same redshift as in Marchesini et al. 2019). We note that at magnitude  $V = 18.8$  the source was much fainter than in Desai et al. (2019), which may explain our weak detection and their non-detection. Considering the low significance of the detection, we consider this redshift value as tentative.

### 6.15. 1RXS J195815.6–301119

Two independent publications give the same value for its redshift,  $z = 0.119$ . This value is reported in the DR3 6dF catalogue (Jones et al. 2004, 2009) and in the Reflex Cluster survey (Guzzo et al. 2009). The public 6dF spectrum has a very low S/N (see Fig. A.1), while the Reflex spectrum is not available.

We thus took a 30 min spectrum with EFOSC2 in order to confirm or disprove this result. The resulting high S/N spectrum is a combination of a power law with the emission of an elliptical galaxy modelled with an 11 Gyr template (Bruzual & Charlot 2003) at  $z = 0.1190 \pm 0.0003$ . Weak [NII] emission is also detected (see Fig. 2, fourth row, left). The rest-frame power law-to-galaxy ratio at  $5500 \text{ \AA}$  is  $1.2 \pm 0.2$  and the galaxy has absolute magnitude  $M_R = -22.1 \pm 0.2$ . This is thus a firm redshift value.

### 6.16. MAGIC J2001+435

MAGIC J2001+435 is a BL Lac object (Bassani et al. 2009) detected at TeV energies (Aleksić et al. 2014). The host galaxy has been detected in deep imaging secured during low blazar activity. The host magnitude was measured to be  $I = 17.15 \pm 0.06$ , which was translated into a photometric redshift  $z = 0.18 \pm 0.04$  (Aleksić et al. 2014). A high S/N ( $\sim 150$ – $200$ ) optical Keck/Low Resolution Imaging Spectrometer (LRIS) spectrum was taken by Shaw et al. (2013) on August 18, 2009, but no spectral features could be detected. From the spectrum (R. Romani, priv. comm.) the source had an observed magnitude  $R_c = 14.8$ , more than 15 times stronger than the host.

In May 2018 the Tuorla blazar monitoring light curve<sup>15</sup> showed that MAGIC J2001+435 was in a low state with an observed magnitude more than two magnitudes fainter than at the epoch of the Keck/LRIS spectrum. It was decided, therefore, to observe it with Keck/ESI in place of a previously scheduled target.

The exposure time was 3480 s and the resulting S/N was 105. In the final spectrum we estimate a redshift  $z = 0.1739 \pm 0.0002$  from the detection of the CaHK, Mgb, CaFe, and possibly NaID absorption features (Fig. 2, fourth row, right). At a  $4.6\sigma$  level we also detect the [NII]  $\lambda 6583$  emission line at the same redshift. The result is consistent with the imaging estimate in Aleksić et al. (2014) at less than  $1\sigma$ . Our redshift measurement was possible even though the S/N of our spectrum is much lower

than that from Shaw et al. (2013), demonstrating that waiting for an optical low state may be very effective for blazar redshift measurement.

### 6.17. 1RXS J203650.9–332817

1RXS J203650.9–332817 was observed by Álvarez Crespo et al. (2016b) who took a 20 min spectrum with the Goodman spectrograph on the SOAR telescope. The spectrum has a S/N of about 50 and a general power-law shape consistent with classification as a BL Lac object. A weak feature around  $5000 \text{ \AA}$ , if interpreted as CaHK, sets a redshift value of  $z = 0.237$ .

Due to the weakness of the feature we observed the source with EFOSC2 in order to confirm or disprove this result with a higher S/N spectrum. However, due to bad weather, we were only able to obtain a low S/N featureless spectrum (Fig. 3, first row, left). The redshift of 1RXS J203650.9–332817 remains undetermined.

### 6.18. RBS 1751

RBS 1751 was classified as a BL Lac object on the basis of a noisy spectrum from the 6dF survey (Jones et al. 2004, 2009). We observed the source with EFOSC2 for one hour in difficult atmospheric conditions. The resulting spectrum has a S/N of 40, higher than the 6dF value, and has a power-law shape (Fig. 3, first row, right). Only after normalisation can a weak absorption feature with  $EW = 2.3 \pm 0.5 \text{ \AA}$  centred around  $\sim 4530 \text{ \AA}$  be found. If we interpret this feature as being due to an unresolved MgII system, then the source would be located at  $z \geq 0.618$ . Unfortunately, the grism we used, Grism 6, does not have enough spectral resolution to separate the two components of a possible MgII absorber. Interestingly, the only possible feature present in the 6dF spectrum is at the same wavelength; however, it is very weak. A deeper observation with better spectral resolution and S/N is needed to solve this issue. We conclude that RBS 1751 is tentatively at a redshift  $z \geq 0.618$ . We provide no error estimate on this value as the possible feature could not be fitted.

### 6.19. 1ES 2322–409

1ES 2322–409 is a bright TeV emitter, serendipitously discovered by H.E.S.S. (see Abdalla et al. 2019). The redshift of this source is uncertain; a value  $z = 0.174$  has been suggested on the basis of a low S/N 6dF spectrum (Jones et al. 2004, 2009). We note that the redshift measurement was mostly based on a weak feature at  $\sim 6900 \text{ \AA}$ , in a region where there is strong telluric absorption. A higher S/N spectrum taken with the Focal Reducer and low dispersion Spectrograph (FORS) at the Very Large Telescope was featureless (Landoni et al. 2013). The spectrum is available on the ZBLAC<sup>16</sup> database and is quite noisy. When the FORS spectrum was taken the source was at magnitude  $R = 15.7$  (Landoni et al. 2013); there is no information on the source magnitude during the 6dF observation.

In order to clarify this issue we took a high S/N EFOSC2 spectrum (Fig. 3, second row). The source was at magnitude  $R_c = 16.1$  and the spectrum is dominated by a non-thermal continuum. In the continuum we were able to detect several weak features: CaHK, CaIG, Mgb, and NaID at a common redshift  $z = 0.1736 \pm 0.0008$ . We consider this a firm redshift. We thus

<sup>15</sup> [http://users.utu.fi/kani/1m/MG4\\_J200112+4352.html](http://users.utu.fi/kani/1m/MG4_J200112+4352.html)

<sup>16</sup> <http://web.oapd.inaf.it/zblac/>

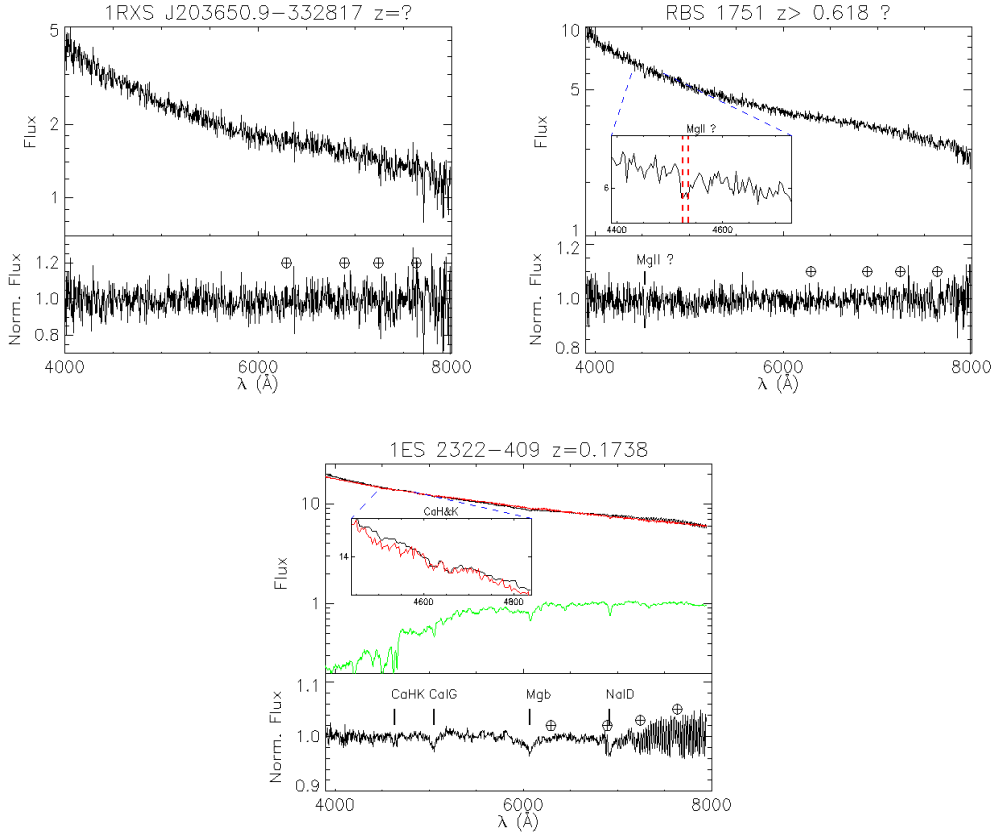


Fig. 3. Same as Fig. 1 for the last three sources in Table 1.

confirm the 6dF result, despite the low S/N of their spectrum. The detection of the NaID feature may have been possible due to the source being in a low state.

## 7. Discussion and conclusions

We observed 19 BL Lacs detected at  $E \geq 10$  GeV with the *Fermi*-LAT satellite and/or by ground-based Cherenkov telescopes. The observations were performed with the ESI spectrograph at the Keck Observatory, with the RSS spectrograph at the South African Astronomical Observatory and with the EFOSC2 spectrograph at the ESO/La Silla observatory. The observing strategy called for obtaining, whenever possible, S/N values per pixel greater than 100. We aimed to measure or constrain the redshift and measure the properties of the host galaxy of each of our target BL Lacs. In the following subsections we discuss these results in more detail and conclude with a summary.

### 7.1. Spectral signal-to-noise ratio

Only nine of our targets reached the target S/N: 1RXS J015658.6-530208, 1RXS J020922.2-522920, PKS 1440-389, TXS 1515-273, GB6 J1838+4802, MITG J184126+2910, 1RXS J195815.6-301119, MAGIC J2001+435, and 1ES 2322-409. We measured the redshift for eight of them. The other ten have S/N values between 20 and 80, and from them we could measure only three firm redshifts and one tentative value. In addition we measured two firm lower limits and one tentative lower limit.

A high S/N spectrum thus appears to be necessary to allow the measurement of the redshift for the gamma-ray BL Lacs in

our sample. On the other hand, a low S/N spectrum, while useful for classification, is not likely to succeed in this effort. The comparative success of high S/N spectra with respect to those with low S/N supports our strategy. However, we also note that in the case of GB6 J1838 + 4802, even a very high S/N spectrum did not lead to a successful measurement of the redshift. This is in line with several previous works (Paiano et al. 2017a; Pita et al. 2014) where it was not possible to measure redshifts of well-known blazars, even with very high-quality spectra. Therefore, while for gamma-ray BL Lacs reaching a S/N of at least 100 is a good indicator of the success of a redshift measurement, it is not a guarantee. Additionally, both of our SALT/RSS spectra and three of our four Keck/ESI spectra reached our target S/N, while only 4 out of 13 EFOSC2 spectra reached that goal. While the sample is still small, this suggests that 10 m class telescopes are the best tools for our programme. Finally, technical or weather issues prevented us from reaching the target S/N in ten of our spectra; three of them have no redshift measurement or lower limit. We will re-observe these three sources to secure a spectrum with a S/N of 100. For any sources for which the redshift will not have been measured even then, we plan to organise dedicated target of opportunity programmes to obtain high S/N spectroscopy during their optical low states.

### 7.2. Optical extension of the sources

We had eight targets with extended counterparts: four of them have extended NIR counterparts in the 2MASS Extended Source Catalogue (2MASX) (Jarrett et al. 2000). Of the remaining four, the extension has been detected for three of them in dedicated observations and the last one, WISE J154824.39+145702.8, is classified as extended in the SDSS database. The 2MASX

sources are TXS 1515–273, PMN J1457–4642, MITG J184126+2910, and 1RXS J195815.6–301119. We measured the redshift for all of them with values between 0.1116 and 0.288 and an average value  $\langle z(2\text{MASX}) \rangle = 0.162$ . This may be an indication that gamma-ray BL Lacs whose counterparts belong to the 2MASX catalogue are good candidates for redshift measurement provided a sufficient S/N is obtained. The low average value of the redshift may be due to the relatively low sensitivity of the 2MASX survey that can reliably detect the extension only for sources with  $J \geq 15.0$ <sup>17</sup>, which is the case for the four sources mentioned above.

The remaining extended sources are GB6 J1838+4802 (Nilsson et al. 2003), WISE J154824.39+145702.8, PKS 1440–389 (Fallah Ramazani et al., in prep.), and MAGIC J2001+435 (Aleksić et al. 2014). The redshifts of the last three were measured. As discussed above, our observations of PKS 1440–389 are much deeper than previous observations, but no information on its optical state during previous observations is available, and we confirmed the previous redshift measurement of WISE J154824.39+145702.8. The situation is different for the other two sources; MAGIC J2001+435 was in an optical low state, which eased the redshift measurement, while a previous more sensitive observation failed to detect any features. Conversely GB6 J1838+4802 was in an optically high state, which likely hampered the redshift measurement. Overall, we measured seven redshifts for eight extended targets, which confirm previous results (Nilsson et al. 2003) of the  $\sim 90\%$  efficiency of redshift measurements for extended counterparts. Moreover, the cases of GB6 J1838+4802 and MAGIC J2001+435 show that the optical state of the BL Lac object can be relevant for our purpose, and that using the optical state of the BL Lac object as a further parameter of our strategy may improve the efficiency of our programme.

### 7.3. Properties of the host galaxies

The average magnitude of the 11 firmly detected host galaxies is  $M_R = -22.6 \pm 0.4$ , slightly brighter than but compatible with the value reported in Shaw et al. (2013), and fainter than the values in Sbarufatti et al. (2005) and Pita et al. (2014). The best fit template was the local one (Mannucci et al. 2001) in nine cases. In the remaining cases the best fit was produced by a simple stellar population template (Bruzual & Charlot 2003) with age 11 Gyr (PMN J1457–4642 and 1RXS J195815.6–301119) and 2.5 Gyr (1RXS J011501.3–340008). This is broadly in agreement with previous results that host galaxies of BL Lacs are normal elliptical galaxies (Urry et al. 2000). We detected faint and narrow emission lines in only five of our targets, in line with previous results at similar sensitivities (Pita et al. 2014). We did not estimate upper limits on the magnitudes of the non-detected host galaxies because the S/N of their spectra was not high enough to yield a meaningful limit.

### 7.4. Comparison with other campaigns

We cannot present an exhaustive literature review here, but we briefly compare the first results from our programme to those of the main previous and ongoing spectroscopic observations aimed at measuring the redshifts of gamma-ray blazars. While we present here observations of 19 sources, our campaign

targets 165 sources; we thus limit our comparison to campaigns targeting extensive samples with about 100 targets or more: the Shaw et al. (2013) campaign, the campaign described by Peña-Herazo et al. (2020), and the recently undertaken campaign by Paiano et al. (2020). With respect to the first one, our sample contains brighter sources with harder gamma-ray spectra, but it contains fewer targets, while our spectra have on average higher and more uniform S/N. Regarding the campaign by Peña-Herazo et al. (2020), we target brighter objects with harder gamma-ray spectra and with better determined classifications; we have fewer targets and we have much higher S/N spectra. With respect to both of these campaigns we can thus expect to have a higher redshift detection efficiency. Finally, the campaign started by Paiano et al. (2020) includes 91 bright objects from the 3FHL catalogue and the authors plan to take high S/N spectra for all of them from the GTC telescope in the Canary Islands. Their sample is smaller than ours and is limited to the Northern Hemisphere, and its gamma-ray selection is much simpler than ours. Their selection is based on simple extrapolations of the flux and spectral index provided by the catalogue, while ours relies on full Monte Carlo simulations of the spectra and EBL absorption using the CTA instrumental response. We think we will have a similar redshift detection efficiency on a larger sample extended to the Southern Hemisphere. With respect to all of these campaigns, we will also add imaging information which will deepen our knowledge of the targets, for example by allowing us to determine (or constrain) more precisely the magnitude of the host galaxy.

### 7.5. Summary

We report on the first results of our programme aimed at measuring the redshift of gamma-ray bright blazars likely to be detected with CTA. Our main results are the following:

1. We performed spectroscopic observations of 19 gamma-ray blazars. Seventeen of them were previously observed in spectroscopy and seven uncertain redshift values were proposed in the literature; we confirmed four and disproved two of them, while the remaining one is still dubious. We measured 11 firm redshifts and 1 tentative redshift with values ranging from  $z = 0.1116$  to  $z = 0.4824$ .
2. Six of the blazars are at redshifts of  $z > 0.2$ , where fewer than 15 VHE BL Lacs are currently known. In particular, we measured firm spectroscopic redshifts of three known TeV sources: PKS 1440–389, 1ES 2322–409, and MAGIC J2001+435. After our observations the number of TeV BL Lacs without redshift values decreases from 13 to 10.
3. We confirmed two previously suggested spectroscopic lower limits and detected one tentative lower limit with values ranging from  $z \geq 0.4480$  to  $z \geq 0.8680$ .
4. We achieved high efficiency (eight out of nine) in redshift measurement for high S/N spectra, and low efficiency (three-four out of ten) in redshift measurement for low S/N spectra.
5. We achieved high efficiency (seven out of eight) in redshift measurement for BL Lacs with extended optical counterparts.
6. We measured the redshift of MAGIC J2001+435 during an optical low state. This is an indication that spectroscopic observations triggered during a spectroscopic minimum can improve redshift measurement efficiency.
7. We measured the average magnitude of the host galaxies,  $M_R = -22.6 \pm 0.4$ , which is compatible with previous measurements of this quantity for gamma-ray blazars. The

<sup>17</sup> [https://old.ipac.caltech.edu/2mass/releases/allsky/doc/sec6\\_1h.html](https://old.ipac.caltech.edu/2mass/releases/allsky/doc/sec6_1h.html)

properties of the hosts are consistent with normal elliptical galaxies.

The overall redshift detection efficiency of 58% or (11/19) is higher than that in the most complete survey of gamma-ray blazars (Shaw et al. 2013) of 44%. This preliminary difference may be due to the different sample selection criteria and/or to the higher S/N of our spectra on average. Due to bad weather and, in some cases, limited sensitivity, not all our spectra matched our criteria.

Our programme of redshift observations will continue on the objects from our sample with the goal of obtaining a spectrum with S/N 100 for each object, in an optical low state if needed. This will help to shape the CTA Key Science Programme on AGN of blazar type. This programme will also serve the astronomical community at large. On the one hand, it will enable scientists studying blazar emission and populations to better tune their models. On the other hand, scientists applying for time not only with CTA, but also with present-day VHE observatories and with other multiwavelength facilities such as *Fermi*-LAT, *XMM-Newton*, *Chandra*, ALMA, and VLA will be able to select optimal targets for their observing programmes from a larger pool of blazars with known redshifts.

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## References

Abdalla, H., Aharonian, F., Ait Benkhali, F., et al. 2019, *MNRAS*, **482**, 3011  
 Abdalla, H., Adam, R., Aharonian, F., et al. 2020, *MNRAS*, **494**, 5590  
 Abdo, A. A., Ackermann, M., Agudo, I., et al. 2010, *ApJ*, **716**, 30  
 Abeyskara, A. U., Archer, A., Benbow, W., et al. 2019, *ApJ*, **885**, 150  
 Abramowski, A., Acero, F., Aharonian, F., et al. 2013, *Phys. Rev. D*, **88**, 102003  
 Acciari, V. A., Ansoldi, S., Antonelli, L. A., et al. 2019, *MNRAS*, **486**, 4233  
 Acero, F., Donato, D., Ojha, R., et al. 2013, *ApJ*, **779**, 133  
 Ackermann, M., Ajello, M., Baldini, L., et al. 2018, *ApJS*, **237**, 32  
 Aharonian, F. A., Coppi, P. S., & Voelck, H. J. 1994, *ApJ*, **423**, L5  
 Ahumada, R., Allende Prieto, C., Almeida, A., et al. 2020, *ApJS*, **249**, 3

Ajello, M., Atwood, W. B., Baldini, L., et al. 2017, *ApJS*, **232**, 18  
 Aleksić, J., Ansoldi, S., Antonelli, L. A., et al. 2014, *A&A*, **572**, A121  
 Álvarez Crespo, N., Masetti, N., Ricci, F., et al. 2016a, *AJ*, **151**, 32  
 Álvarez Crespo, N., Massaro, F., Milisavljevic, D., et al. 2016b, *AJ*, **151**, 95  
 Alves Batista, R., Saveliev, A., & de Gouveia Dal Pino, E. M. 2019, *MNRAS*, **489**, 3836  
 Angel, J. R. P., & Stockman, H. S. 1980, *ARA&A*, **18**, 321  
 Angelakis, E., Hovatta, T., Blinov, D., et al. 2016, *MNRAS*, **463**, 3365  
 Arsioli, B., & Chang, Y. L. 2017, *A&A*, **598**, A134  
 Arsioli, B., Fraga, B., Giommi, P., Padovani, P., & Marrese, P. M. 2015, *A&A*, **579**, A34  
 Ballet, J., Burnett, T. H., Digel, S., & Lott, B. 2020, ArXiv e-prints [arXiv:2005.11208]  
 Bassani, L., Landi, R., Masetti, N., et al. 2009, *MNRAS*, **397**, L55  
 Becerra González, J., Acosta-Pulido, J. A., Boschini, W., et al. 2021, *MNRAS*, **504**, 5258  
 Biteau, J., & Williams, D. A. 2015, *ApJ*, **812**, 60  
 Blanton, M. R., Bershad, M. A., Abolfathi, B., et al. 2017, *AJ*, **154**, 28  
 Breeveld, A. A., Curran, P. A., Hoversten, E. A., et al. 2010, *MNRAS*, **406**, 1687  
 Bruzual, G., & Charlot, S. 2003, *MNRAS*, **344**, 1000  
 Buckley, D. A. H., Andreoni, I., Barway, S., et al. 2018, *MNRAS*, **474**, L71  
 Burgh, E. B., Nordsieck, K. H., Kobulnicky, H. A., et al. 2003, in Prime Focus Imaging Spectrograph for the Southern African Large Telescope: Optical Design, eds. M. Iye, & A. F. M. Moorwood, *SPIE Conf. Ser.*, **4841**, 1463  
 Buzzoni, B., Delabre, B., Dekker, H., et al. 1984, *The Messenger*, **38**, 9  
 Cappellari, M. 2017, *MNRAS*, **466**, 798  
 Carswell, R. F., & Webb, J. K. 2014, Astrophysics Source Code Library [record ascl:1408.015]  
 Cerruti, M., Zech, A., Boisson, C., & Inoue, S. 2015, *MNRAS*, **448**, 910  
 Cherenkov Telescope Array Consortium 2019, *Science with the Cherenkov Telescope Array* (World Scientific Publishing Co)  
 Cherenkov Telescope Array Consortium 2021, *JCAP*, **02**, 048  
 Costamante, L., & Ghisellini, G. 2002, *A&A*, **384**, 56  
 Costamante, L., Ghisellini, G., Giommi, P., et al. 2001, *A&A*, **371**, 512  
 Covino, S., Stefanon, M., Sciuto, G., et al. 2004, in REM: A Fully Robotic Telescope for GRB Observations, eds. A. F. M. Moorwood, & M. Iye, *SPIE Conf. Ser.*, **5492**, 1613  
 Crawford, S. M., Still, M., Schellart, P., et al. 2010, *SPIE Conf. Ser.*, **7737**, 773725  
 de Angelis, A., Galanti, G., & Roncadelli, M. 2011, *Phys. Rev. D*, **84**, 105030  
 Deil, C., Zanin, R., Lefaucheur, J., et al. 2017, in *35th International Cosmic Ray Conference (ICRC2017)*, Int. Cosm. Ray Conf., **301**, 766  
 Desai, A., Marchesi, S., Rajagopal, M., & Ajello, M. 2019, *ApJS*, **241**, 5  
 Domínguez, A., & Ajello, M. 2015, *ApJ*, **813**, L34  
 Domínguez, A., Primack, J. R., Rosario, D. J., et al. 2011, *MNRAS*, **410**, 2556  
 Drake, A. J., Djorgovski, S. G., Mahabal, A., et al. 2009, *ApJ*, **696**, 870  
 Falomo, R. 1996, *MNRAS*, **283**, 241  
 Falomo, R., & Kotilainen, J. K. 1999, *A&A*, **352**, 85  
 Falomo, R., Pian, E., & Treves, A. 2014, *A&ARv*, **22**, 73  
 Fermi-LAT Collaboration (Abdollahi, S., et al.) 2018, *Science*, **362**, 1031  
 Fitzpatrick, E. L. 1999, *PASP*, **111**, 63  
 Foffano, L., Prandini, E., Franceschini, A., & Paiano, S. 2019, *MNRAS*, **486**, 1741  
 Fosbury, R. A. E., & Disney, M. J. 1976, *ApJ*, **207**, L75  
 Franckowiak, A., Garrappa, S., Paliya, V., et al. 2020, *ApJ*, **893**, 162  
 Fujinaga, Y., Niinuma, K., Kimura, A., et al. 2016, *PASJ*, **68**, 70  
 Furniss, A., Williams, D. A., Danforth, C., et al. 2013, *ApJ*, **768**, L31  
 Gehrels, N., Chincarini, G., Giommi, P., et al. 2004, *ApJ*, **611**, 1005  
 Ghisellini, G., Tavecchio, F., Foschini, L., & Ghirlanda, G. 2011, *MNRAS*, **414**, 2674  
 Ghisellini, G., Righi, C., Costamante, L., & Tavecchio, F. 2017, *MNRAS*, **469**, 255  
 Giommi, P., & Padovani, P. 2015, *MNRAS*, **450**, 2404  
 Giommi, P., Piranomonte, S., Perri, M., & Padovani, P. 2005, *A&A*, **434**, 385  
 Giommi, P., Padovani, P., Oikonomou, F., et al. 2020, *A&A*, **640**, L4  
 Giovannini, G., Falomo, R., Scarpa, R., Treves, A., & Urry, C. M. 2004, *ApJ*, **613**, 747  
 Goldoni, P., Pita, S., Boisson, C., et al. 2016, *A&A*, **586**, L2  
 Guzzo, L., Schuecker, P., Böhringer, H., et al. 2009, *A&A*, **499**, 357  
 H.E.S.S. Collaboration (Abdalla, H., et al.) 2017, *A&A*, **606**, A59  
 Halpern, J. P., Eracleous, M., & Forster, K. 1997, *AJ*, **114**, 1736  
 Hauser, M. G., & Dwek, E. 2001, *ARA&A*, **39**, 249  
 IceCube Collaboration (Aartsen, M. G., et al.) 2018a, *Science*, **361**, eaat1378  
 IceCube Collaboration (Aartsen, M. G., et al.) 2018b, *Science*, **361**, 147  
 Jarrett, T. H., Chester, T., Cutri, R., et al. 2000, *AJ*, **119**, 2498  
 Jones, D. H., Saunders, W., Colless, M., et al. 2004, *MNRAS*, **355**, 747  
 Jones, D. H., Read, M. A., Saunders, W., et al. 2009, *MNRAS*, **399**, 683  
 Jorstad, S. G., Marscher, A. P., Morozova, D. A., et al. 2017, *ApJ*, **846**, 98



- Kaur, A., Rau, A., Ajello, M., et al. 2017, *ApJ*, **834**, 41
- Kaur, A., Rau, A., Ajello, M., et al. 2018, *ApJ*, **859**, 80
- Kaur, A., Ajello, M., Marchesi, S., & Omodei, N. 2019, *ApJ*, **871**, 94
- Kausch, W., Noll, S., Smette, A., et al. 2015, *A&A*, **576**, A78
- Kewley, L. J., Groves, B., Kauffmann, G., & Heckman, T. 2006, *MNRAS*, **372**, 961
- Kifune, T. 1999, *ApJ*, **518**, L21
- Klindt, L., van Soelen, B., Meintjes, P. J., & Väisänen, P. 2017, *MNRAS*, **467**, 2537
- Kobulnicky, H. A., Nordsieck, K. H., Burgh, E. B., et al. 2003, in *SPIE Conf. Ser.*, eds. M. Iye, & A. F. M. Moorwood, *Proc. SPIE*, **4841**, 1634
- Landoni, M., Falomo, R., Treves, A., et al. 2013, *AJ*, **145**, 114
- Landoni, M., Massaro, F., Paggi, A., et al. 2015, *AJ*, **149**, 163
- Landoni, M., Paiano, S., Falomo, R., Scarpa, R., & Treves, A. 2018, *ApJ*, **861**, 130
- Landt, H., Padovani, P., & Giommi, P. 2002, *MNRAS*, **336**, 945
- Lefaucheur, J., & Pita, S. 2017, *A&A*, **602**, A86
- Lister, M. L., Aller, M., Aller, H., et al. 2011, *ApJ*, **742**, 27
- Mannucci, F., Basile, F., Poggianti, B. M., et al. 2001, *MNRAS*, **326**, 745
- Marchesi, S., Kaur, A., & Ajello, M. 2018, *AJ*, **156**, 212
- Marchesini, E. J., Masetti, N., Chavushyan, V., et al. 2016, *A&A*, **596**, A10
- Marchesini, E. J., Peña-Herazo, H. A., Álvarez Crespo, N., et al. 2019, *Ap&SS*, **364**, 5
- Markwardt, C. B. 2009, in *Non-linear Least-squares Fitting in IDL with MPFIT*, eds. D. A. Bohlender, D. Durand, & P. Dowler, *ASP Conf. Ser.*, **411**, 251
- Masetti, N., Sbarufatti, B., Parisi, P., et al. 2013, *A&A*, **559**, A58
- Massaro, F., Paggi, A., Errando, M., et al. 2013, *ApJS*, **207**, 16
- Massaro, F., D'Abrusco, R., Landoni, M., et al. 2015a, *ApJS*, **217**, 2
- Massaro, F., Landoni, M., D'Abrusco, R., et al. 2015b, *A&A*, **575**, A124
- Massaro, E., Maselli, A., Leto, C., et al. 2015c, *Ap&SS*, **357**, 75
- Meisner, A. M., & Romani, R. W. 2010, *ApJ*, **712**, 14
- Mirizzi, A., Raffelt, G. G., & Serpico, P. D. 2007, *Phys. Rev. D*, **76**, 023001
- Mirzoyan, R. 2019, *ATel*, **12538**, 1
- Mücke, A., & Protheroe, R. J. 2001, *Astropart. Phys.*, **15**, 121
- Nigro, C., Deil, C., Zanin, R., et al. 2019, *A&A*, **625**, A10
- Nilsson, K., Pursimo, T., Heidt, J., et al. 2003, *A&A*, **400**, 95
- Nilsson, K., Lindfors, E., Takalo, L. O., et al. 2018, *A&A*, **620**, A185
- Padovani, P., & Giommi, P. 1995, *ApJ*, **444**, 567
- Paggi, A., Milisavljevic, D., Masetti, N., et al. 2014, *AJ*, **147**, 112
- Paiano, S., Falomo, R., Franceschini, A., Treves, A., & Scarpa, R. 2017a, *ApJ*, **851**, 135
- Paiano, S., Landoni, M., Falomo, R., et al. 2017b, *ApJ*, **837**, 144
- Paiano, S., Falomo, R., Treves, A., & Scarpa, R. 2018, *ApJ*, **854**, L32
- Paiano, S., Falomo, R., Treves, A., & Scarpa, R. 2020, *MNRAS*, **497**, 94
- Paliya, V. S., Böttcher, M., Olmo-García, A., et al. 2020, *ApJ*, **902**, 29
- Peña-Herazo, H. A., Marchesini, E. J., Álvarez Crespo, N., et al. 2017, *Ap&SS*, **362**, 228
- Peña-Herazo, H. A., Amaya-Almazán, R. A., Massaro, F., et al. 2020, *A&A*, **643**, A103
- Perlman, E. S., Stocke, J. T., Schachter, J. F., et al. 1996, *ApJS*, **104**, 251
- Perlman, E. S., Padovani, P., Giommi, P., et al. 1998, *AJ*, **115**, 1253
- Piranomonte, S., Perri, M., Giommi, P., Landt, H., & Padovani, P. 2007, *A&A*, **470**, 787
- Pita, S., Goldoni, P., Boisson, C., et al. 2014, *A&A*, **565**, A12
- Poole, T. S., Breeveld, A. A., Page, M. J., et al. 2008, *MNRAS*, **383**, 627
- Raghunathan, S., Clowes, R. G., Campusano, L. E., et al. 2016, *MNRAS*, **463**, 2640
- Rau, A., Schady, P., Greiner, J., et al. 2012, *A&A*, **538**, A26
- Rector, T. A., Gabuzda, D. C., & Stocke, J. T. 2003, *AJ*, **125**, 1060
- Roming, P. W. A., Kennedy, T. E., Mason, K. O., et al. 2005, *Space Sci. Rev.*, **120**, 95
- Salamon, M. H., Stecker, F. W., & de Jager, O. C. 1994, *ApJ*, **423**, L1
- Sbarufatti, B., Treves, A., & Falomo, R. 2005, *ApJ*, **635**, 173
- Sbarufatti, B., Falomo, R., Treves, A., & Kotilainen, J. 2006, *A&A*, **457**, 35
- Scarpa, R., Urry, C. M., Falomo, R., et al. 1999, *ApJ*, **521**, 134
- Schlaflly, E. F., & Finkbeiner, D. P. 2011, *ApJ*, **737**, 103
- Schwoppe, A., Hasinger, G., Lehmann, I., et al. 2000, *Astron. Nachr.*, **321**, 1
- Sembach, K. R., & Savage, B. D. 1992, *ApJS*, **83**, 147
- Shaw, M. S., Romani, R. W., Cotter, G., et al. 2013, *ApJ*, **764**, 135
- Sheinis, A. I., Bolte, M., Epps, H. W., et al. 2002, *PASP*, **114**, 851
- Skrutskie, M. F., Cutri, R. M., Stiening, R., et al. 2006, *AJ*, **131**, 1163
- Smette, A., Sana, H., Noll, S., et al. 2015, *A&A*, **576**, A77
- Sol, H., Zech, A., Boisson, C., et al. 2013, *Astropart. Phys.*, **43**, 215
- Stickel, M., Padovani, P., Urry, C. M., Fried, J. W., & Kuehr, H. 1991, *ApJ*, **374**, 431
- Takeuchi, Y., Kataoka, J., Maeda, K., et al. 2013, *ApJS*, **208**, 25
- Tody, D. 1986, in *The IRAF Data Reduction and Analysis System*, ed. D. L. Crawford, *SPIE Conf. Ser.*, **627**, 733
- Urry, C. M., & Padovani, P. 1995, *PASP*, **107**, 803
- Urry, C. M., Scarpa, R., O'Dowd, M., et al. 2000, *ApJ*, **532**, 816
- Vermeulen, R. C., & Cohen, M. H. 1994, *ApJ*, **430**, 467
- Vladilo, G., Centurion, M., Falomo, R., & Molaro, P. 1997, *A&A*, **327**, 47
- Wagner, S. J., & Witzel, A. 1995, *ARA&A*, **33**, 163
- Zerbi, R. M., Chincarini, G., Ghisellini, G., et al. 2001, *Astron. Nachr.*, **322**, 275

## Appendix A: Redshift values in our simulation sample modified with respect to the 3FHL

### A.1. Notes on each source in Table A.1

*1RXSJ002200.9+000659*. This source has no redshift in the 3FHL catalogue, but its spectrum from the SDSS database (Fig. A.1, first row, left) clearly shows the features of an elliptical galaxy at  $z = 0.3057$ . This is the value that we adopted in our simulations.

*KUV00311–1938*. This source has redshift  $z = 0.61$  (Piranomonte et al. 2007) in the 3FHL catalogue. Several other high S/N spectra (see e.g. Pita et al. 2014) did not show any features except for an MgII absorption system at  $z = 0.506$ . We thus adopted  $z = 0.506$  for this source in our simulations.

*PKS 0118–272*. The redshift of PKS 0118–272 is unknown in the 3FHL catalogue, but several authors (see e.g. Vladilo et al. 1997) have detected a rich absorption system at  $z = 0.558$ . This is the value that we adopted in our simulations for the redshift of the source.

*MG1 J021114+1051*. This source has  $z = 0.2$  (Meisner & Romani 2010) in the 3FHL catalogue. This value is obtained from statistical estimations using imaging observations that we do not consider accurate enough for our purposes. We therefore consider its redshift unknown.

*RBS 334*. This source has no redshift in the 3FHL catalogue. Its redshift has been measured to be  $z = 0.411$  in Pita et al. (2014) and we used this value in our simulations.

*NVSS J033859–284619*. The redshift of NVSS J033859–284619 is unknown in the 3FHL catalogue. Halpern et al. (1997) determined that its redshift is  $z = 0.251$ . This is the value that we used in our simulations

*PKS 0447–439*. This source has  $z = 0.205$  (Perlman et al. 1998) in the 3FHL catalogue. This value was obtained from the detection of weak features, but it has never been independently confirmed with optical spectroscopy. Furthermore, higher S/N spectra (see e.g. Pita et al. 2014) obtained later failed to confirm it, we thus consider the redshift of PKS 0447–439 to be unknown.

*1ES 0502+675*. This source has  $z = 0.416$  (Landt et al. 2002) in the 3FHL catalogue. A value of  $z = 0.314$  has also been reported but not discussed in Scarpa et al. (1999). This puzzling situation may be linked to a contaminating object, identified as a star (Giovannini et al. 2004), which lies at just 0.33 arcsec from the BL Lac object (Scarpa et al. 1999). This distance is smaller than atmospheric seeing, which implies that in every ground observation the spectrum of the BL Lac object is contaminated by the features of the star, which should be subtracted before attempting to measure the BL Lac redshift. We conclude therefore that there is no convincing, precise measurement of the redshift of 1ES 0502+675, and we consider it unknown.

*TXS 0518+211*. This source has  $z = 0.108$  (Shaw et al. 2013) in the 3FHL catalogue obtained from weak emission features. This value has not been confirmed independently with spectroscopy. Furthermore, higher S/N spectra (see e.g. Paiano et al. 2017a) obtained later failed to confirm this result. We thus conclude that the redshift of TXS 0518+211 is unknown.

*PKS 0548–322*. The redshift of PKS 0548–322 is unknown in the 3FHL catalogue, but its spectrum corresponds to that of an

elliptical galaxy at  $z = 0.069$  (see e.g. Fosbury & Disney 1976) and we took this value as its redshift.

*B3 0609+413*. This source has no redshift in the 3FHL catalogue, but an MgII system at  $z = 1.108$  has been detected in his spectrum (Shaw et al. 2013). We adopted this value as its redshift.

*PMN J0622–2605*. The redshift of PMN J0622–2605 in the 3FHL catalogue is  $z = 0.414$  obtained from a spectrum in the 6dF database (Jones et al. 2004, 2009). However an examination of the spectrum (see Fig. A.1, first row, right) reveals no convincing feature in the 6dF spectrum, and we consider its redshift unknown.

*1ES 0647+250*. This source has redshift  $z = 0.203$  in the 3FHL catalogue. This value (Rector et al. 2003) has been classified as tentative. Recent high S/N spectra (Paiano et al. 2017a) failed to show any features, and therefore we consider its redshift unknown.

*PMN J0816–1311*. PMN J0816–1311 has redshift  $z = 0.046$  in the 3FHL catalogue obtained from a 6dF spectrum (Jones et al. 2004, 2009). However, an MgII system at  $z = 0.288$  has been detected in its spectrum with X-shooter (Pita et al. 2014). Its redshift is thus higher than 0.288, and we adopted  $z = 0.3$  as its redshift in our simulations.

*NVSS J110735+022225*. The redshift of NVSS J110735+022225 is unknown in the 3FHL catalogue. An MgII system at  $z = 1.0743$  has been detected in its SDSS spectrum (Fig. A.1, second row, left), and we took this value as its redshift.

*RBS 0970*. This source has redshift  $z = 0.124$  in the 3FHL catalogue. This value (Perlman et al. 1996) has not been confirmed in several subsequent observations (see e.g. Paiano et al. 2017b). We therefore consider its redshift unknown.

*S4 1250+53*. The redshift of S4 1250+53 is unknown in the 3FHL catalogue. An MgII system at  $z = 0.664$  has been detected in its SDSS spectrum (Fig. A.1, second row, right) and we adopted this value as its redshift.

*NVSS J141028+143841*. This source has no redshift in the 3FHL catalogue, but its SDSS spectrum shows the typical features of an elliptical galaxy at  $z = 0.144$  (Fig. A.1, third row, left), which we took as its redshift.

*PKS 1424+240*. The redshift of PKS 1424+240 is unknown in the 3FHL catalogue. A hydrogen Ly- $\alpha$  absorption system at  $z = 0.604$  has been detected in its spectrum (Furniss et al. 2013), and we adopted this value as its redshift in our simulations.

*RBS 1409*. This source has redshift  $z = 0.15$  (Schwope et al. 2000) in the 3FHL catalogue, which has not been confirmed in subsequent observations (Shaw et al. 2013). We therefore consider its redshift unknown.

*SUMSS J144236–462302*. This source has no redshift in the 3FHL catalogue, but its SDSS spectrum is that of an elliptical galaxy at  $z = 0.1026$  (Fig. A.1, third row, right), which we took as its redshift.

*PKS 1440–389*. The redshift of PKS 1440–389 in the 3FHL catalogue is  $z = 0.069$ , obtained from a low S/N spectrum in the 6dF database (Jones et al. 2004, 2009). Subsequent high S/N spectra failed to show any features, and we considered its redshift as unknown in our simulations. In this paper we report that its redshift is  $z = 0.1385$  (see Sect. 6.4 for details).

**Table A.1.** Sources in our sample for which we modified the redshift value with respect to the 3FHL catalogue.

3FHL		Simulations			Measured		
3FHL name	Source name	redshift	redshift reference	redshift lower limit	adopted redshift	adopted redshift/ lower limit reference	redshift
3FHL J0022.0+0006	1RXS J002200.9+000659	–	–	–	0.3057	(1)	–
3FHL J0033.5–1921	KUV 00311–1938	0.61	(2)	0.506	0.506	(3)	–
3FHL J0120.4–2701	PKS 0118–272	–	–	0.558	0.558	(5)	–
3FHL J0211.2+1051	MG1 J021114+1051	0.2	(6)	–	0.3	–	–
3FHL J0237.6–3602	RBS 334	–	–	–	0.411	(3)	–
3FHL J0338.9–2848	NVSS J033859–284619	–	–	–	0.251	(4)	–
3FHL J0449.4–4350	PKS 0447–439	0.205	(7)	–	0.3	(3),(8)	–
3FHL J0508.0+6737	1ES 0502+675	0.416	(9)	–	0.3	(10),(11)	–
3FHL J0521.7+2112	TXS 0518+211	0.108	(12)	–	0.3	(8)	–
3FHL J0550.5–3215	PKS 0548–322	–	–	–	0.069	(13)	–
3FHL J0612.8+4122	B3 0609+413	–	–	1.108	1.108	(12)	–
3FHL J0622.4–2606	PMN J0622–2605	0.414	(14)	–	0.3	–	–
3FHL J0650.7+2503	1ES 0647+250	0.203	(15)	–	0.3	(8)	–
3FHL J0816.4–1311	PMN J0816–1311	0.046	(14)	0.288	0.3	(3)	–
3FHL J1107.4+0221	NVSS J110735+022225	–	–	1.0743	1.0743	(1)	–
3FHL J1120.8+4212	RBS 0970	0.124	(16)	–	0.3	(8)	–
3FHL J1253.1+5300	S4 1250+53	–	–	0.664	0.664	(1)	–
3FHL J1410.5+1438	NVSS J141028+143841	–	–	–	0.144	(1)	–
3FHL J1427.0+2348	PKS 1424+240	–	–	0.604	0.604	(17)	–
3FHL J1436.9+5639	RBS 1409	0.15	(18)	–	0.3	(12)	–
3FHL J1442.5–4621	SUMSS J144236–462302	–	–	–	0.1026	(14)	–
3FHL J1443.9–3908	PKS 1440–389	0.069	(14)	–	0.3	(12)	0.1385
3FHL J1447.9+3608	RBS 1432	–	–	0.739	0.739	(1)	–
3FHL J1603.8–4903	PMN J1603–4904	–	–	–	0.232	(19)	–
3FHL J1610.6–6649	PMN J1610–6649	–	–	0.447	0.447	(1)	–
3FHL J1903.2+5540	TXS 1902+556	–	–	0.727	0.727	(12)	–
3FHL J1918.2–4111	PMN J1918–4111	–	–	1.591	1.591	(12).	–
3FHL J1931.1+0937	RX J1931.1+0937	–	–	0.476	0.476	(12)	–
3FHL J1958.3–3011	1RXS J195815.6–301119	0.119	(14)	–	0.3	–	0.119
3FHL J2243.9+2020	RGB J2243+203	–	–	0.395	0.395	(12)	–
3FHL J2255.2+2410	MITG J225517+2409	–	–	0.864	0.864	(1)	–
3FHL J2324.7–4040	1ES 2322–409	0.174	(14)	–	0.3	–	0.1738

**Notes.** The columns contain the source name, the redshift value in the 3FHL catalogue, the reference for that redshift value, the lower limit value, the adopted redshift value, the reference justifying the adopted value, and finally, for sources discussed in this paper, the measured redshift. If the entry (redshift, lower limit, reference) is unknown, the legend is ‘–’. There is a total of 32 sources with 7 new redshifts and 14 lower limits. The lower limit of PMN J0816–1311 ( $z \geq 0.288$ ) is smaller than the adopted value for unknown redshift  $z=0.3$ ; we therefore used  $z=0.3$  for this source in our simulations. The spectra of ten sources have not been published yet to our knowledge, and we present them for reference in Fig. A.1.

**References.** (1) Blanton et al. (2017); (2) Piranomonte et al. (2007); (3) Pita et al. (2014); (4) Halpern et al. (1997); (5) Vladilo et al. (1997); (6) Meisner & Romani (2010), imaging redshift; (7) Perlman et al. (1998), weak features, unconfirmed; (8) Paiano et al. (2017a,b); (9) Landt et al. (2002); (10) Scarpa et al. (1999); (11) Giovannini et al. (2004); (12) Shaw et al. (2013); (13) Fosbury & Disney (1976); (14) Jones et al. (2004, 2009), low S/N spectrum; (15) Rector et al. (2003), tentative value; (16) Perlman et al. (1996); (17) Furniss et al. (2013); (18) Schwoppe et al. (2000); (19) Goldoni et al. (2016).

**RBS 1432.** The redshift of RBS 1432 is unknown in the 3FHL catalogue. An MgII system at  $z = 0.739$  has been detected in its SDSS spectrum (Fig. A.1, fourth row, left), and we took this value as its redshift.

**PMN J1603–4904.** This source has no redshift in 3FHL. Its redshift has been measured as  $z = 0.232$  (Goldoni et al. 2016), and we used this value in our simulations.

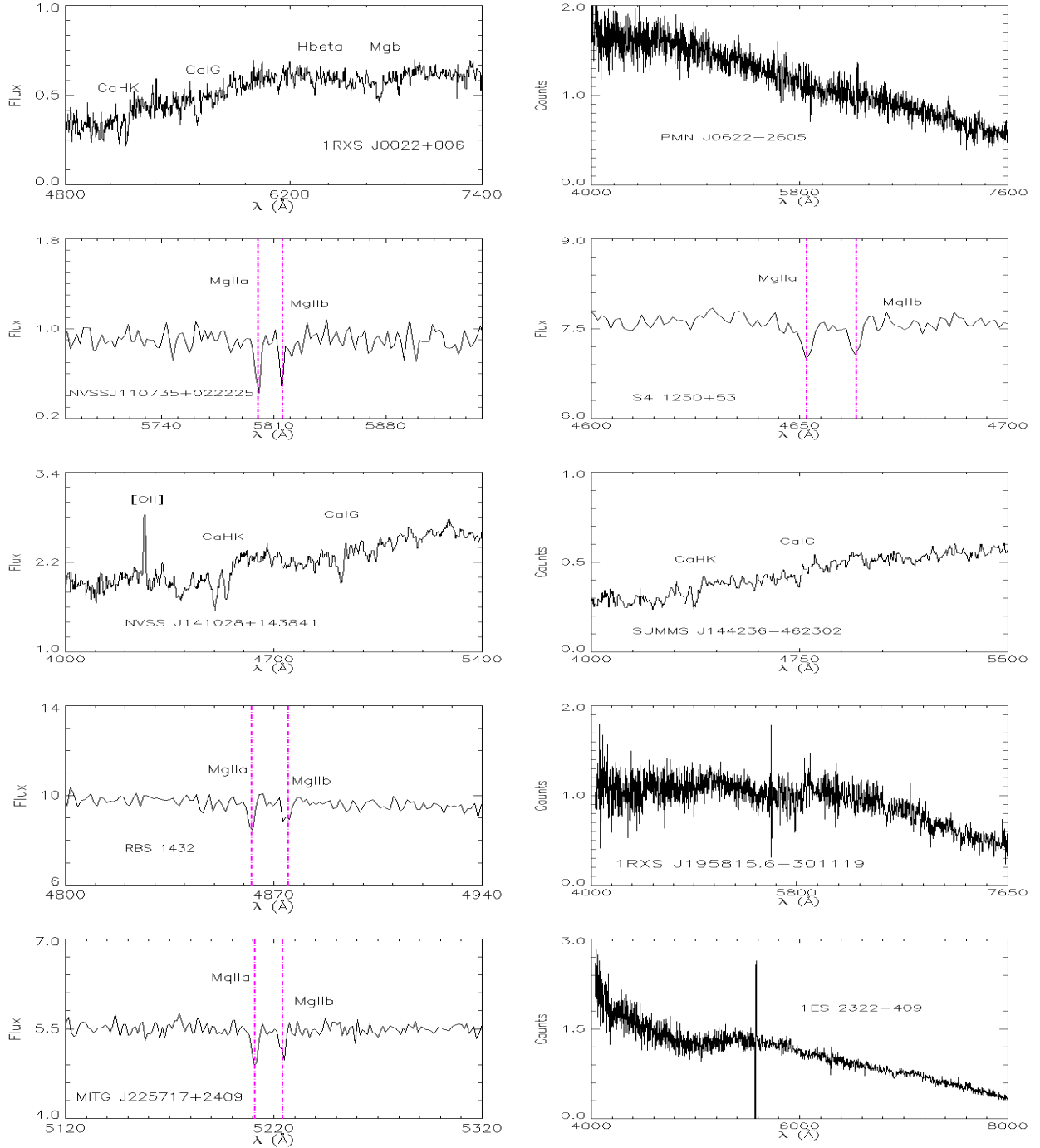
**PMN J1610–6649.** The redshift of PMN J1610–6649 is unknown in the 3FHL catalogue. An MgII system at  $z = 0.447$  has been detected in its spectrum (Shaw et al. 2013), and we took this value as its redshift.

**TXS 1902+556.** Same as the source above, the MgII system is at  $z = 0.727$ .

**PMN J1918–4111.** Same as the source above, the MgII system is at  $z = 1.591$ .

**RX J1931.1+0937.** Same as the source above, the MgII system is at  $z = 0.476$ .

**1RXS J195815.6–301119.** 1RXS J195815.6–301119 has a redshift of  $z = 0.119$  in the 3FHL catalogue from a low S/N spectrum in the 6dF database (Fig. A.1, fourth row, right). Its redshift value was considered unknown in our simulations, but we confirmed it at  $z = 0.119$  in this paper (see Sect. 6.15 for details).



**Fig. A.1.** Public SDSS and 6dF spectra of sources whose redshift values in our sample are different to those in the 3FHL catalogue. When needed, the lines used to measure the redshift are marked in the plots. *First row, left:* 1RXS J0022+0006,  $z = 0.3057$ ; *right:* PMN J0622-2605, redshift unknown. *Second row, left:* NVSS J110735+022225,  $z \geq 1.0743$ ; *right:* S4 1250+53,  $z \geq 0.664$ . *Third row, left:* NVSS J141028+143841,  $z = 0.144$ ; *right:* SUMSS J144236-462302,  $z = 0.1026$ . *Fourth row, left:* RBS 1432,  $z \geq 0.739$ ; *right:* 1RXS J195815.6-301119, redshift unknown (but  $z = 0.119$  confirmed here; see Sect. 6.15). *Fifth row, left:* MITG J225717+2409,  $z \geq 0.864$ ; *right:* 1ES 2322-409, redshift unknown (but  $z = 0.1738$  measured here; see Sect. 6.19).

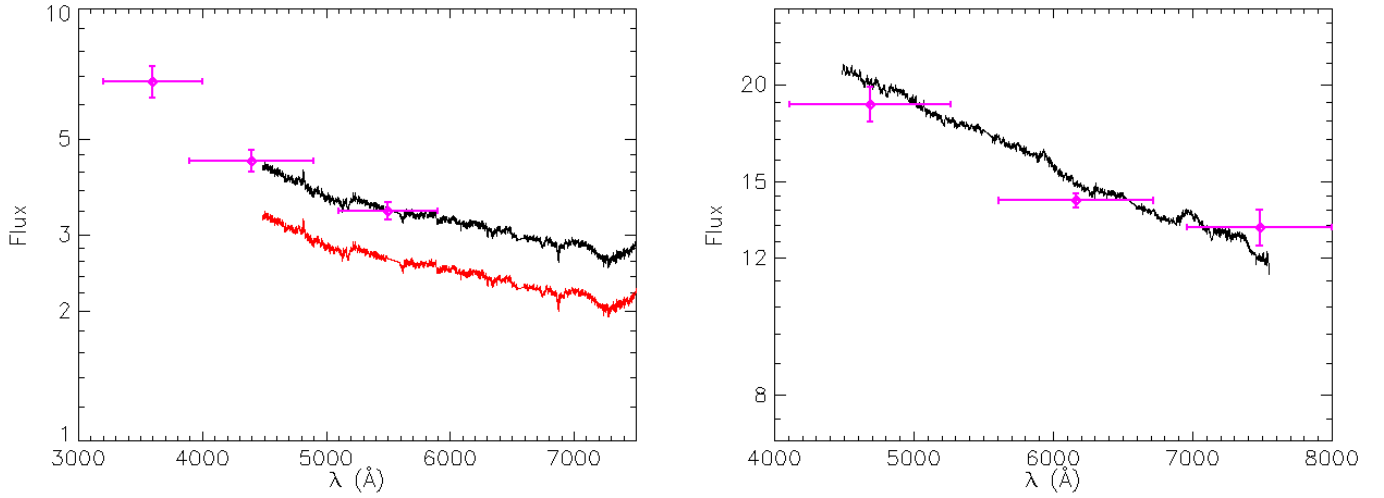
**RGB J2243+203.** This source has no redshift in the 3FHL catalogue. An MgII system at  $z = 0.395$  has been detected in its spectrum (Shaw et al. 2013), and we took this value as its redshift.

**MITG J225517+2409.** The redshift of MITG J225517+2409 is unknown in the 3FHL catalogue. An MgII system at  $z = 0.864$

has been detected in its SDSS spectrum (Fig. A.1, fifth row, left), and we took this value as its redshift.

**1ES 2322-409.** 1ES 2322-409 has redshift  $z = 0.174$  in the 3FHL catalogue from a low S/N spectrum in the 6dF database (Fig. A.1, fifth row, right). Its redshift value was considered unknown in our simulations, but we measured it at  $z = 0.1738$  in this paper (see Sect. 6.19 for details).

## Appendix B: *Swift*-UVOT and REM/ROSS photometry



**Fig. B.1.** Spectra and near-contemporaneous photometry of 1RXS J015658.6–530208 and 1RXS J020922.2–522920. *Left panel:* SALT/RSS flux calibrated and telluric corrected optical spectrum of 1RXS J015658.6–530208 (in red), and the same spectrum scaled to the average UVOT photometry (in black). The  $u$ ,  $b$ , and  $v$  photometry and errors are shown in magenta. *Right panel:* SALT/RSS flux calibrated and telluric corrected optical spectrum of 1RXS J020922.2–522920. The  $g$ ,  $r$ , and  $i$  average photometry is shown in magenta. The photometric points are consistent with the spectral flux within  $2\sigma$  at most.

**Table B.1.** Log and results of *Swift*/UVOT observations of 1RXS J015658.6–530208 in  $v$ ,  $b$ , and  $u$  bands.

Date (UT)	MJD	$u$ (mag)	$b$ (mag)	$v$ (mag)
2019-11-30	58817	$17.12 \pm 0.05$	$18.11 \pm 0.06$	$17.72 \pm 0.09$

**Table B.2.** Log and fitting results of REM observations of 1RXS J020922.2–522920 in  $g$ ,  $r$ , and  $i$  bands.

Date (UT)	MJD	$g$ (mag)	$r$ (mag)	$i$ (mag)
2019-12-25	58842	$15.839 \pm 0.126$	$15.655 \pm 0.067$	$15.334 \pm 0.113$
2019-12-31	58848	$16.139 \pm 0.025$	$15.780 \pm 0.022$	$15.425 \pm 0.084$
2020-01-01	58849	$16.123 \pm 0.150$	$15.858 \pm 0.013$	$15.522 \pm 0.141$
2020-01-02	58850	$16.204 \pm 0.025$	$15.865 \pm 0.029$	$15.507 \pm 0.053$
2020-01-03	58851	$16.218 \pm 0.026$	$15.842 \pm 0.028$	$15.491 \pm 0.060$