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Studying star-forming processes at core and clump scales: the case of the young stellar object G29.862−0.0044 *(Corrigendum)*

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In the original article, a molecular emission line was not correctly identified. One of the used pieces of millimeter data obtained from the Atacama Large Millimeter Array database was incorrectly identified as the CH_3OCHO 20(2,19)–19(2,18)E line. The analyzed emission corresponds to the cyano radical (CN) $N = 2$–1, $J = 5/2$–3/2, $F = 7/2$–5/2 line at 226874.764 MHz, which can be blended with the $F = 5/2$–3/2 and $F = 3/2$–1/2 lines at 226874.183 and 226875.896 MHz, respectively. Hence, in Figs. 4, 5 and 7 of the original paper, the CN integrated emission is presented.

This misunderstanding does not alter any conclusion made in the original paper. The discussion regarding the morphology of the integrated CH_3OCHO emission and its relations with the infrared data still applies to the CN emission in the same way. As some works show (Han et al. 2015; Qiu et al. 2012; Zapata et al. 2008; Beuther et al. 2004), extended CN emission is usually detected towards molecular cores and star-forming regions, and this type of emission usually extends in filaments and concentrates in knots as is the case in G29.862−0.0044. In a following, paper we will study the CN emission towards a large sample of molecular cores, including this particular source (Paron et al., in prep.).

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